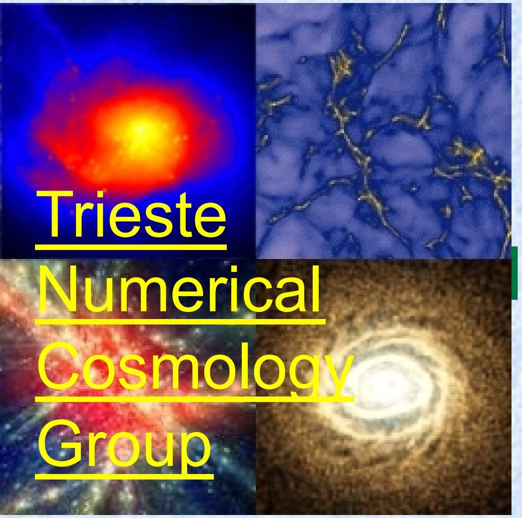
Gas Outflows in Cosmological Hydrodynamical Simulations



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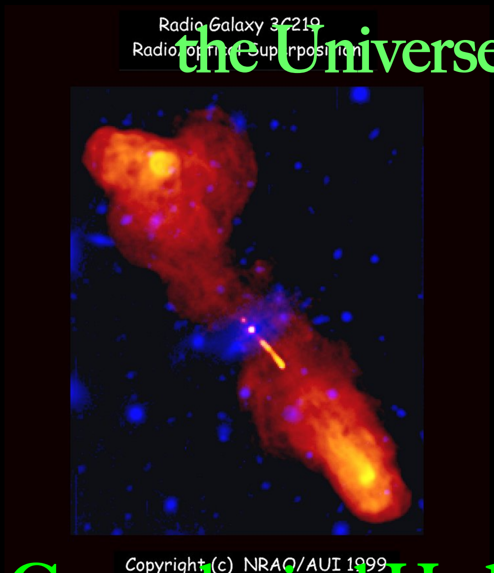
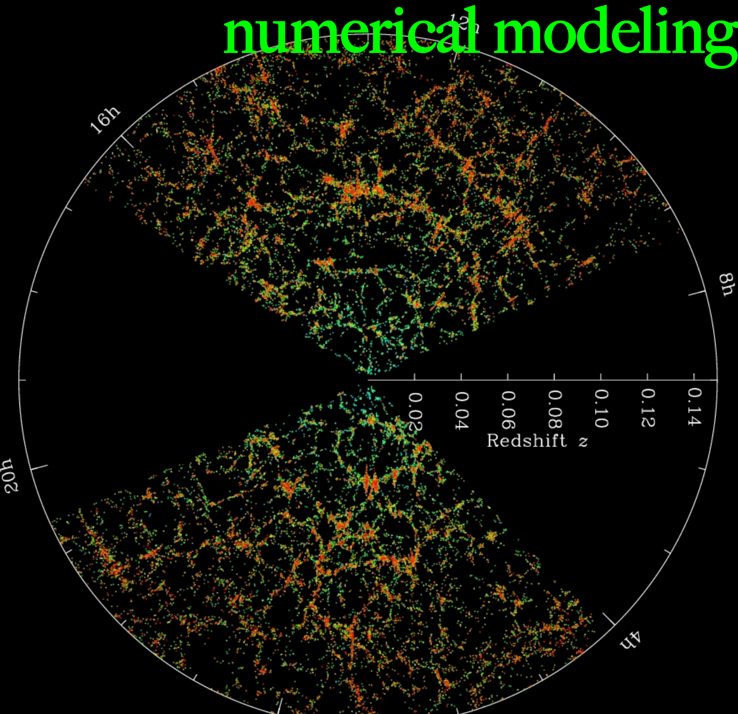
ASTRO @ Trieste

4 June 2015

Matter observed to be distributed in various forms at different scales in the Universe: Galaxy, Cluster, Large-scale structures



Cosmological Hydrodynamic Simulations is a tool for theoretical/ numerical modeling of galaxy formation & evolution



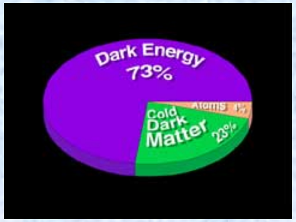
Galaxy Cluster Abell 2218 **HST** • WFPC2
NASA, A. Fruchter and the ERO Team (STScI) • STScI-PRC00-08

The Universe in a Box:
Cosmological Hydrodynamic Simulations

- Computational box ⇔ representative volume of the Universe
- Resolution elements (particles or grid) in box ⇔ matter

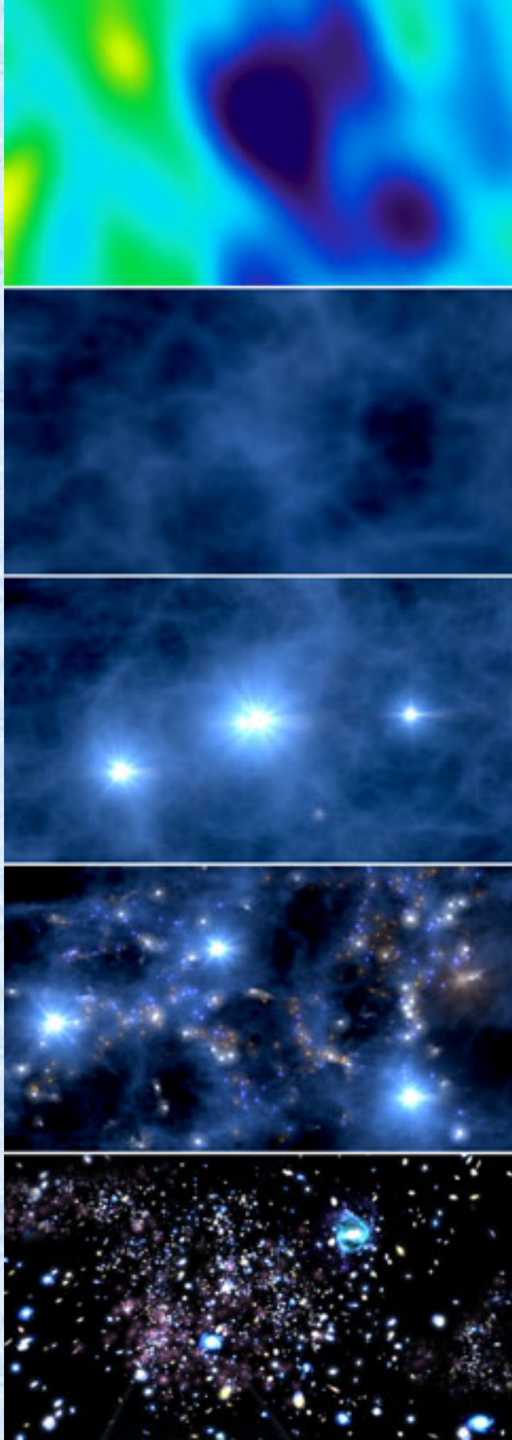
- Model LSS in terms of massive elements each of mass 10^6 - 10^7 $M_\odot$

ΛCDM cosmology

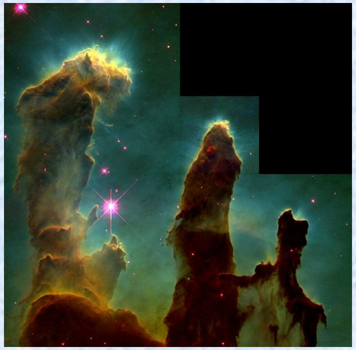


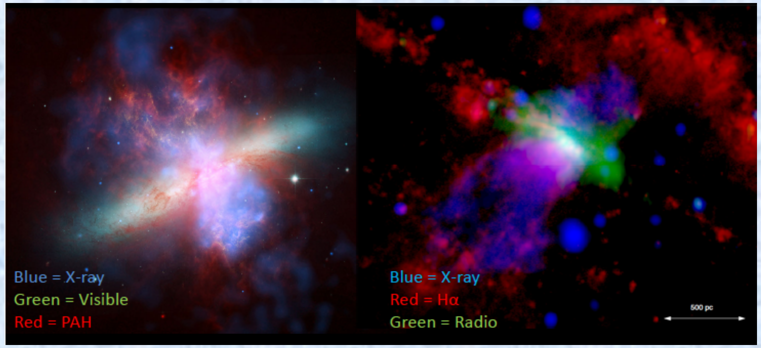
- Steps:
  - Generate the initial condition
    - Primordial density fluctuations (Gaussian) at CMB epoch (z~1100)
  - Follow the non-linear evolution of density fields numerically
  - Identify galaxies, clusters (group finder) at different z

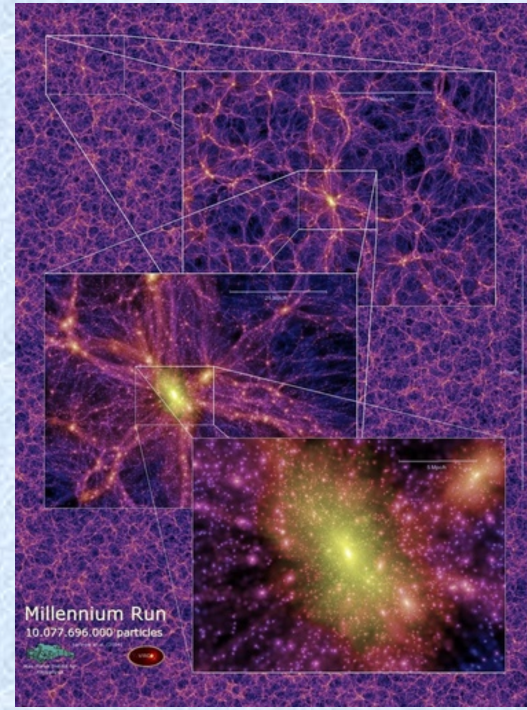
- Run on supercomputers
  - High speed & processing power, days to weeks of time

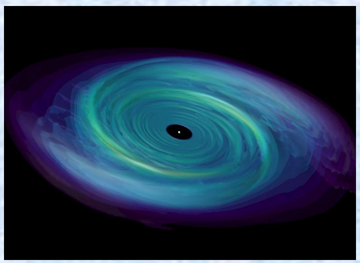


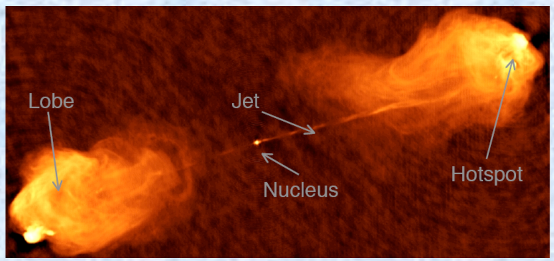
Why Sub-Resolution Models ?











