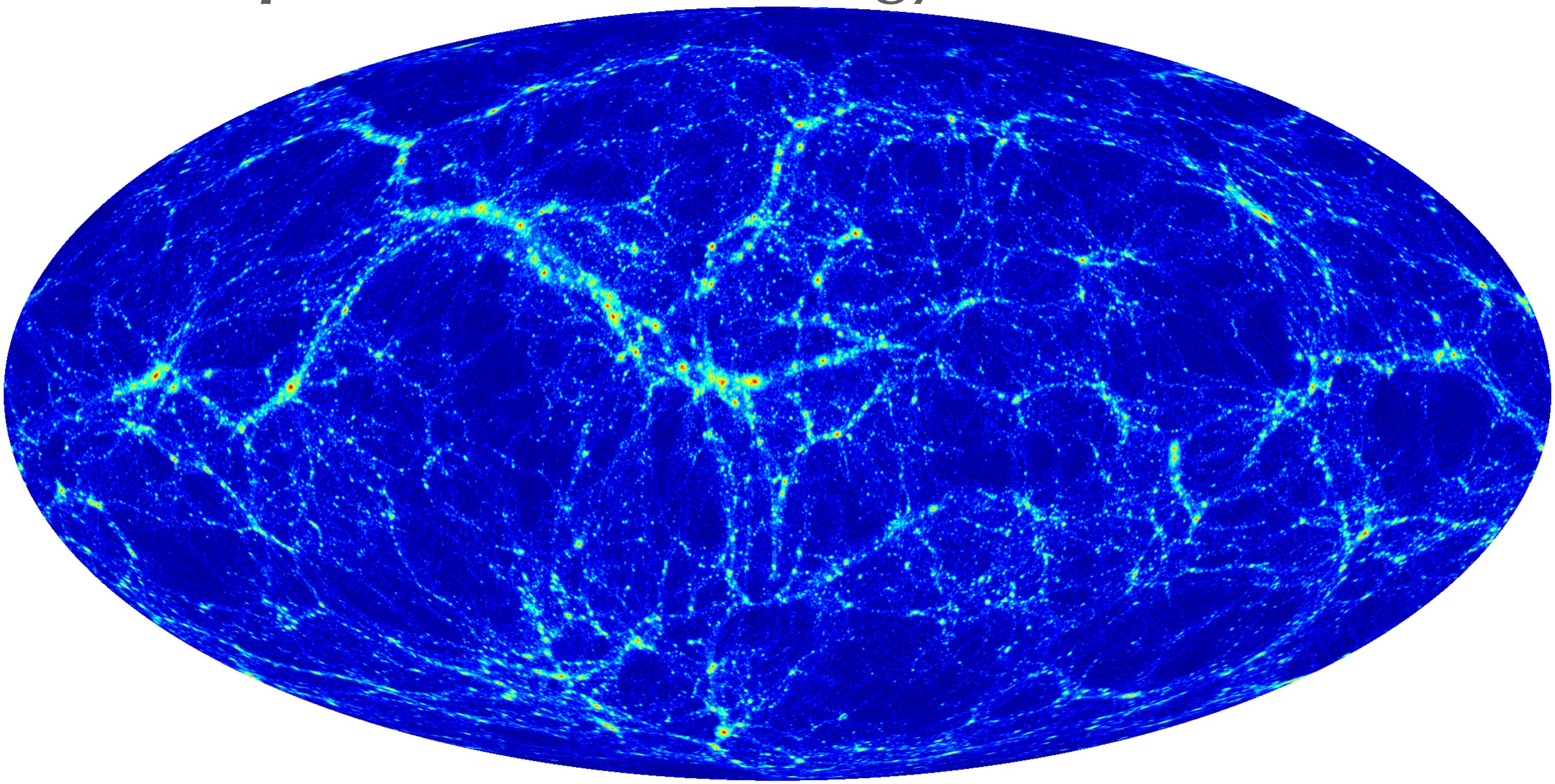
Toward High-Fidelity Synthetic Skies: Computational Cosmology in Adolescence



August (Gus) Evrard

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Departments of Physics and Astronomy