Physics of baryons

- Radiative cooling and (photo + collisional) ionization heating of gas
- Fragmentation, clumping, multiphase ISM
- Star formation
- Metal production & chemical enrichment
- SN feedback, galactic wind
- AGN accretion + feedback
- …

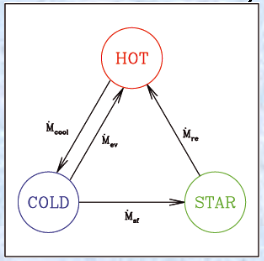
In cosmological hydrodynamical simulations

(few - 10's Mpc) box : Resolution ~ $10^6 M_{Sun}$, 1 kpc

- Baryonic physics, occurring on much smaller scales, are crucial ingredient
- Implemented as sub-resolution models

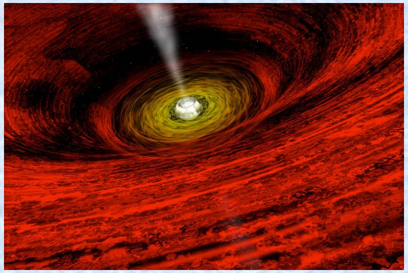
Modified-GADGET3 code: Sub-Resolution Physics

- GADGET : TreePM (gravity) - SPH (hydro)
  - Springel 2005, MNRAS, 364, 1105
- Metal-line cooling & radiative heating (Wiersma et al. 2009, MNRAS, 399, 574) in the presence of UV photoionizing background (Haardt & Madau 2001)



- Star Formation
- Stellar & Chemical Evolution (Tornatore et al. 2007, MNRAS, 382, 1050)
  - Metal (C, Ca, O, N, Ne, Mg, S, Si, Fe) release from SN type-II, type-Ia, & AGB stars; stellar age, mass & yield; different IMF; mass & metal loss from starburst

- SN Feedback
  - Thermal feedback (↑ T) : inefficient, energy radiated away quickly
  - ∴ Kinetic feedback (↑ v)
- AGN accretion + feedback



Star-Formation in Multiphase ISM

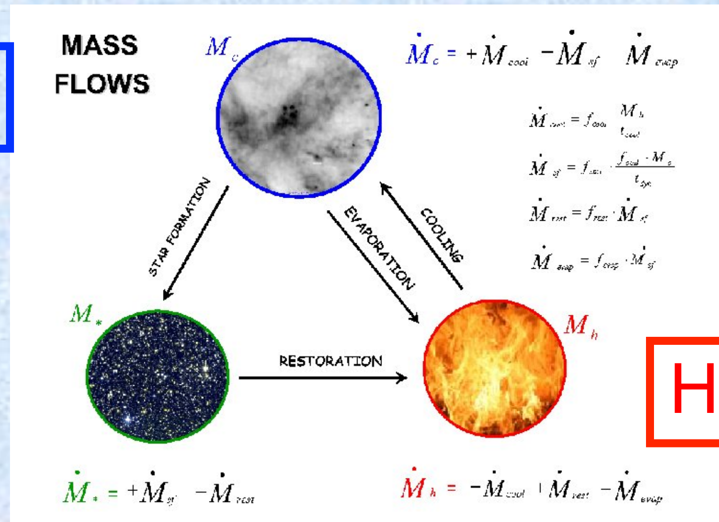
- High-density SPH particle represents a part of ISM
  - Composed of 2 gas phases & stars

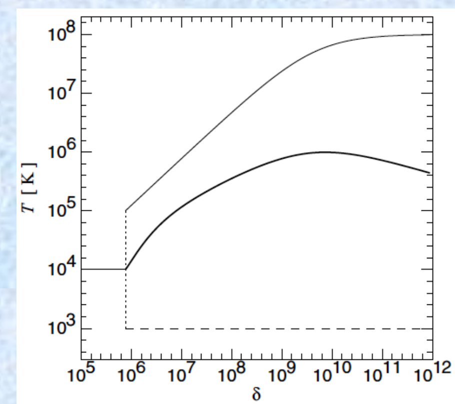
$\rho > \rho_{SF,th}$

COLD gas

STARS

HOT gas





- Effective model *(Springel & Hernquist 2003)*
  - Equilibrium solution
  - Self-regulated SF: constant effective pressure
- MUPPI = MUlti-Phase Particle Integrator *(Murante et al. 2010)*
  - Molecular fraction of gas ∝ Pressure
  - Mass & energy flows between components explicitly followed by numerically integrating system of ODEs within SPH timestep

Existing Models of SN Feedback

- Kinetic feedback : give velocity kick to gas
  - Energy-driven wind
    - Springel & Hernquist (2003)

$$v_w, \eta = \text{constant}$$

$$v_w = 3\sigma\sqrt{\frac{L}{L_{crit}} - 1}$$

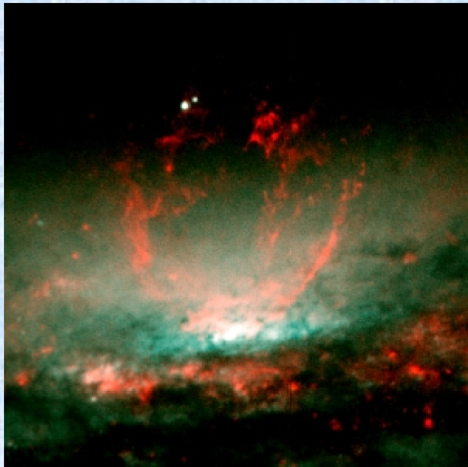
σ_0

Most of the models assume that wind velocity and mass-loading scales with some global galaxy property (mass, velocity dispersion, SFR)

  - Radially-varying wind velocity
    - Barai et al. (2013)
  - Combinations & variations of energy and momentum-driven
    - Schaye et al. (2010)
    - Dave et al. (2013)
    - Volgelsberger et al. (2014)

- Thermal feedback : increase gas temperature
  - Dalla Vecchia & Schaye (2012), Schaye et al. (2014)

- Turn off radiative cooling
  - Stinson et al. (2006)
- …



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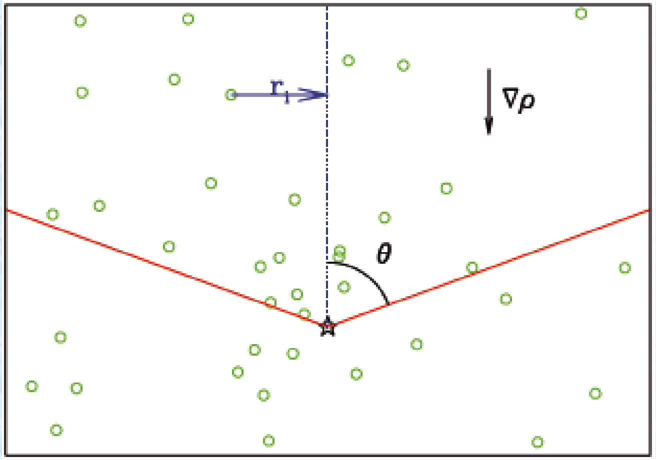
SN Energy Feedback in MUPPI *(Murante et al. 2015)*

- Energy imparted to gas particles
  - Inside SPH smoothing length and cone with semi-aperture angle = 60°
  - Along path of least resistance
    - Negative density gradient



- Direct distribution of
  - Thermal energy
    - Efficiency fraction
    - Injected to local hot phase
  - Kinetic energy
    - Efficiency fraction, Probability

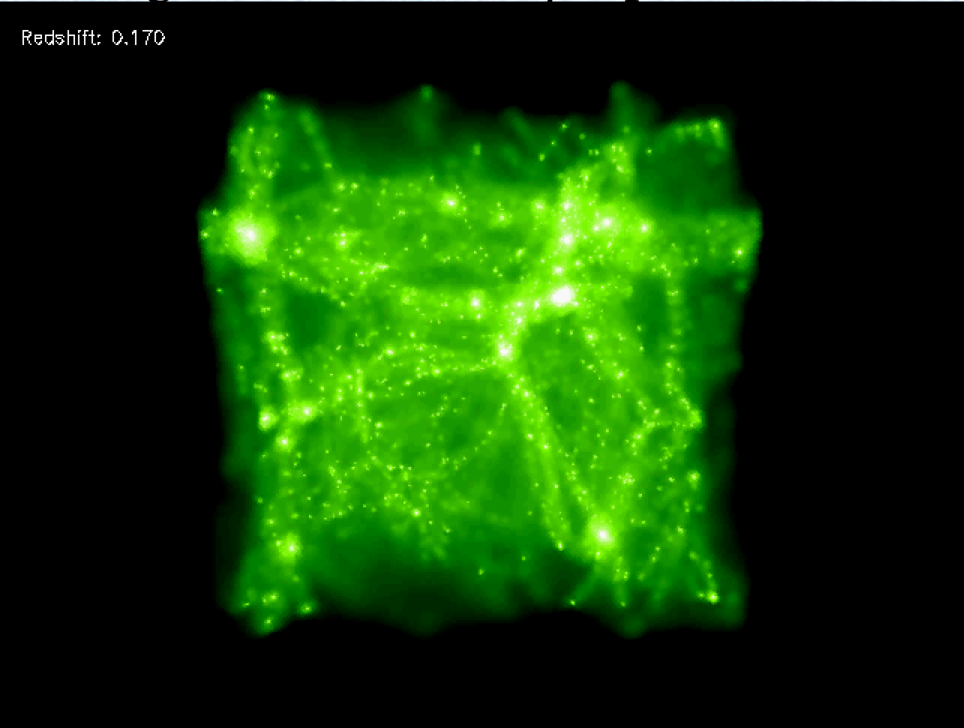
$$E_{th} = E_{SN} f_{fb,th} \frac{\Delta M_*}{M_{*,SN}}$$

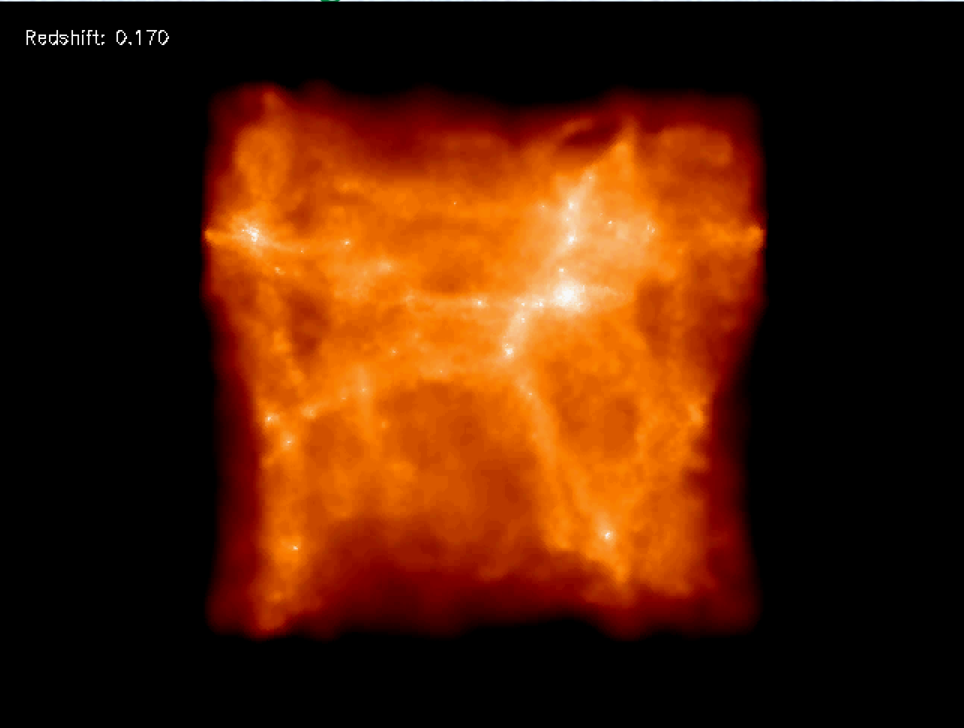
$$E_{kin} = E_{SN} f_{fb,kin}$$

- No direct input expression of wind velocity & outflow mass loading

- Wind particles are hydrodynamically decoupled

Large-scale filaments. (7 Mpc)^3 box at low-z. Dark matter - green, Gas - red, Stars - blue

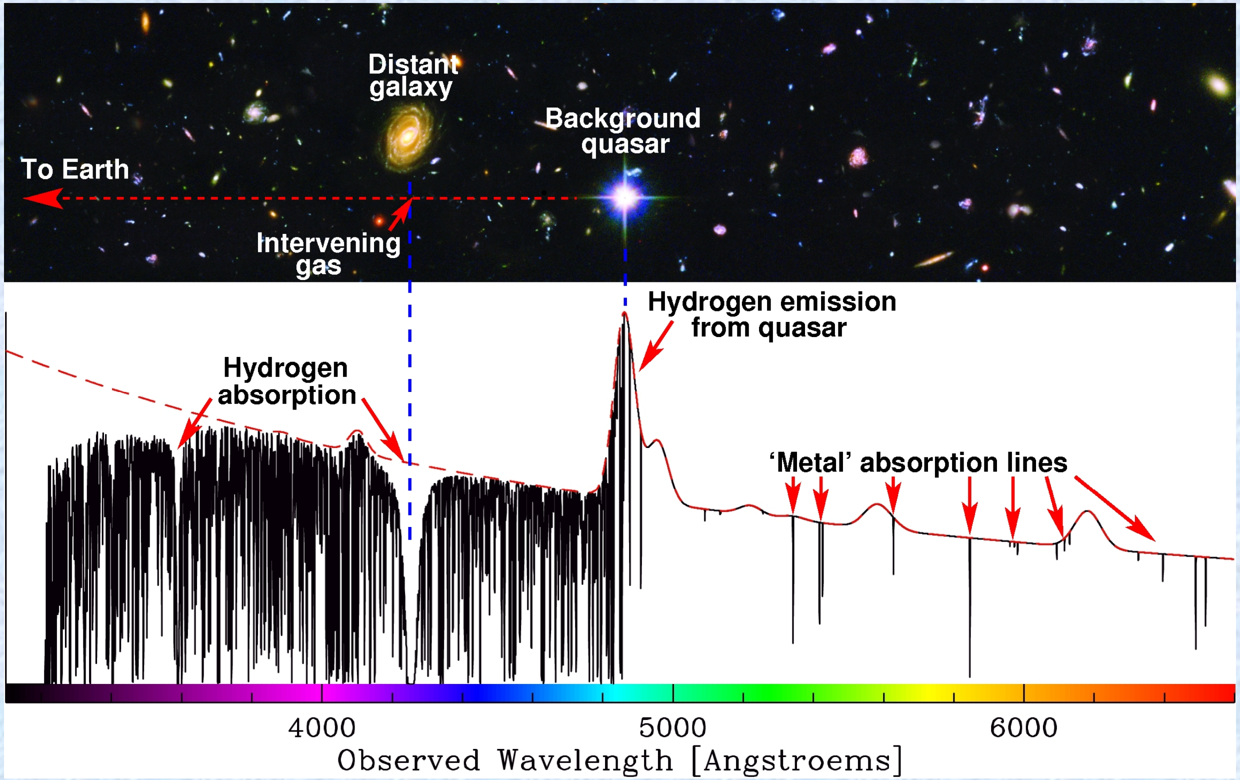




Collaboration with Observers for IGM science

InterGalactic Medium (IGM)