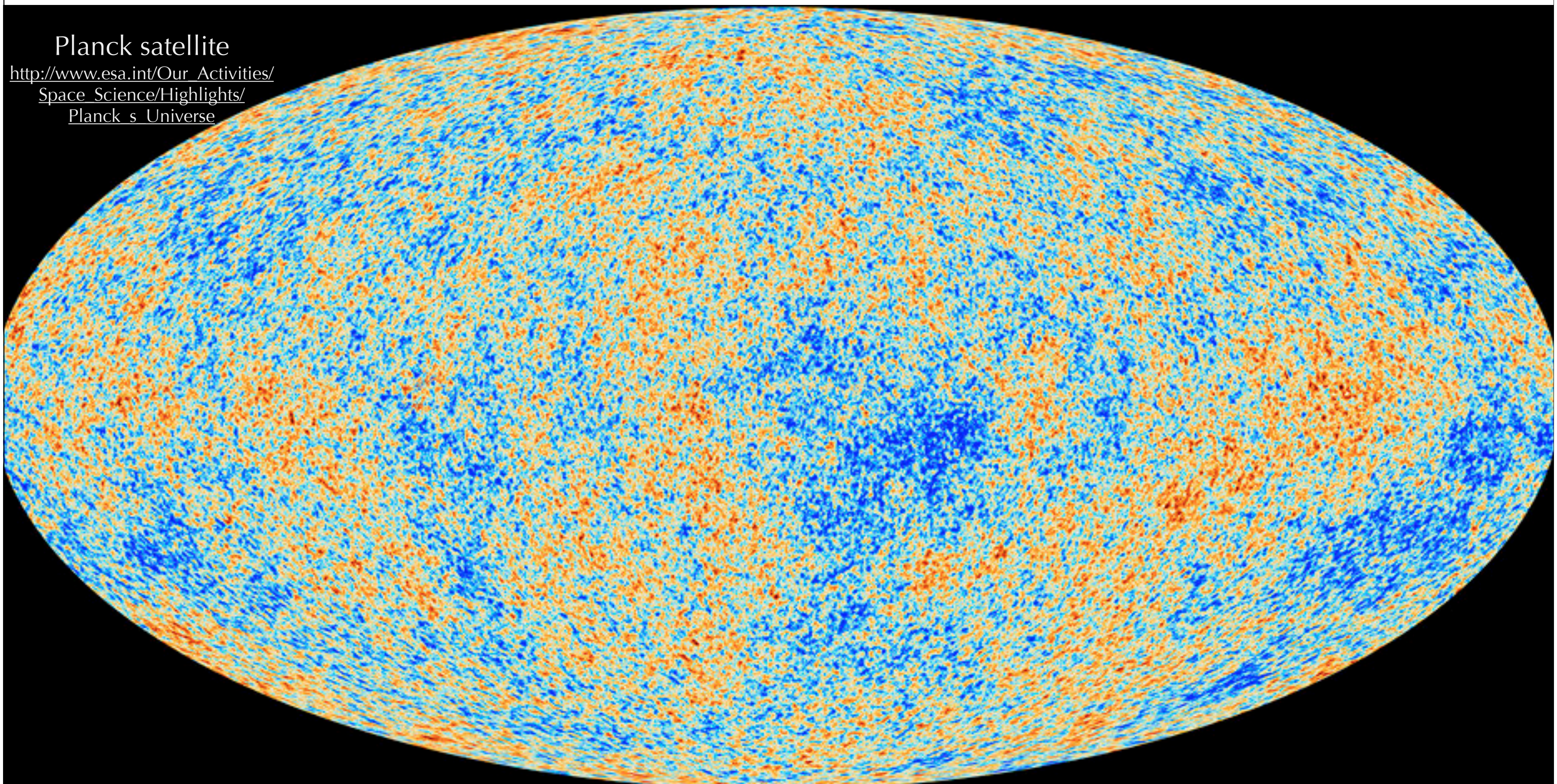
Michigan Center for Theoretical Physics

University of Michigan

outline

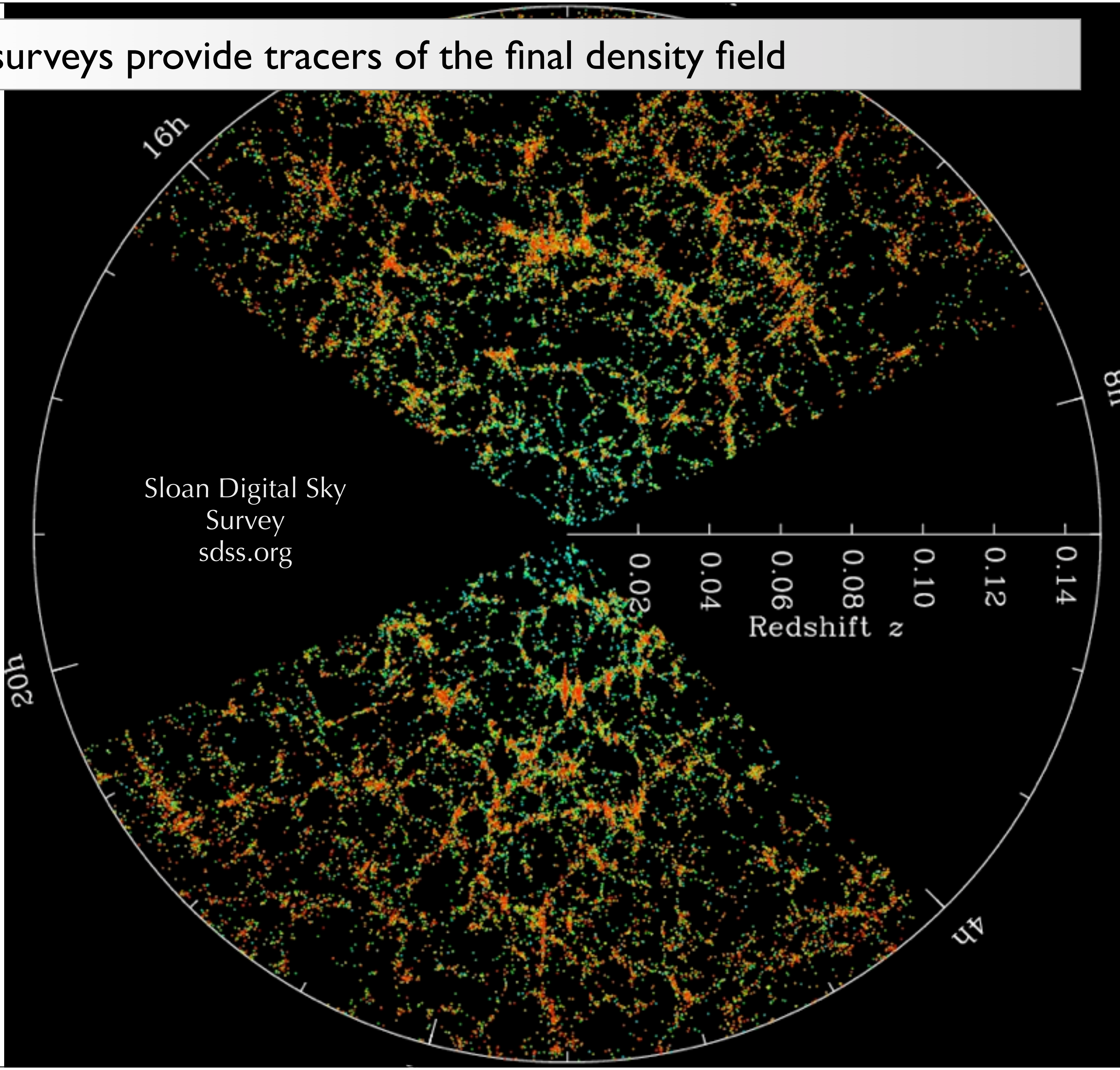
* cosmic structure as a (complex) initial-value problem
  – clusters of galaxies as a motivational example
* the N-body era
  – broad and deep expertise in non-linear behavior
* beyond N-body: tracing baryon evolution
  – galaxy formation: many recipes, improving insights
* synthetic skies for Dark Energy Survey + XMM-XXL
  – empirical, statistical approaches circumvent complex astrophysics
* post-adolescent thoughts
  – community / identify / infrastructure for 21st C. workflows

the cosmic microwave background provides initial conditions



Planck satellite
http://www.esa.int/Our_Activities/Space_Science/Highlights/Planck_s_Universe

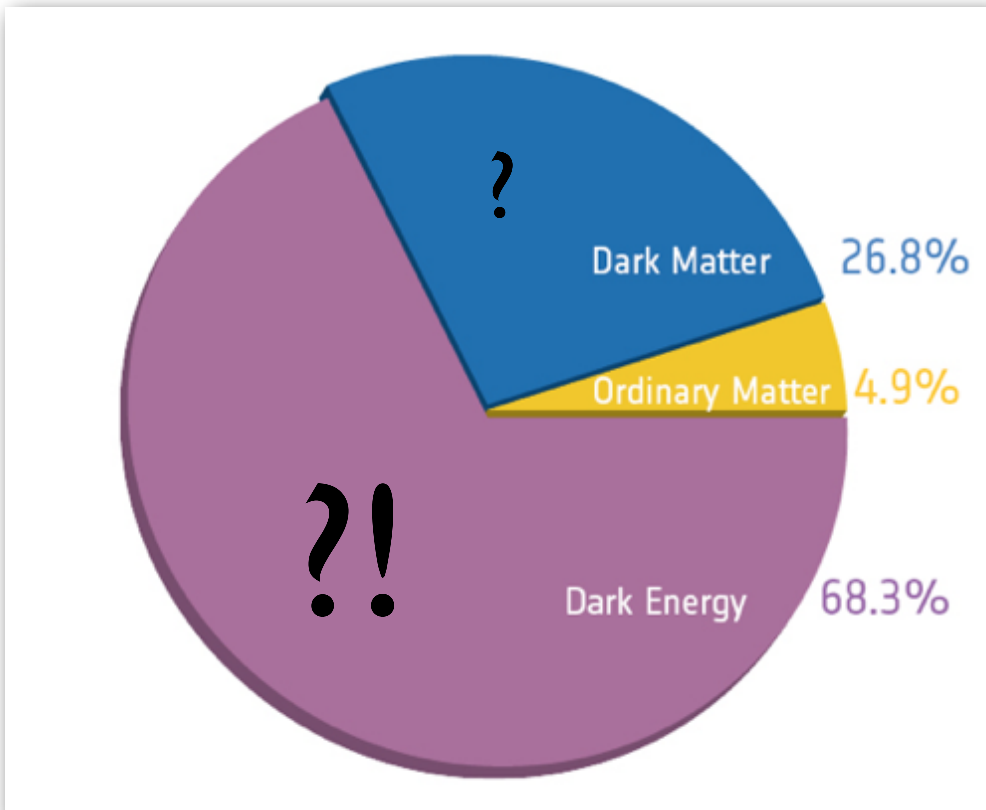
galaxy surveys provide tracers of the final density field