- Diffuse gas between galaxies
- Observable by spectroscopic absorption lines against background quasars



Stefano Cristiani,
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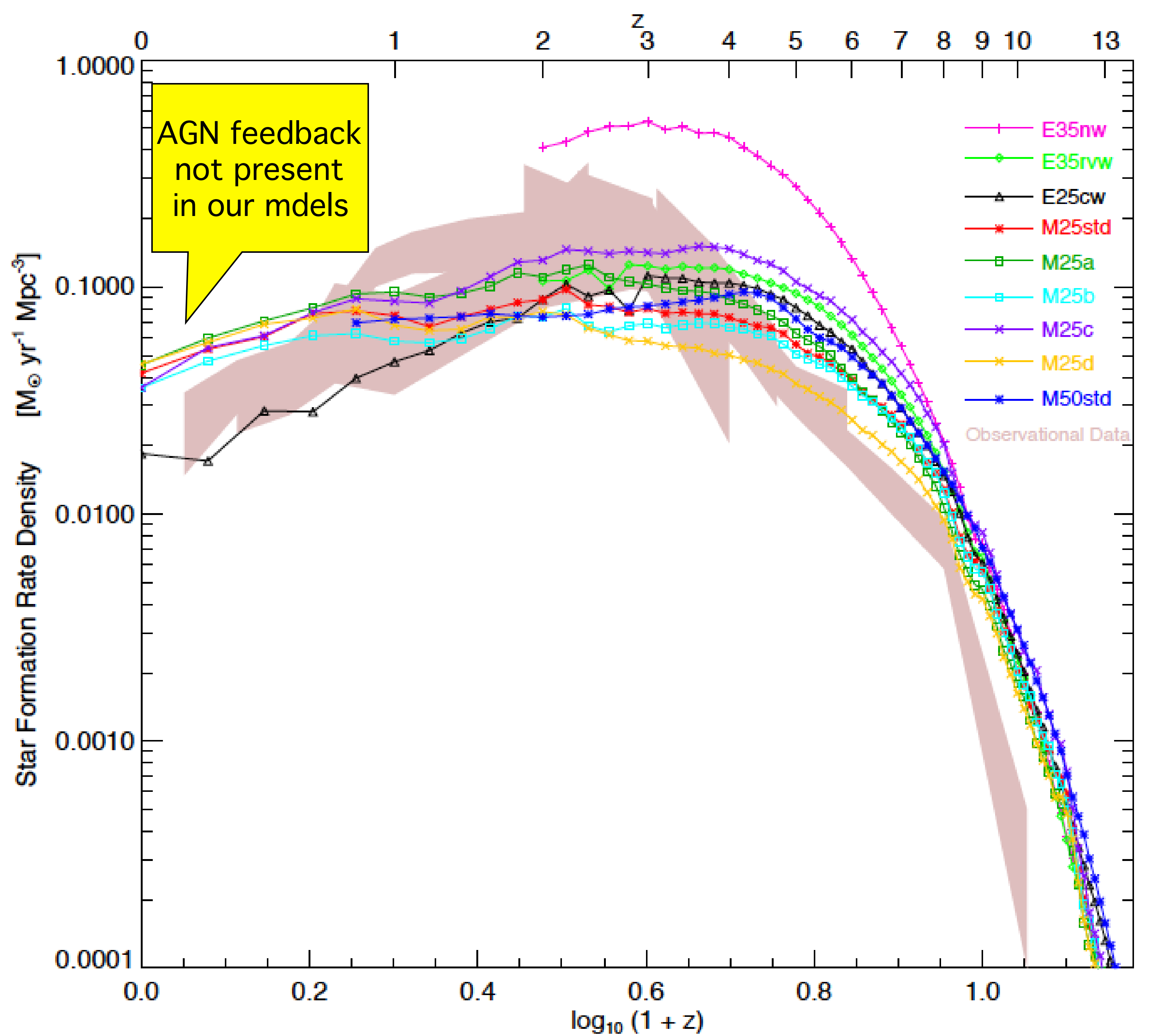
Matteo Viel,

Chiara Mongardi

Next talk

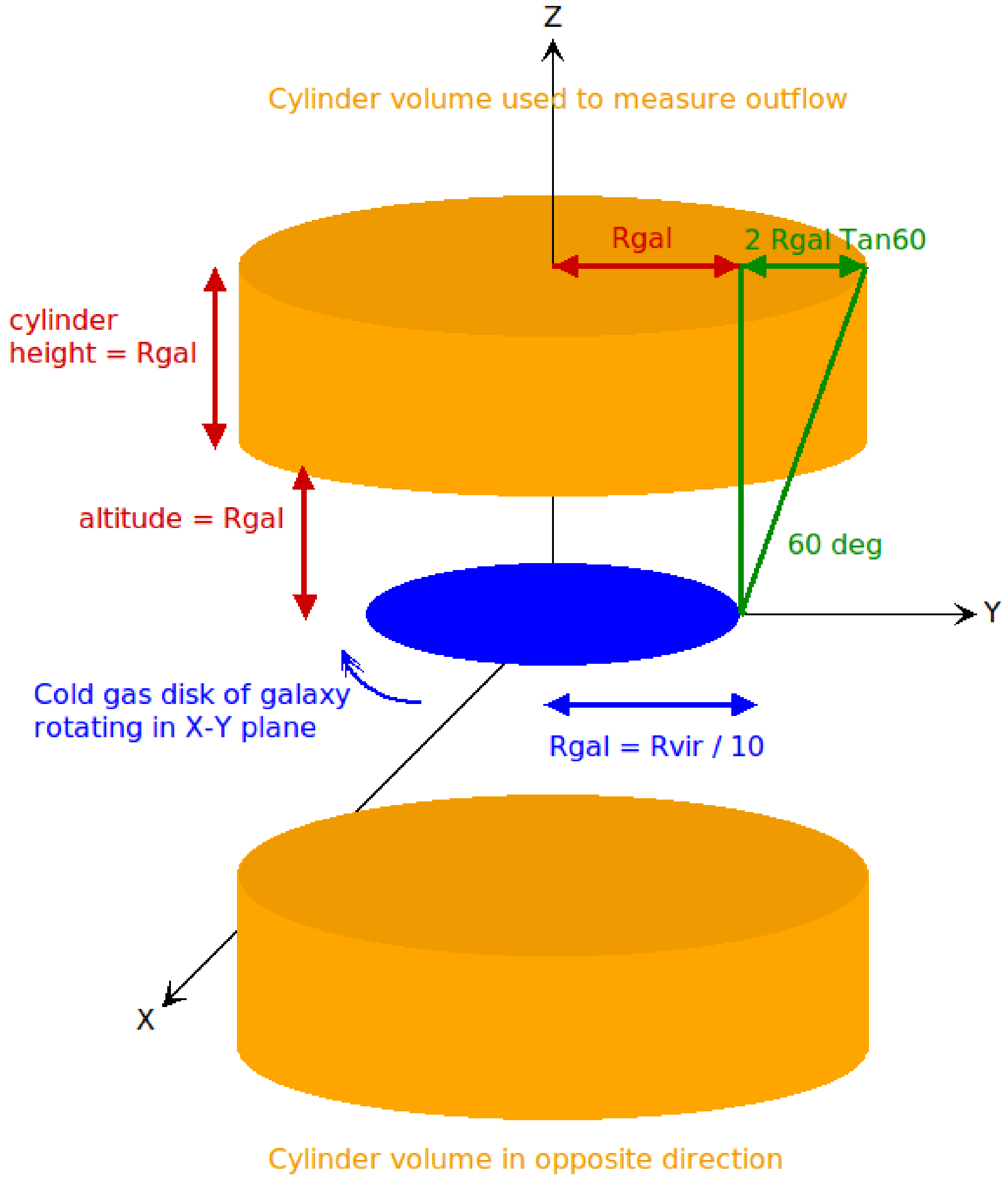
Simulation Runs (Barai et al. 2015)

Run Name	L_{box} [Mpc]	N_{part}	m_{gas} [$M_{\odot}$]	$m_{\star}$ [$M_{\odot}$]	L_{soft} [kpc]	SF & SN feedback sub-resolution physics Model	v_w	$f_{fb,out}$	$f_{fb,kin}$	P_{kin}
E35nw	35.56	2×320^3	8.72×10^6	2.18×10^6	2.77 (comoving)	Effective	0			
E35rvw	35.56	2×320^3	8.72×10^6	2.18×10^6	2.77 (comoving)	Effective	$v_w(r)$			
E25cw	25	2×256^3	5.36×10^6	1.34×10^6	0.69 (physical)	Effective	350			
M25std	25	2×256^3	5.36×10^6	1.34×10^6	0.69 (physical)	MUPPI		0.2	0.6	0.03
M25a	25	2×256^3	5.36×10^6	1.34×10^6	0.69 (physical)	MUPPI		0.4	0.4	0.03
M25b	25	2×256^3	5.36×10^6	1.34×10^6	0.69 (physical)	MUPPI		0.2	0.8	0.03
M25c	25	2×256^3	5.36×10^6	1.34×10^6	0.69 (physical)	MUPPI		0.2	0.6	0.01
M25d	25	2×256^3	5.36×10^6	1.34×10^6	0.69 (physical)	MUPPI		0.2	0.6	0.06
M50std	50	2×512^3	5.36×10^6	1.34×10^6	0.69 (physical)	MUPPI		0.2	0.5	0.03



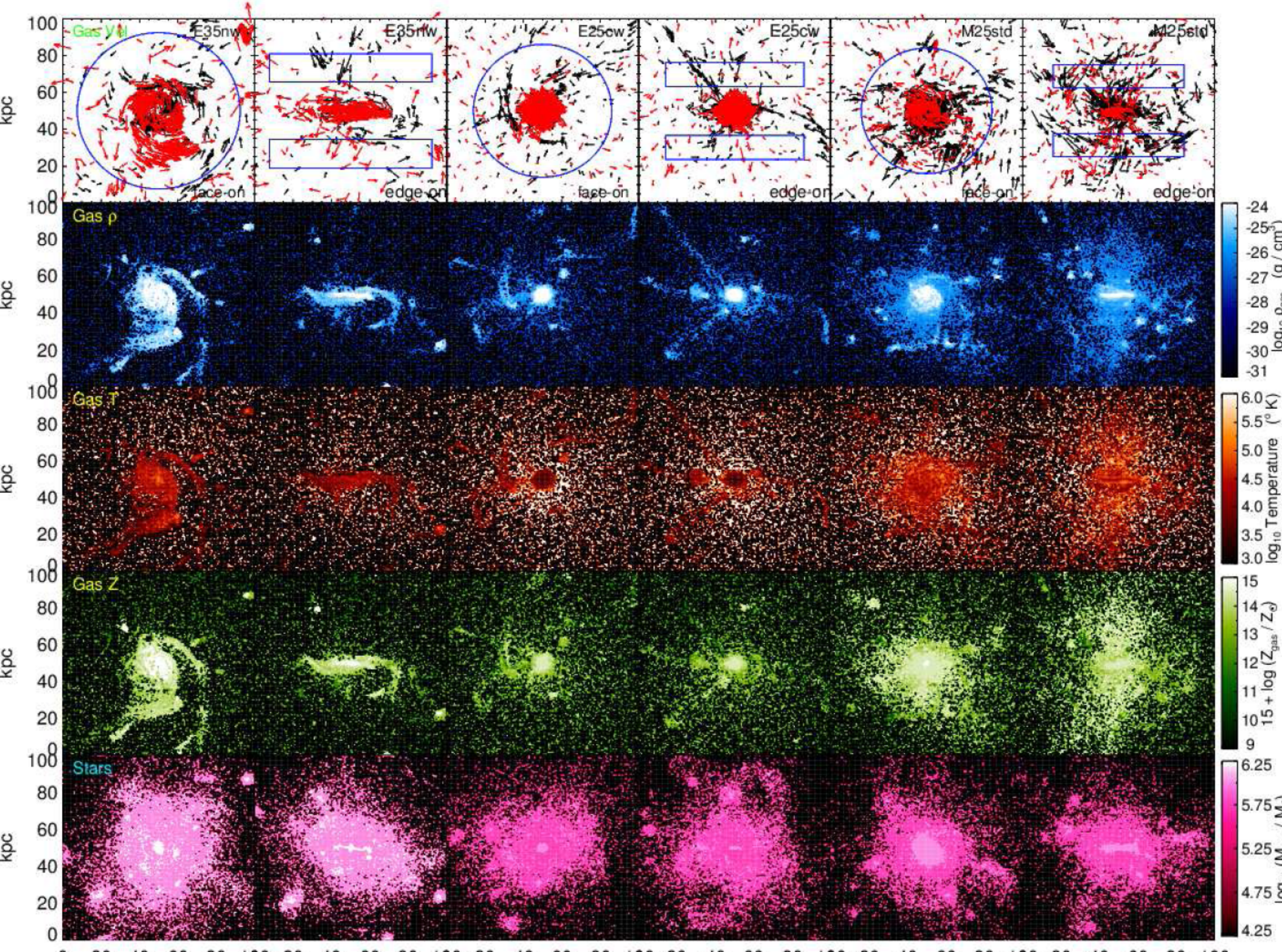
Star Formation Rate Density Evolution

Outflow measurement technique (modified from Antonio Ragagnin 2013, Master thesis)