constraints on the energy density of today's universe

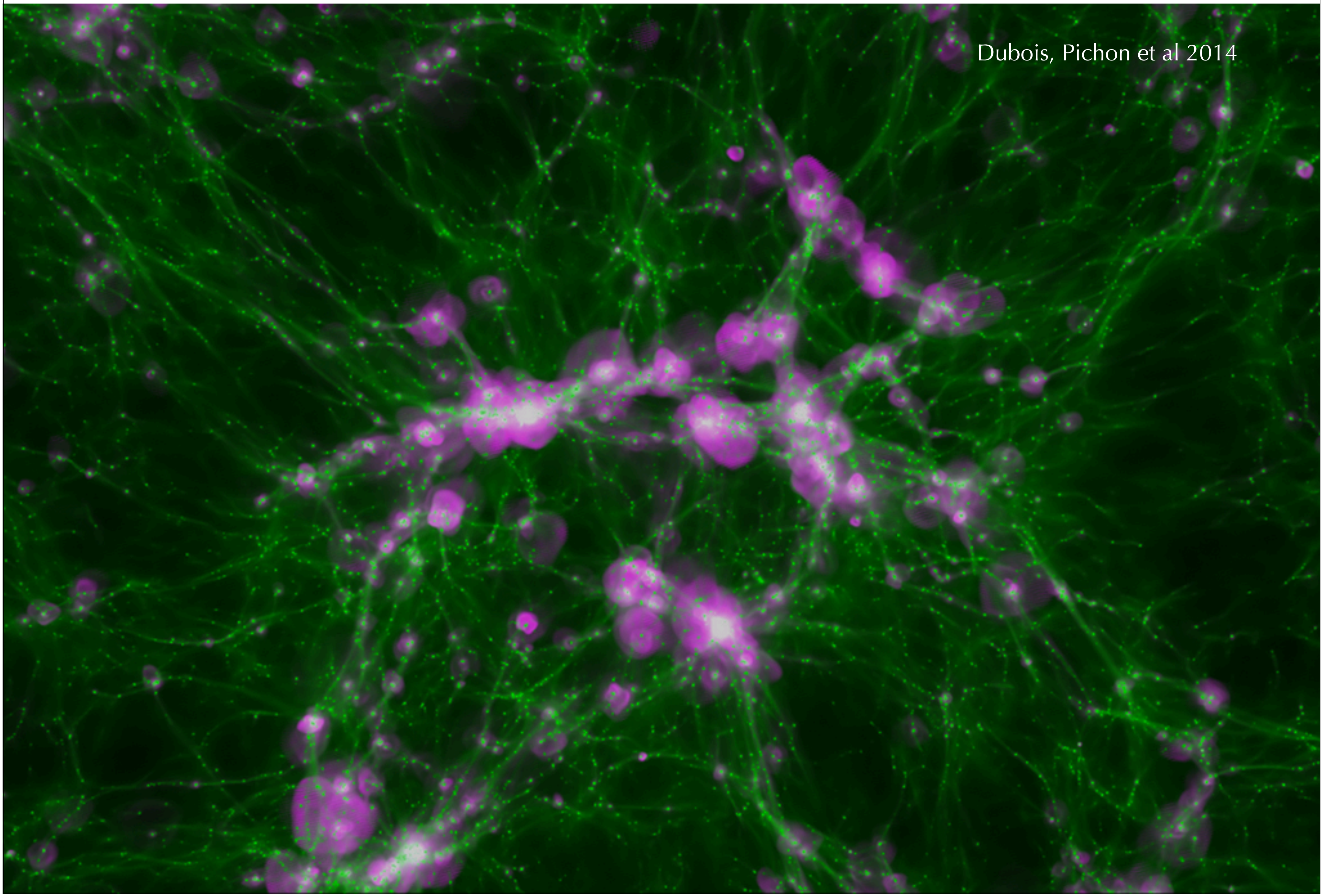
Combined statistical constraints from:

- Cosmic Microwave Background (CMB)
- Large-Scale Structure (LSS) traced by galaxies
- local cosmic expansion rate (Hubble constant)

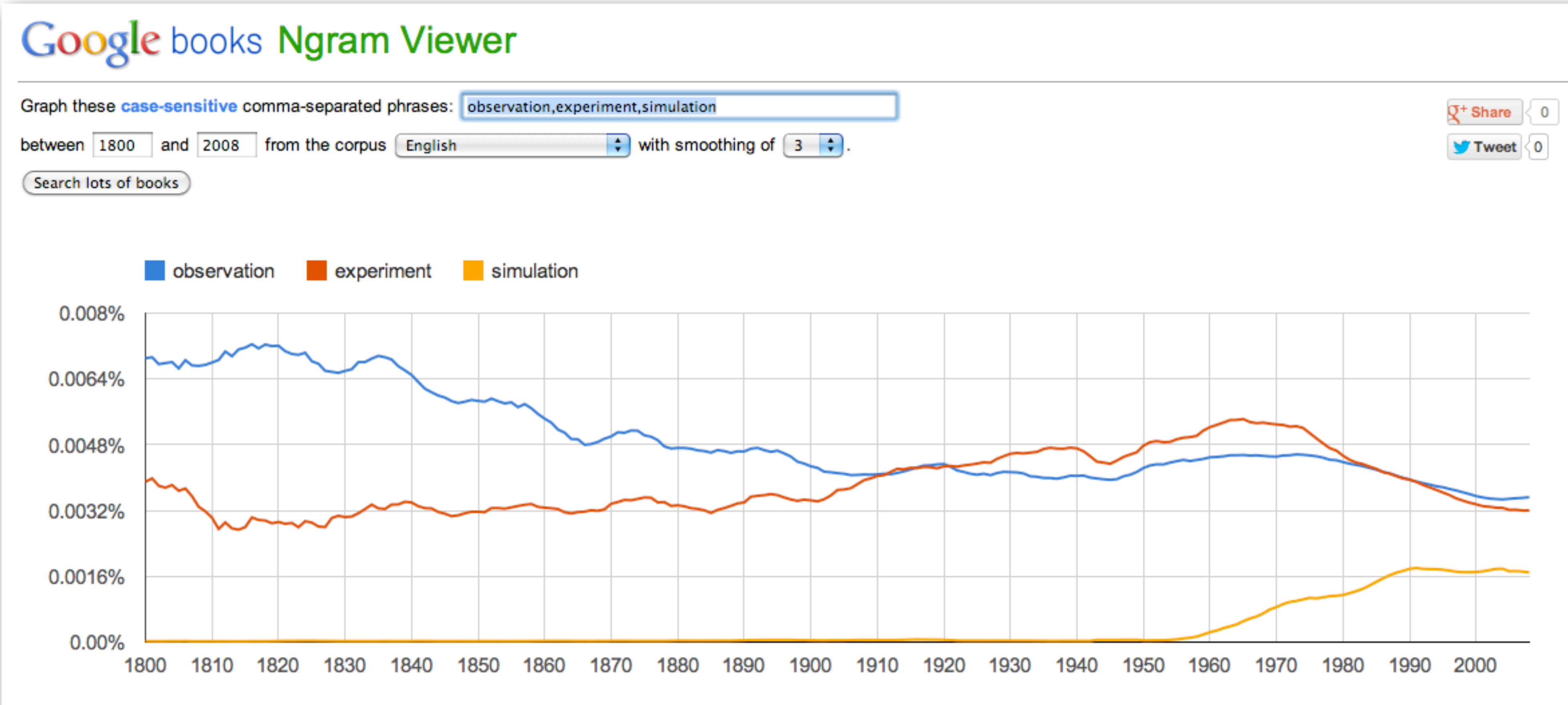


Gravity amplifies
tiny random noise into a
cosmic web
of *hierarchically-evolving*
large-scale structure,
which features
locally bound, ellipsoidal
halos
as an essential component.