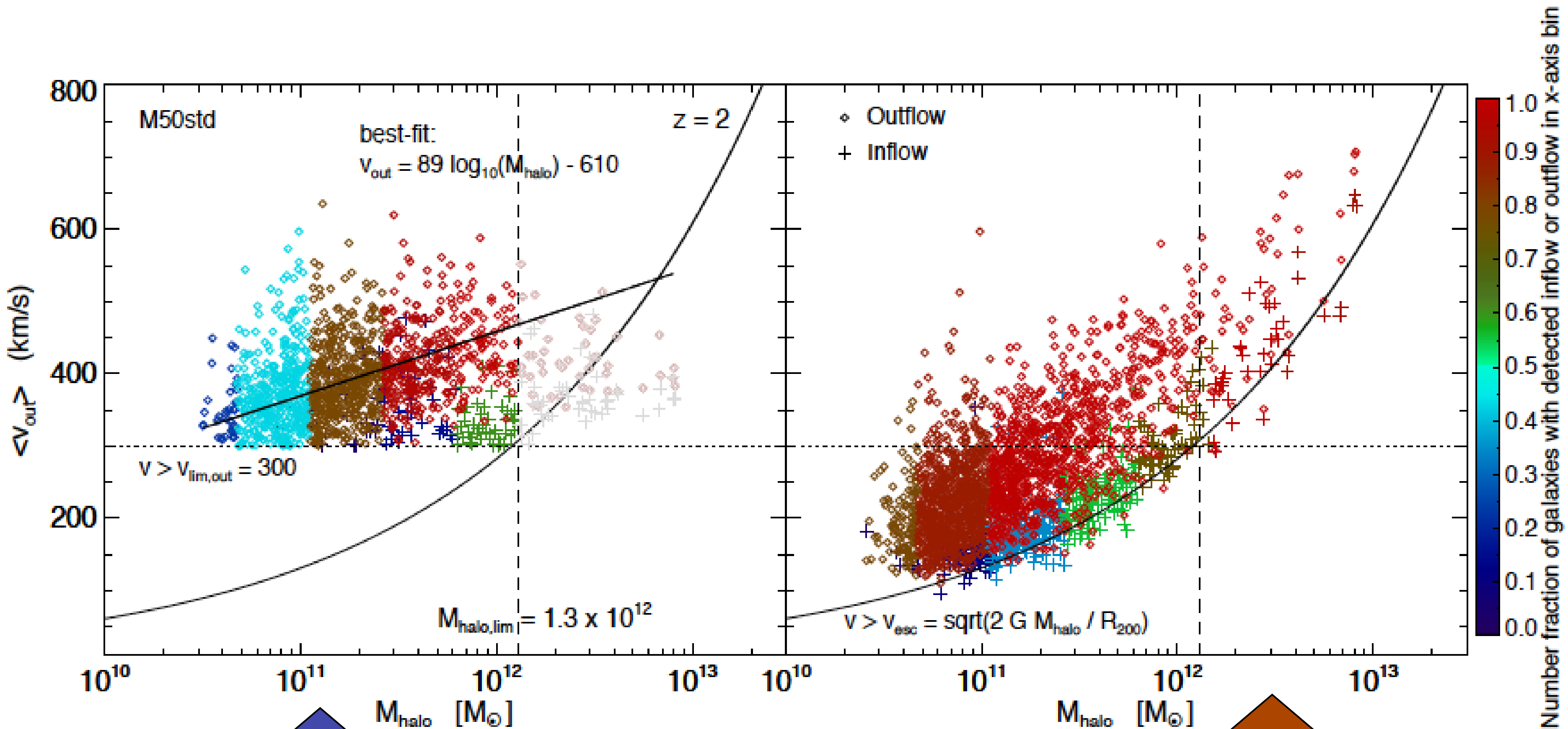
- Transform galaxy coordinates s.t. cold gas disk is rotating in X-Y plane

- Select gas particles:
  - lying inside either cylinder
  - moving at a high-velocity, $|v_z| > V_{limit,outflow}$
  - if $(z*v_z > 0) \Rightarrow$ Outflow
  - if $(z*v_z < 0) \Rightarrow$ Inflow



Gas Vel
E35nw
E25cw
M25std
face-on
edge-on
Gas ρ
Gas T
Gas Z
Stars
kpc
log10 Temperature (° K)
15 + log (Zgas / Z⊙)

Setting the lower velocity threshold for outflow measurement



Fixed value:
Reveals correlations.
Adopt this for
outflow velocity

Escape velocity:
Measures escape outside halo
Adopt this for
mass outflow rate

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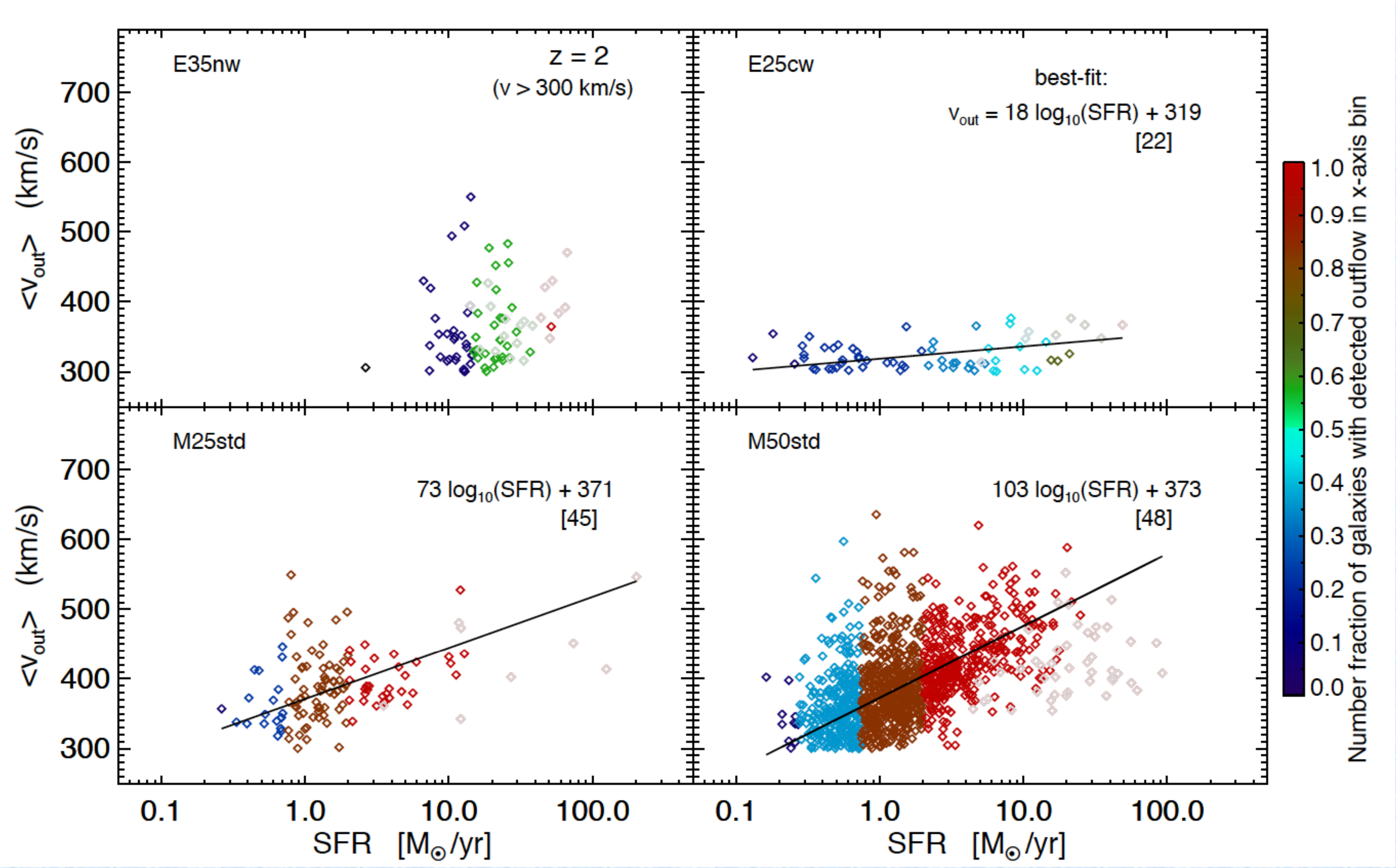
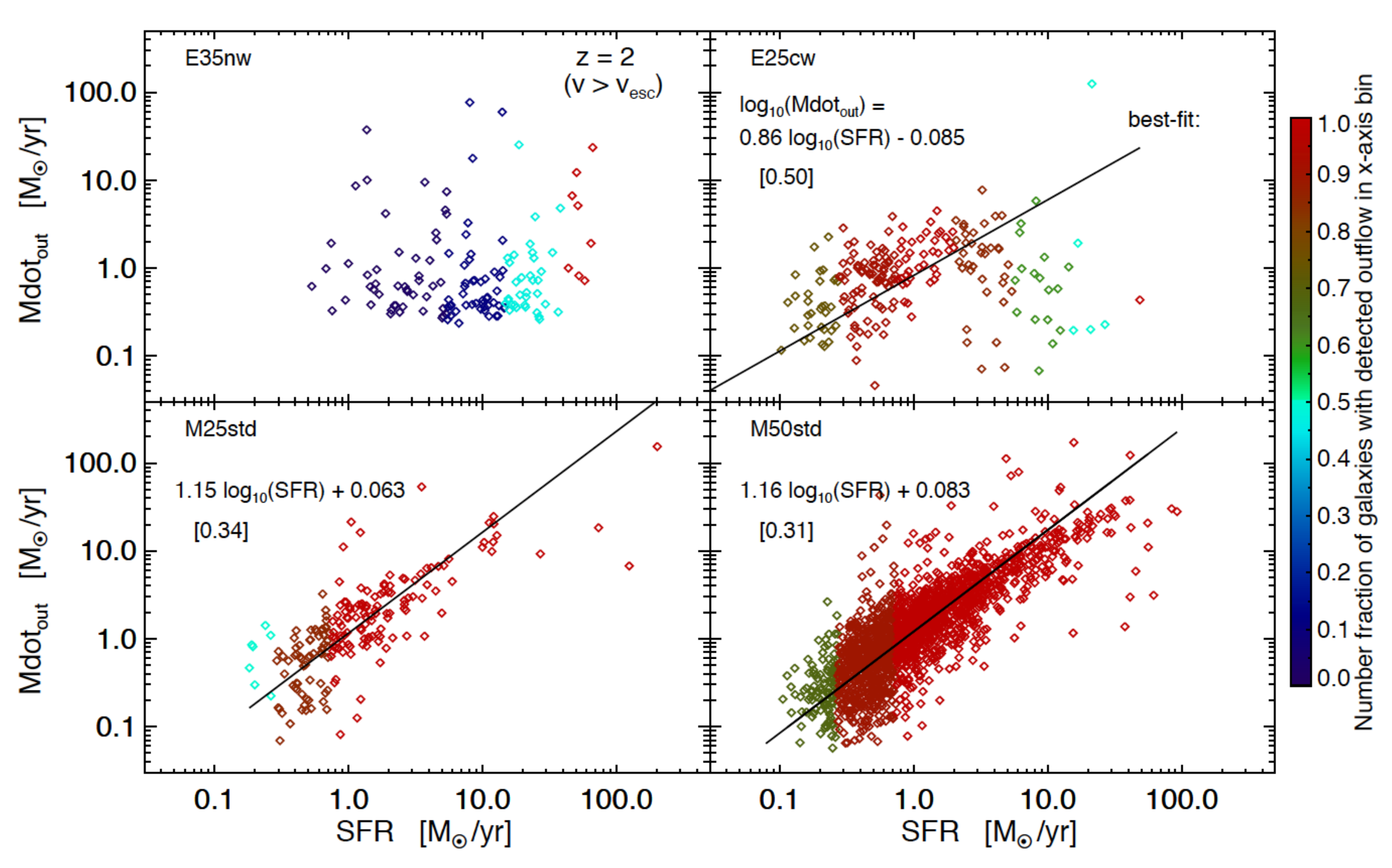
Outflow velocity vs. galaxy SFR