example of baryonic web at z~1.5 http://spine-public.projet-horizon.fr



frequency of "simulation" in the Google Books corpus



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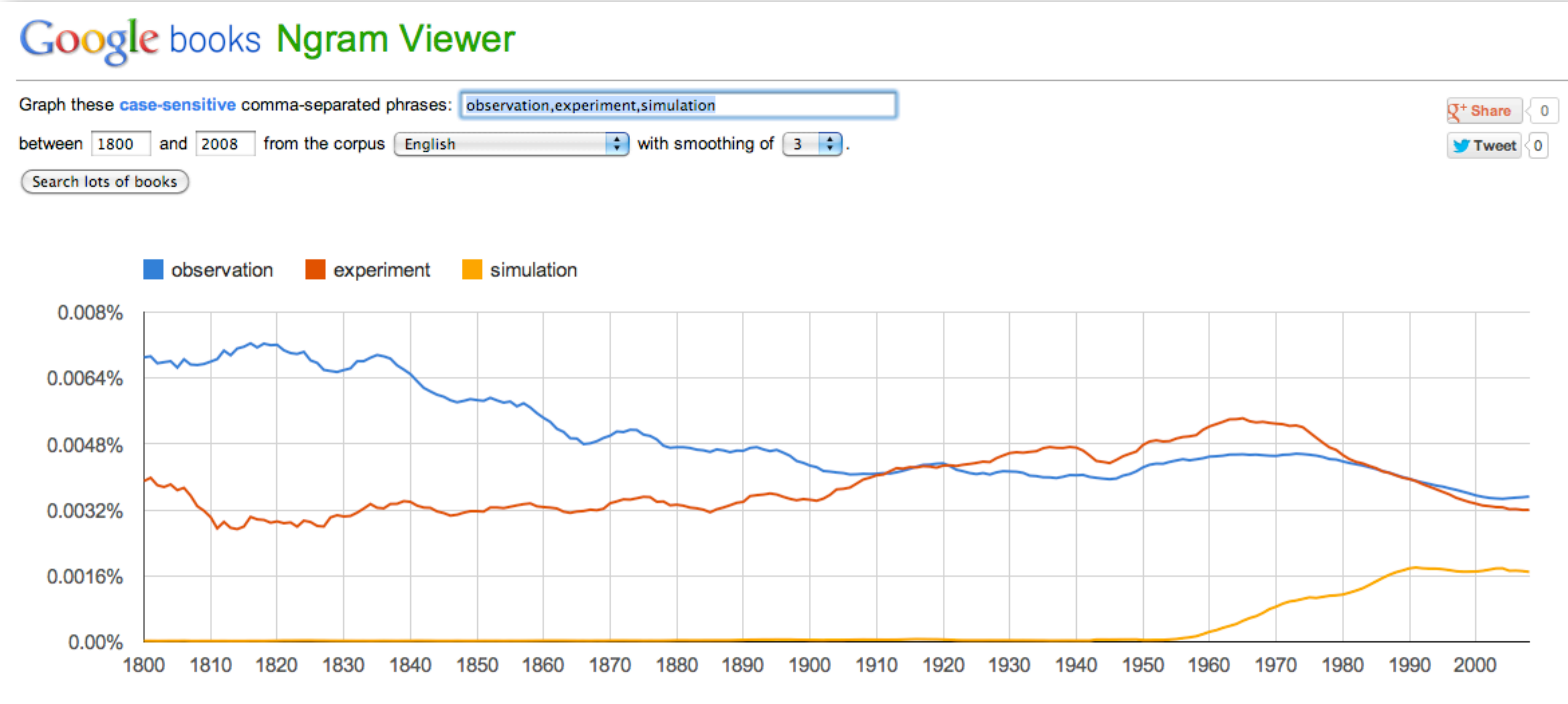
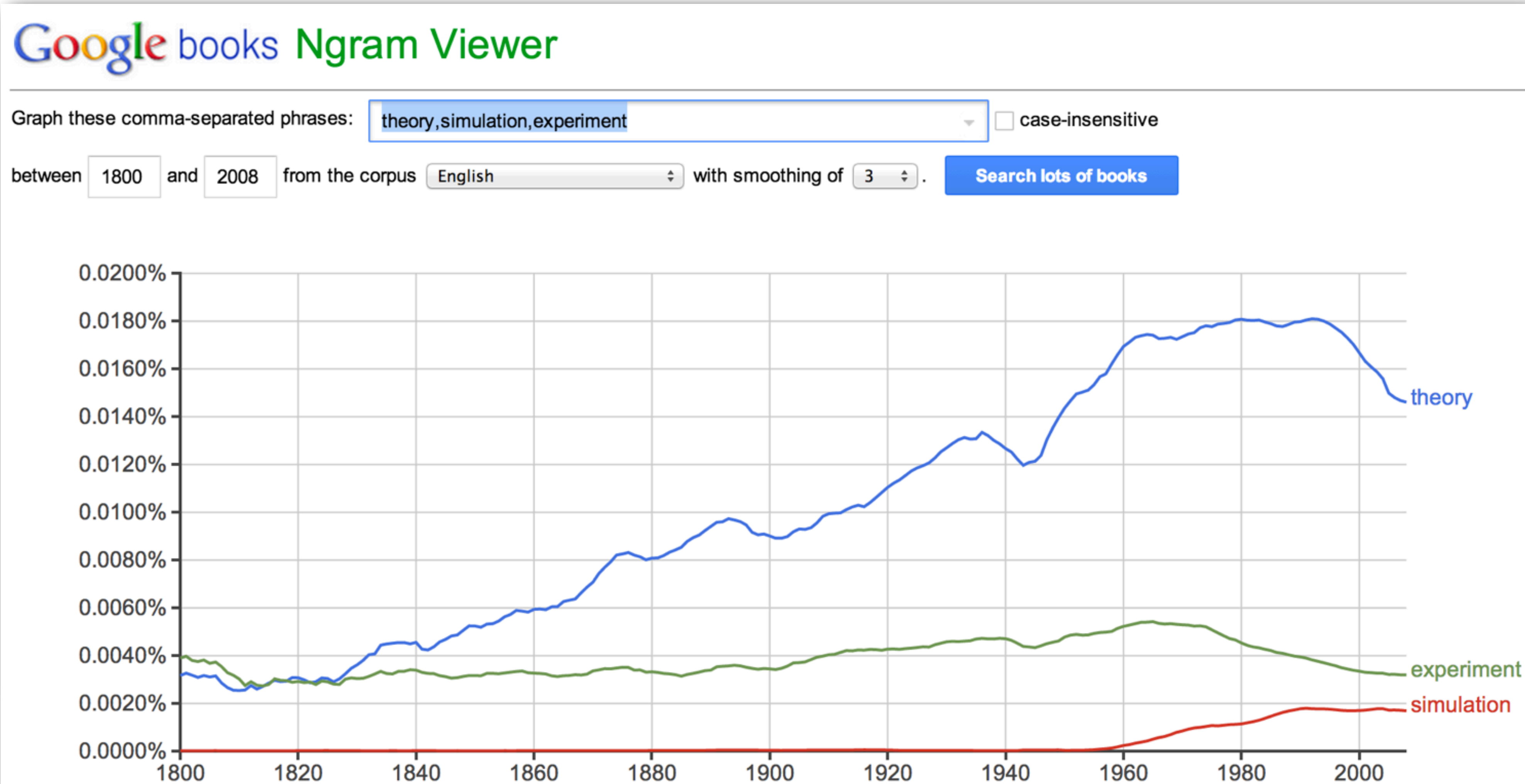


Relative to observation and experiment, simulation is in an adolescent phase.

frequency of "simulation" in the Google Books corpus



clusters of galaxies =
manifestations of massive,
dark matter-dominated halos

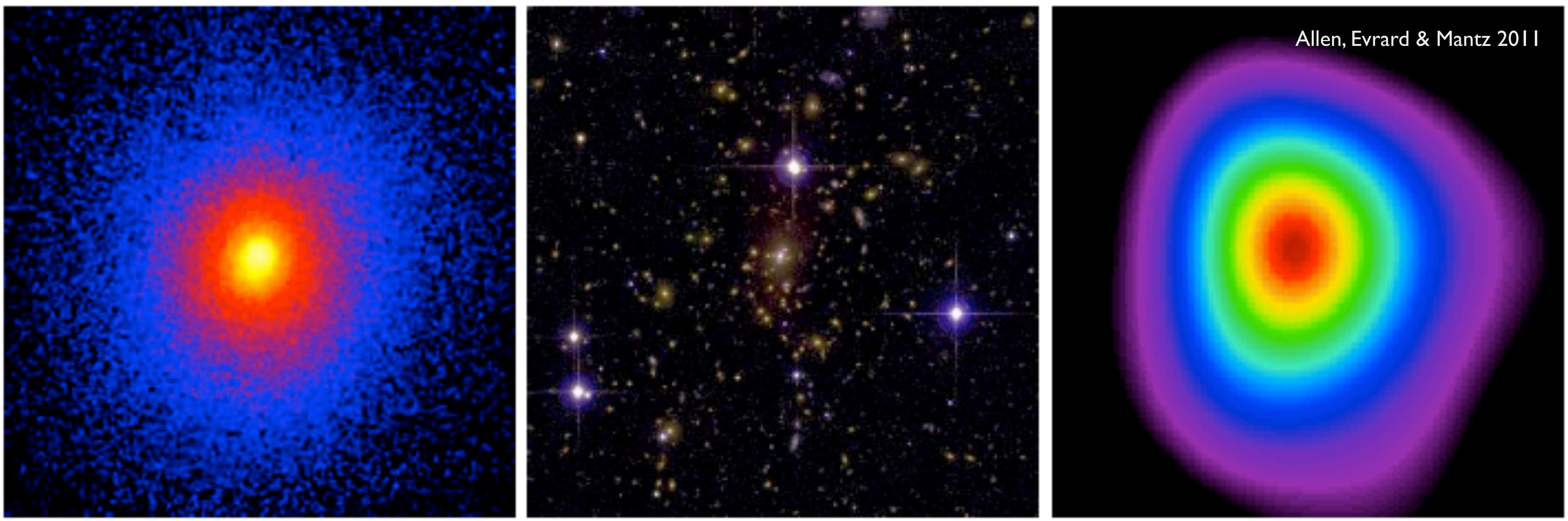
clusters of galaxies are simple (relative to galaxies) astrophysical laboratories

* baryons in M>~1e14 Msun **halos** reside mainly in a

hot, intracluster medium (ICM)

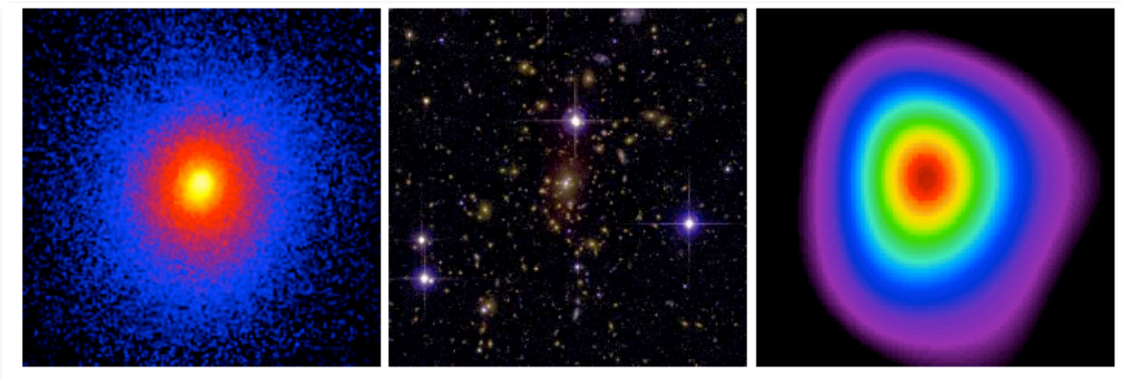
– hot gas outweighs baryons in stars by factors >~5