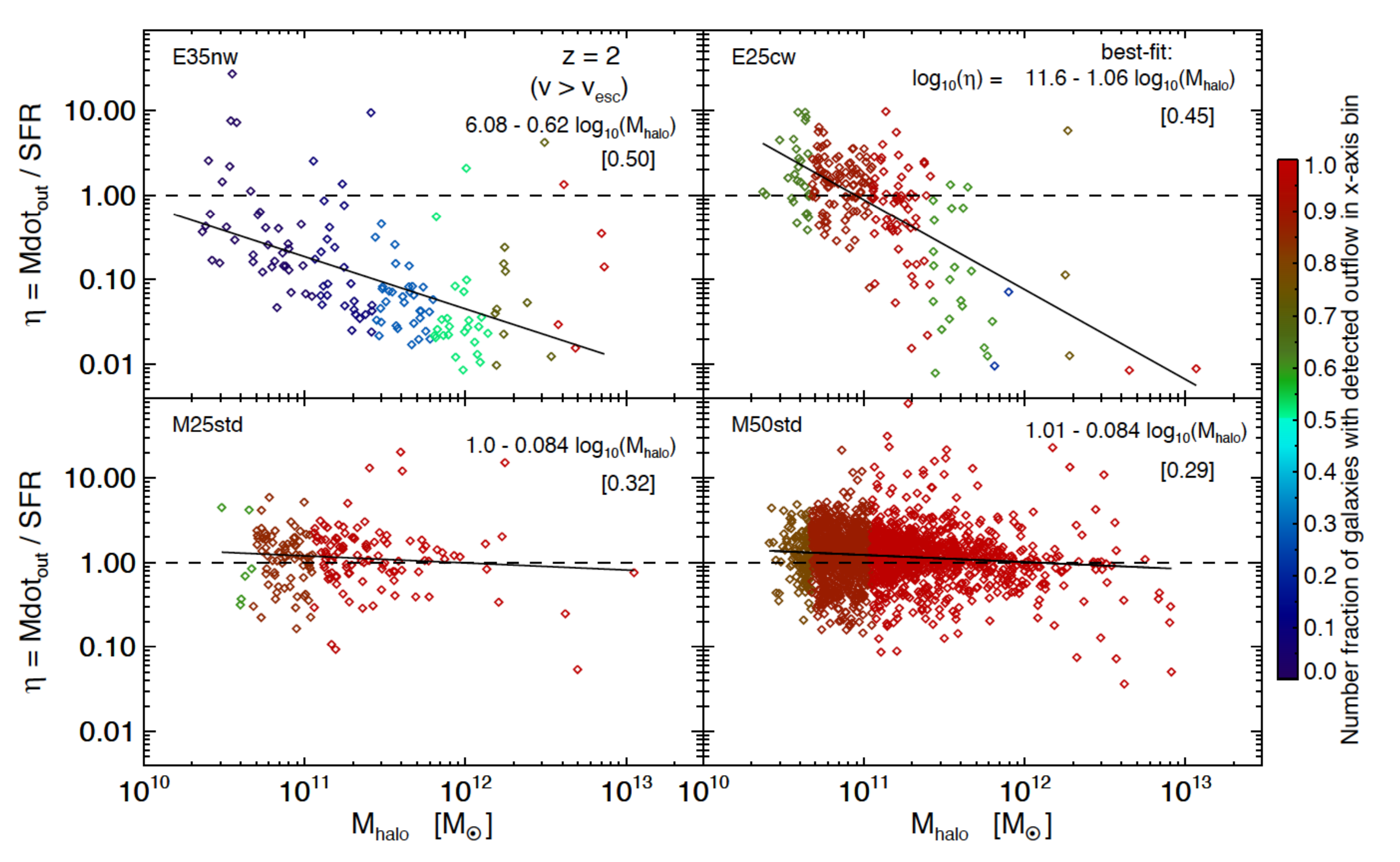


Observation : Martin (2005), Grimes et al. (2009), Banerji et al. (2011), Bordoloi et al. (2013)
- positive correlation of outflow speed with galaxy mass and SFR.

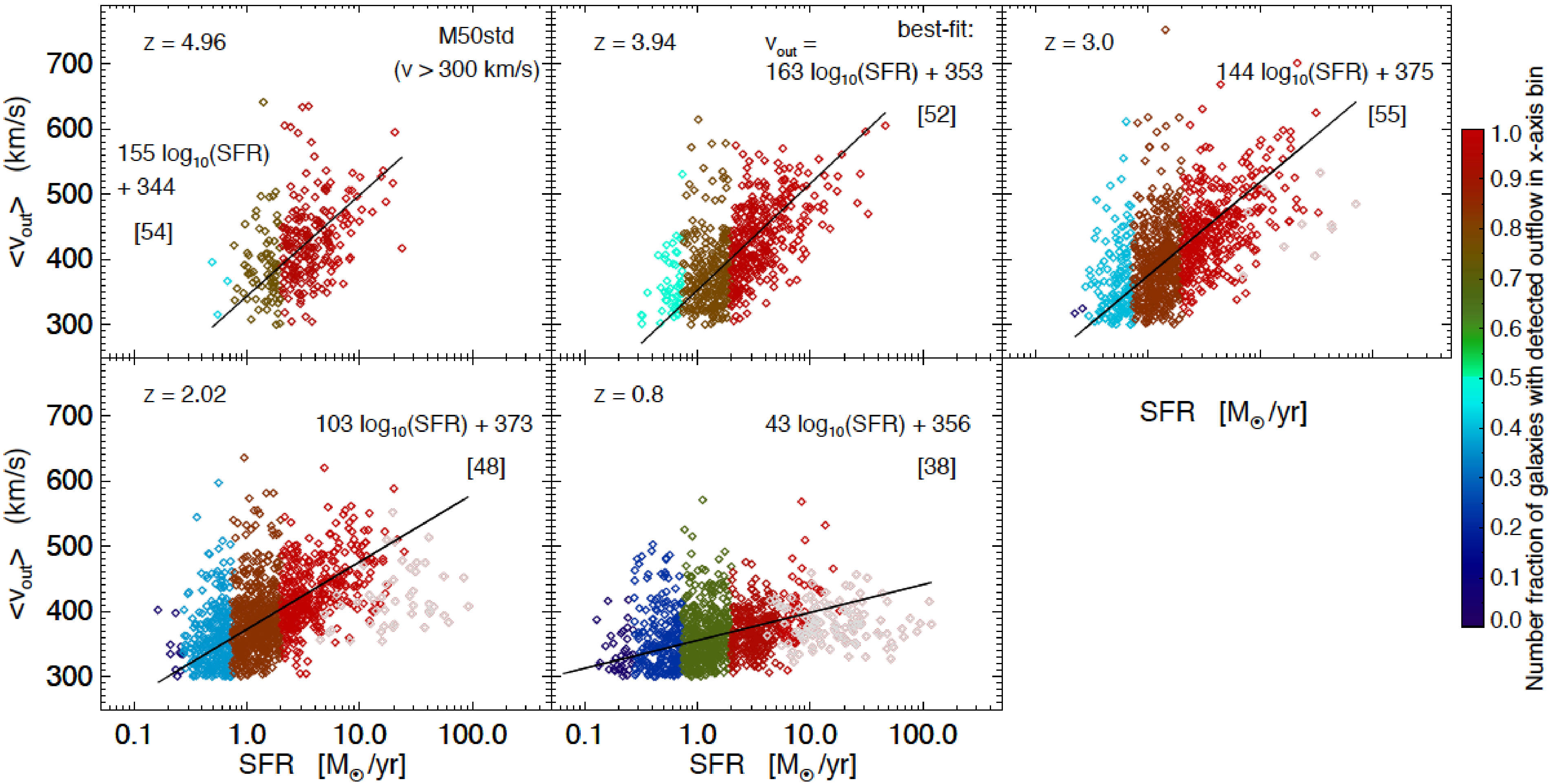
Mass outflow rate vs. galaxy SFR



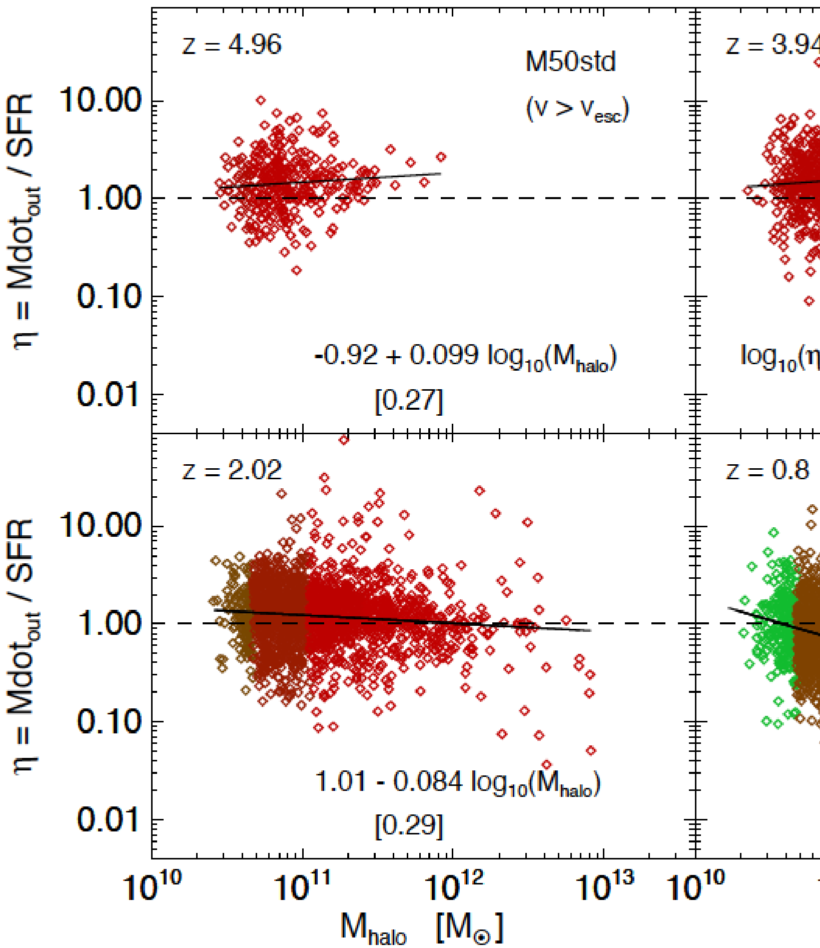
Mass loading factor (η = Mass outflow rate / SFR) vs. halo mass

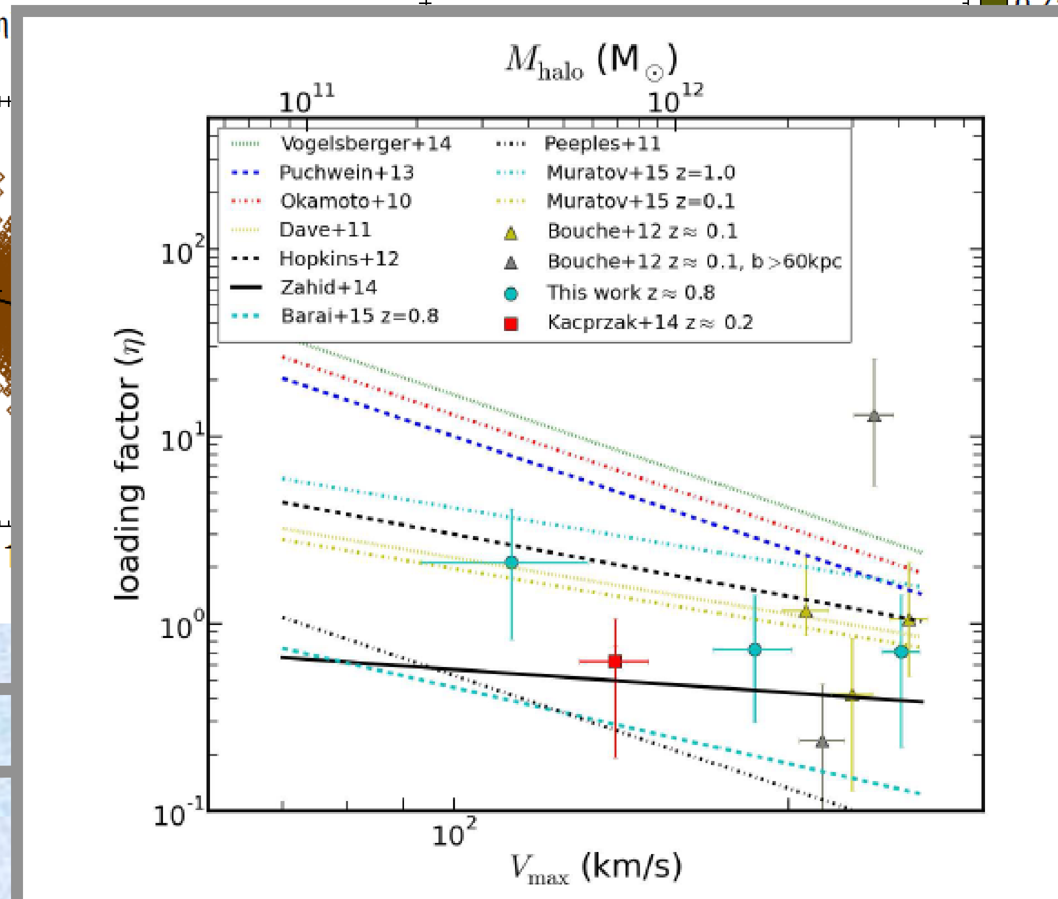


Redshift Evolution of Outflow Velocity vs SFR



Redshift Evolution of Mass-Loading factor vs Halo Mass

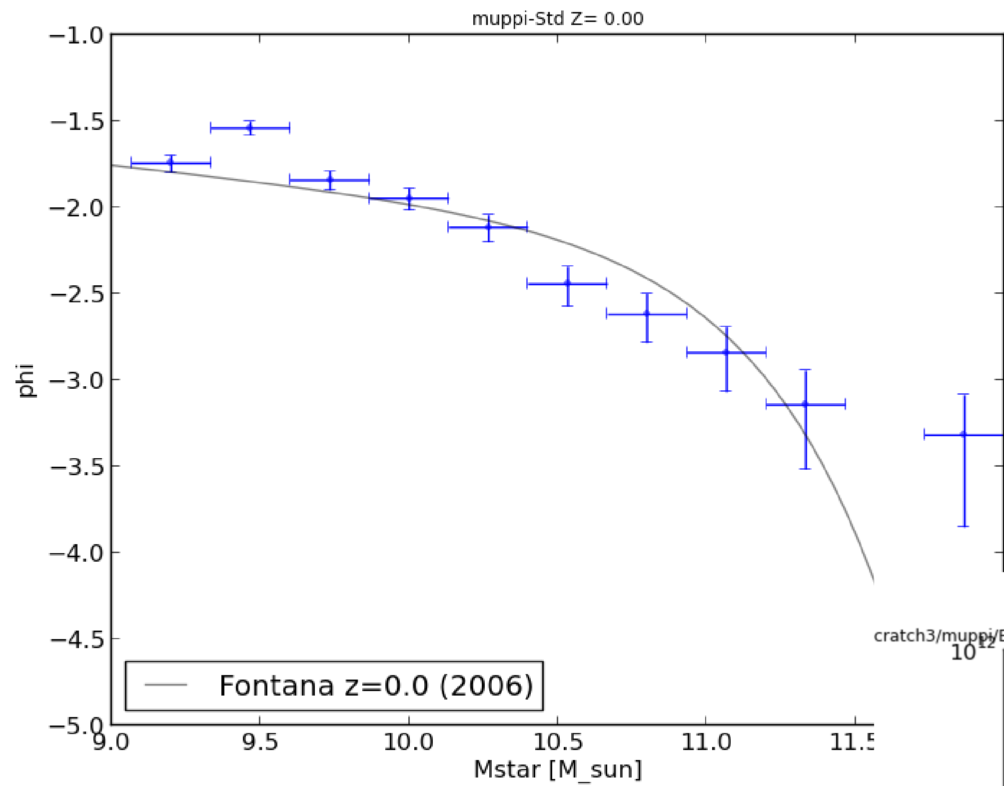


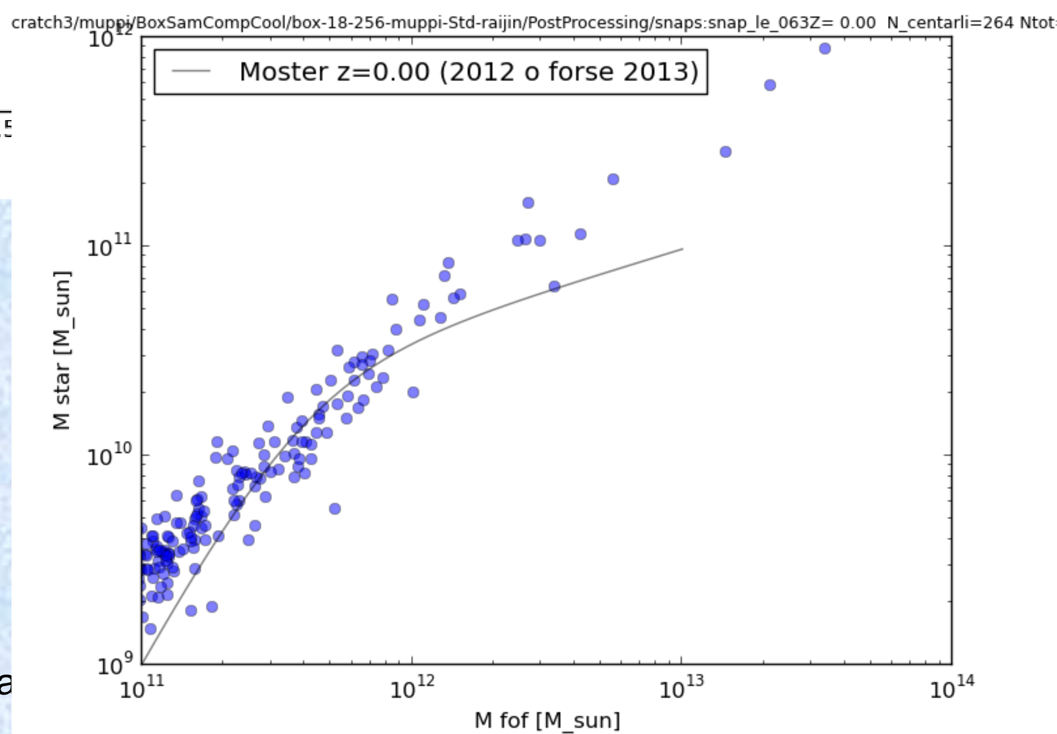


Observations from Schroetter, Bouché, et al. (2015)

Galaxy stellar properties at z=0

(Ragagnin et al. 2015, in prep.)





Simulating realistic disk galaxies with MUPPI, in zoom-in cosmological simulations using moderate resolution *(Murante et al. 2015, MNRAS)*

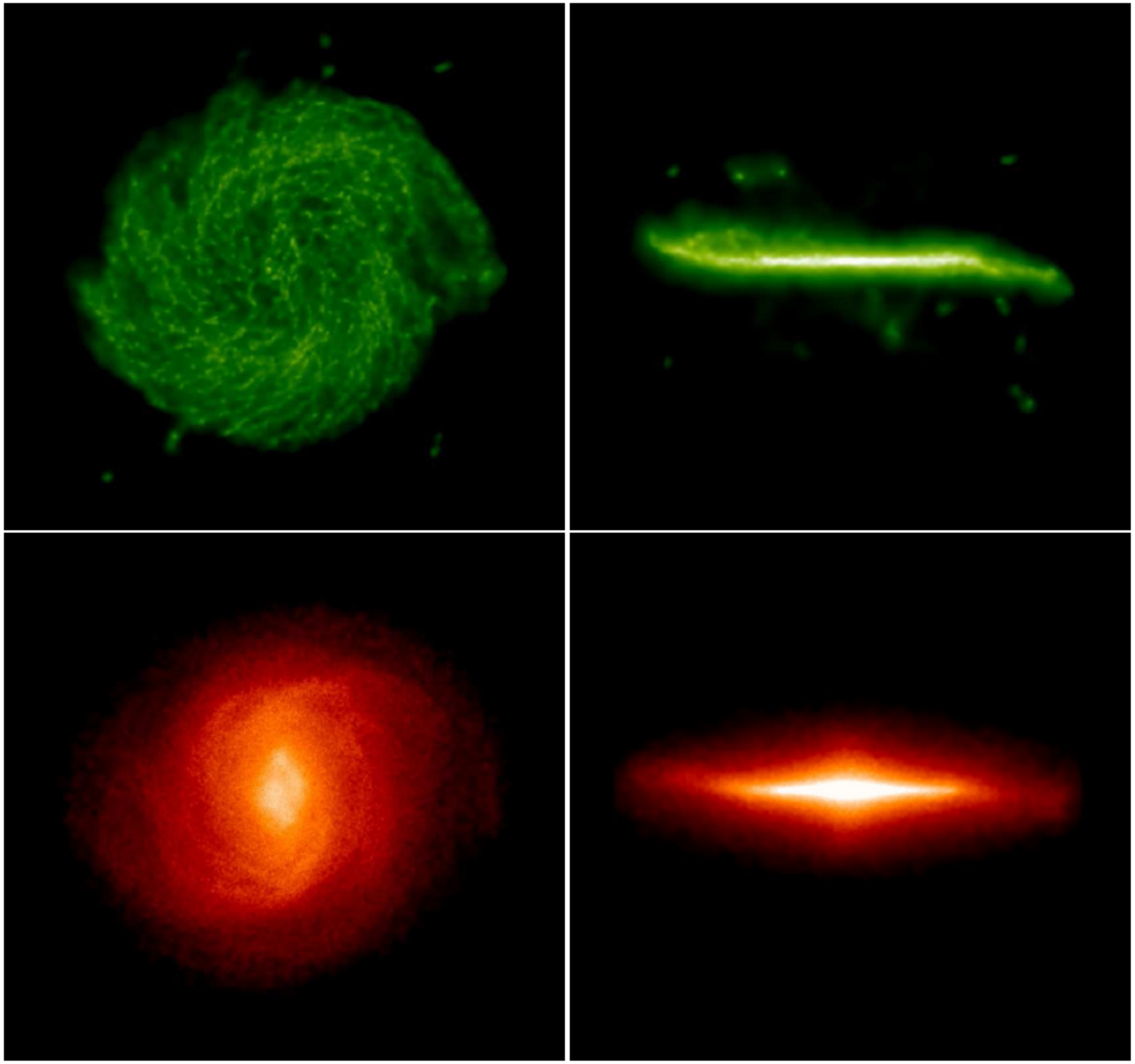
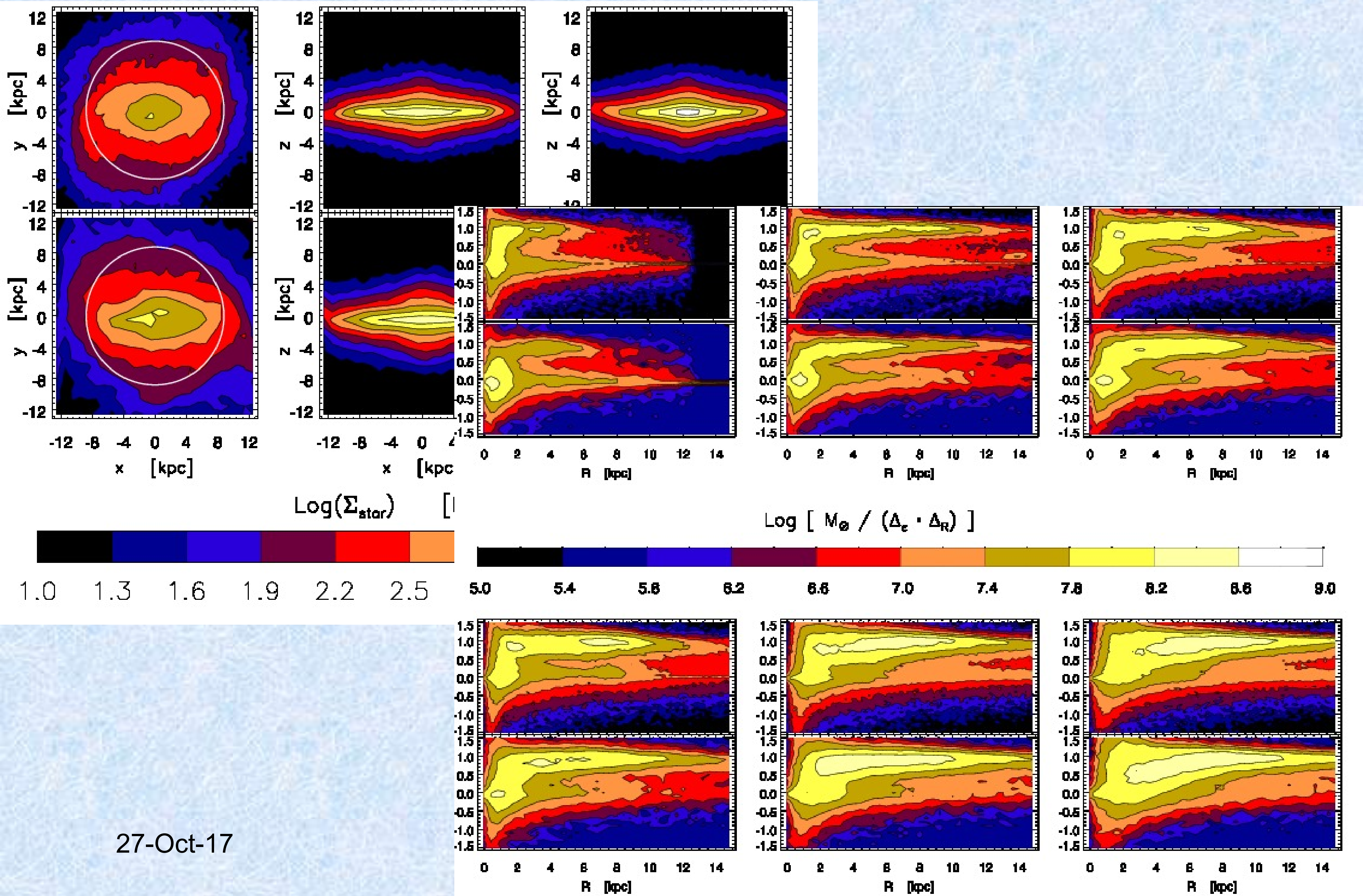


Figure 1. Projected gas (upper panels) and stellar (lower panels) density for the GA2 simulation. The z-axis of the coordinate system is aligned with the angular momentum vector of the gas enclosed within the inner 8 kpc. Left panels show face-on densities, right column shows edge-on densities. Box size is 57 kpc.

Study of barred spiral disk galaxies, in zoom-in cosmological simulations

(Goz et al. 2015, MNRAS)



27-Oct-17

Summary

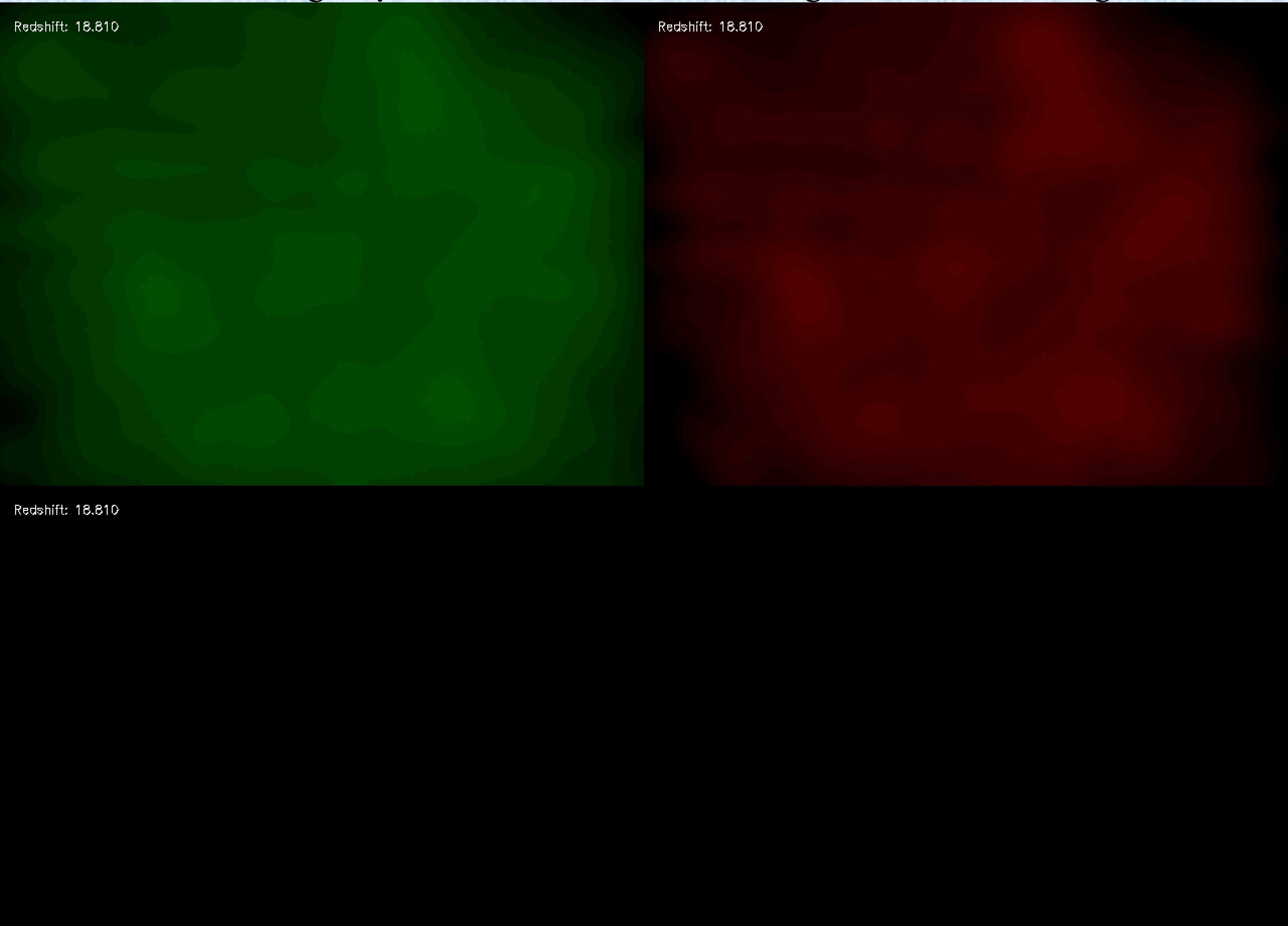
- Cosmological Hydrodynamic Simulations are powerful tool
  - Reproduce the observed large-scale structures
  - Study origin & evolution of galaxies over Hubble time

- Can study impact of galactic winds on galaxy & IGM properties
  - Still far away from self-consistently driving these winds in such sims

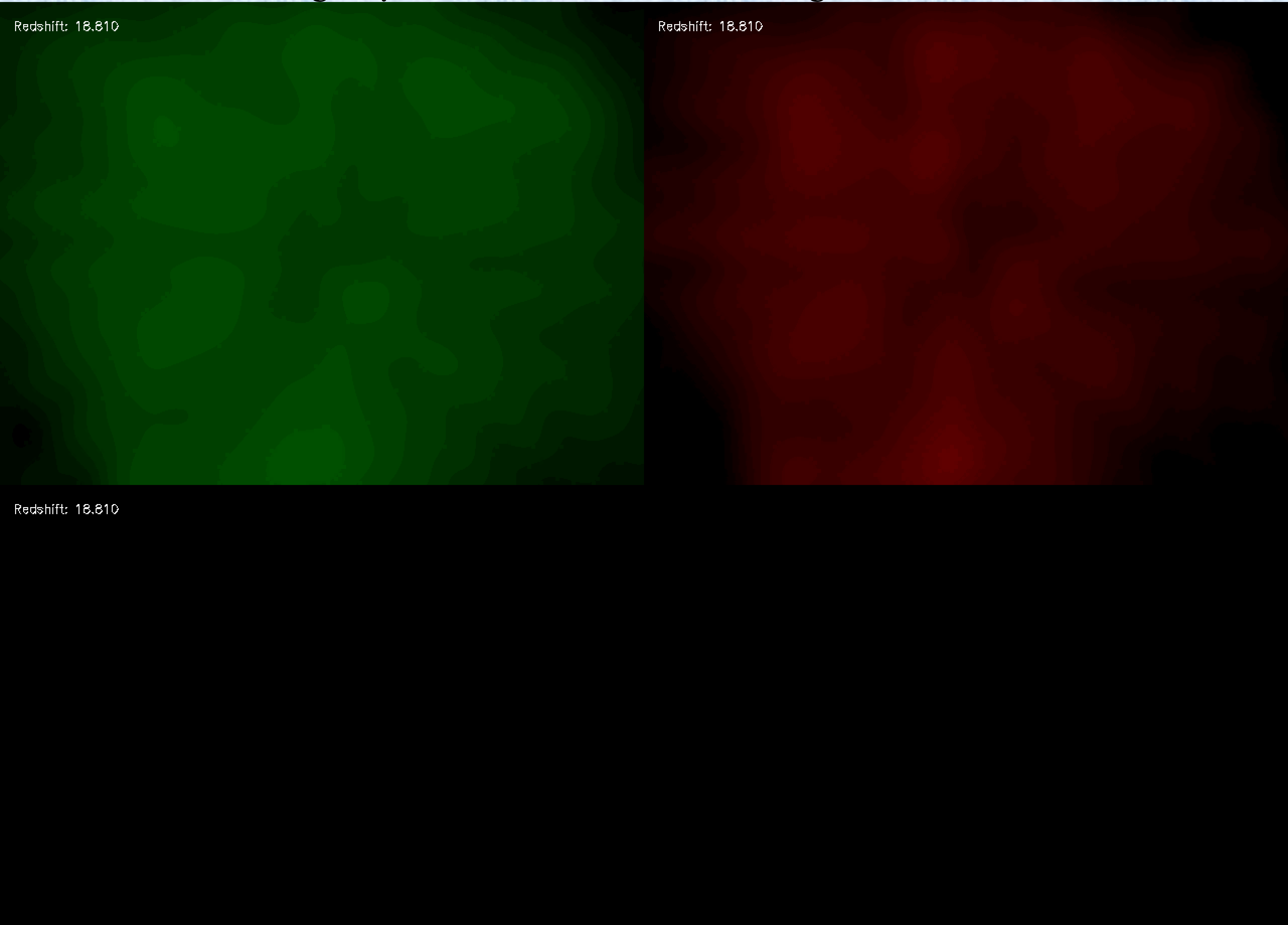
- MUPPI is more physically-motivated sub-resolution model that uses only local properties of gas and generates realistic:
  - Galactic outflows
    - Outflow velocity positive correlation with global galaxy SFR
    - Contant mass-loading value at z=2
    - Redshift evolution predicted over z = 1 - 5
    - Need more observational data
  - Disk galaxies

- Future :
  - Compute further galaxy & IGM observables from sim
  - More physics in code: kinetic AGN feedback, …

Formation of a disk galaxy at z=2. Dark matter, gas, stars. Edge-on view.



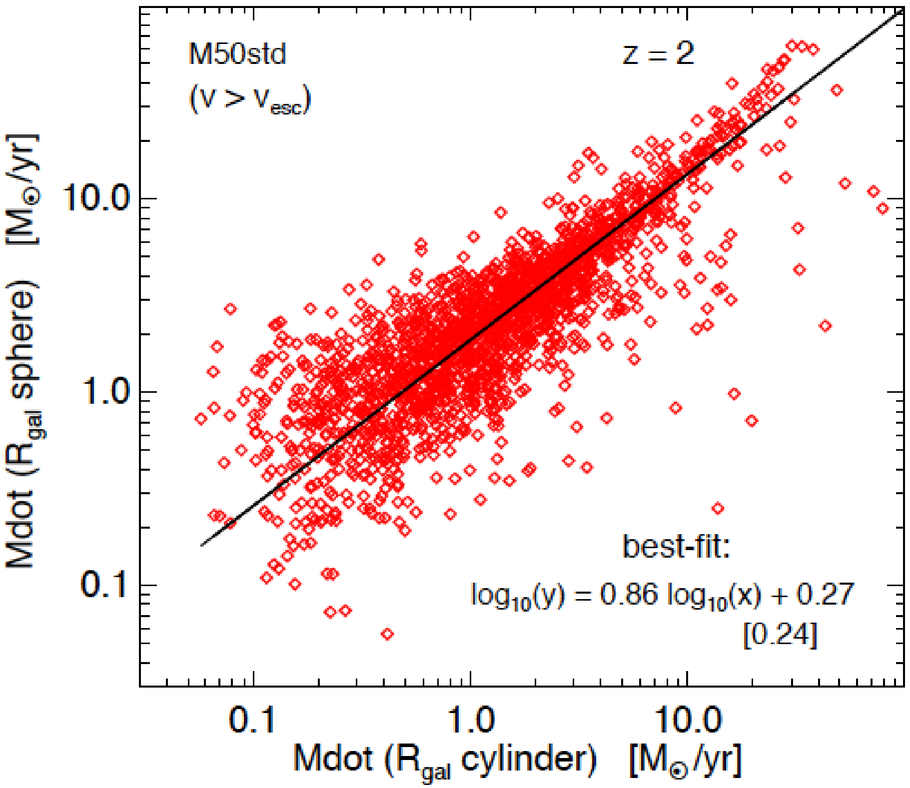
Formation of a disk galaxy at z=2. Dark matter, gas, stars. Face-on view.

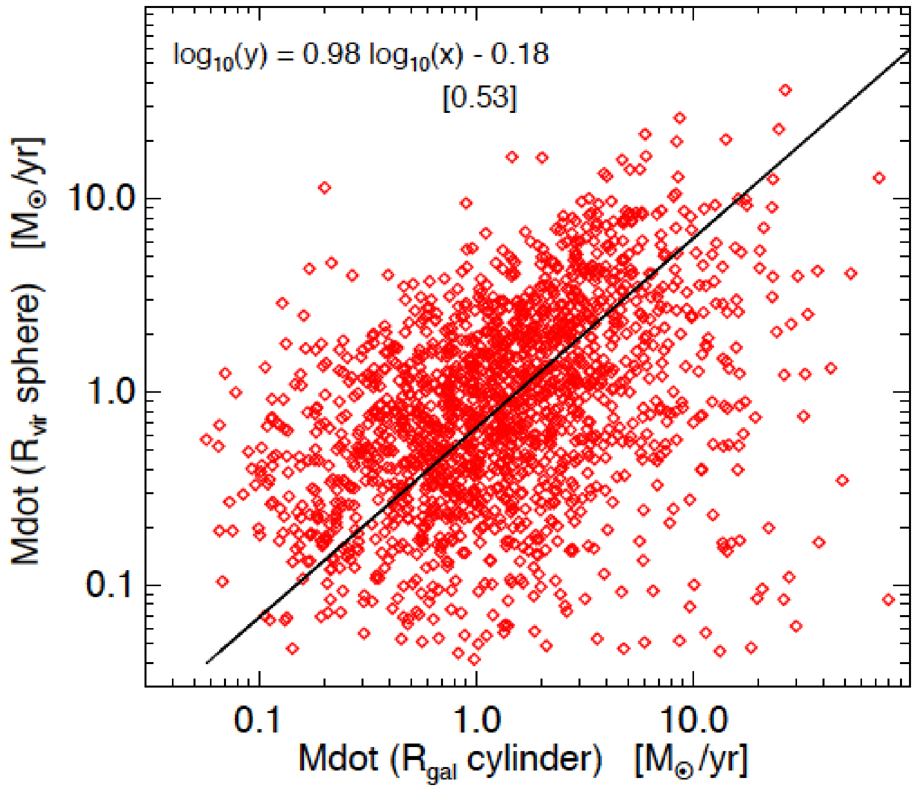


Extra Slides

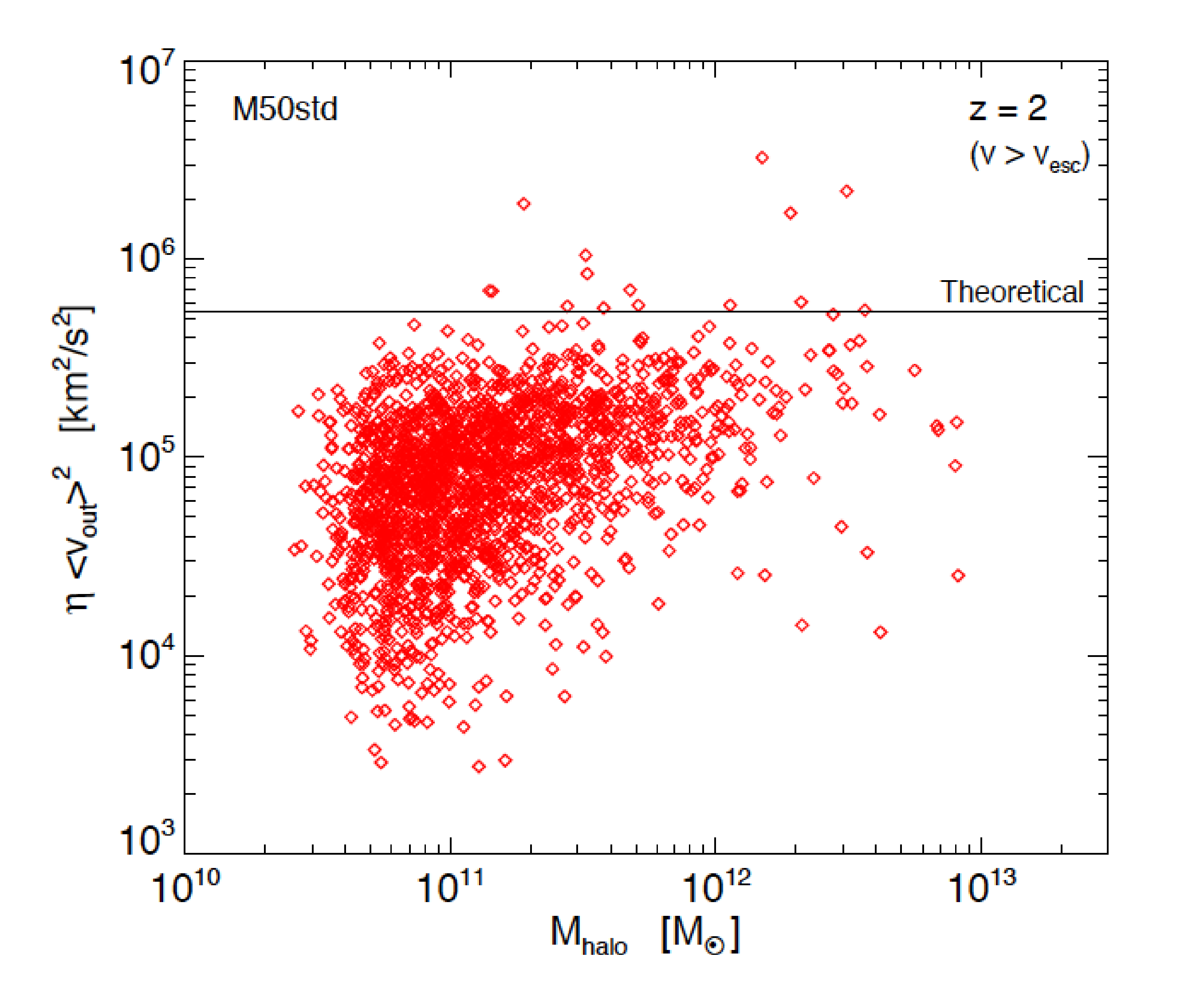
Mass outflow rate at Rgal versus that at Rvir

Method	$N_{\rm outflow}$	$f_{\rm outflow}$
At $R_{\rm gal}$ using $\lvert v_r \rvert > v_{\rm esc}(R_{\rm gal})$, in a cylinder	1842	0.93
At $R_{\rm gal}$ using $\lvert v_r \rvert > v_{\rm esc}(R_{\rm gal})$, in a sphere	1936	0.97
At $R_{\rm vir}$ using $\lvert v_r \rvert > v_{\rm esc}(R_{\rm vir})$, in a sphere	1734	0.87

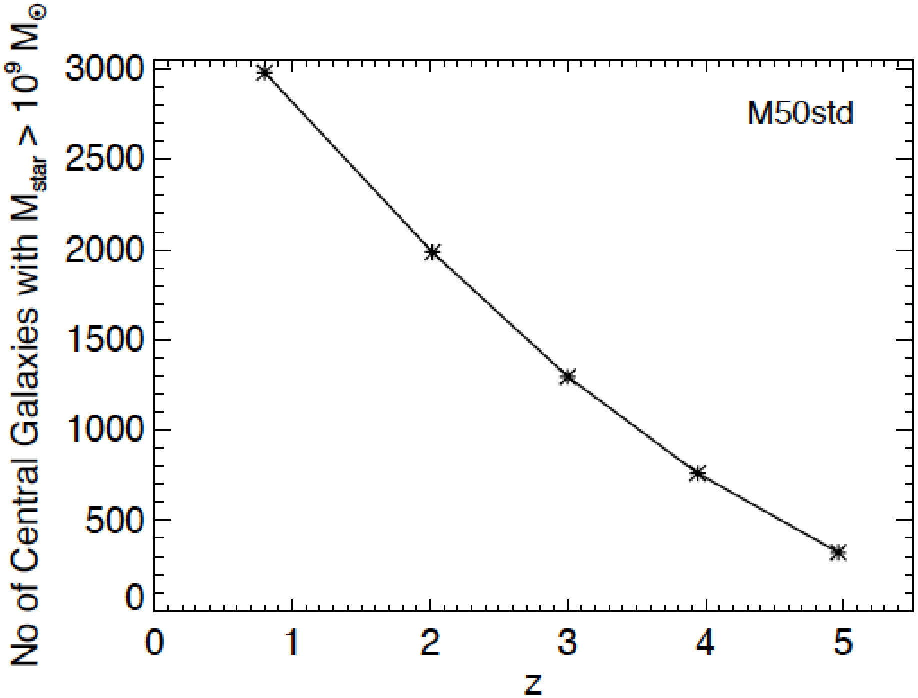


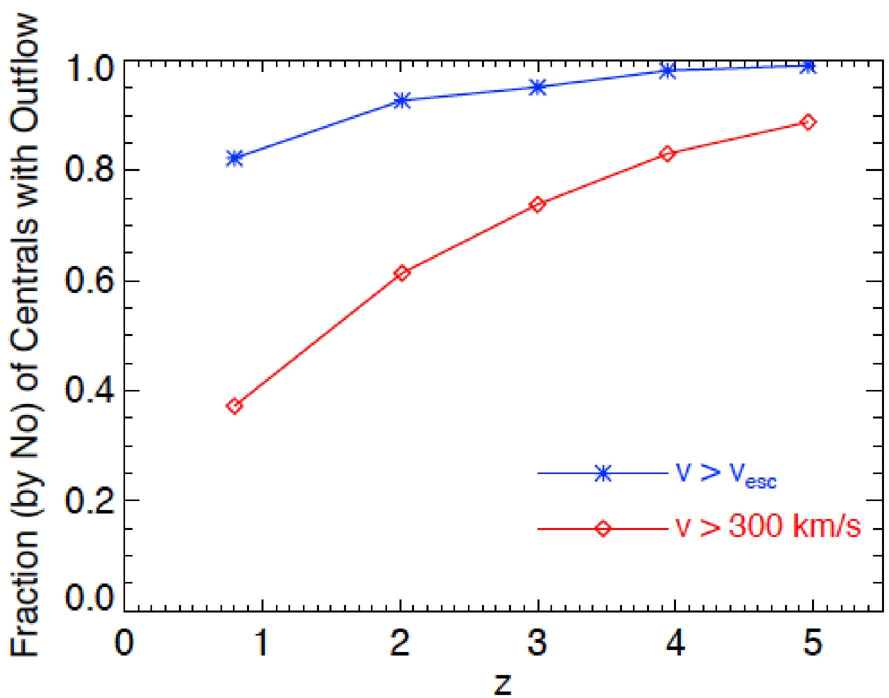


Prediction with Theoretical Estimate of the MUPPI model



Redshift Evolution of Outflow Number Fraction





Observation : Karman et al. (2014) - incidence of large-velocity outflow higher at z~3 than at z<1.