Abell 1835 (z=0.25) 5.2 arcmin ~ 1.2 Mpc images



X-ray optical mm (Sunyaev-Zel'dovich)

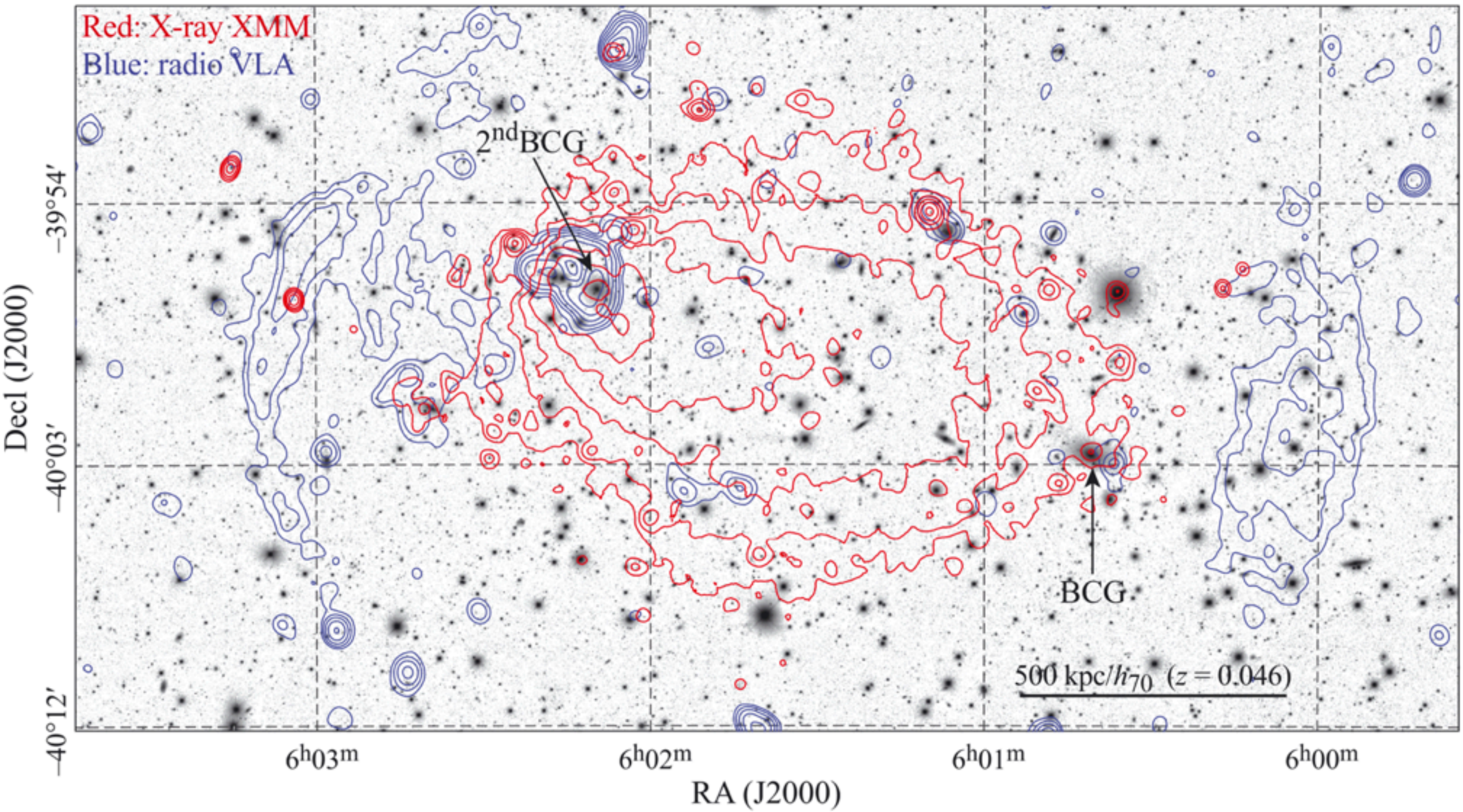
levels of physical modeling



level	approach	pros	cons
0	collisionless N-body	inexpensive, mature methods, DM phenom.	limited physical scope
1	coupled N-body + simple gas dynamics	+ hot gas phenom., good hydro methods	no galaxy phenom., limited applicability
2	+ star/BH formation + SN/AGN feedback	+ galaxy phenom., wide applicability	expensive, large modeling uncertainties

some clusters are more dynamic, like Abell 3376

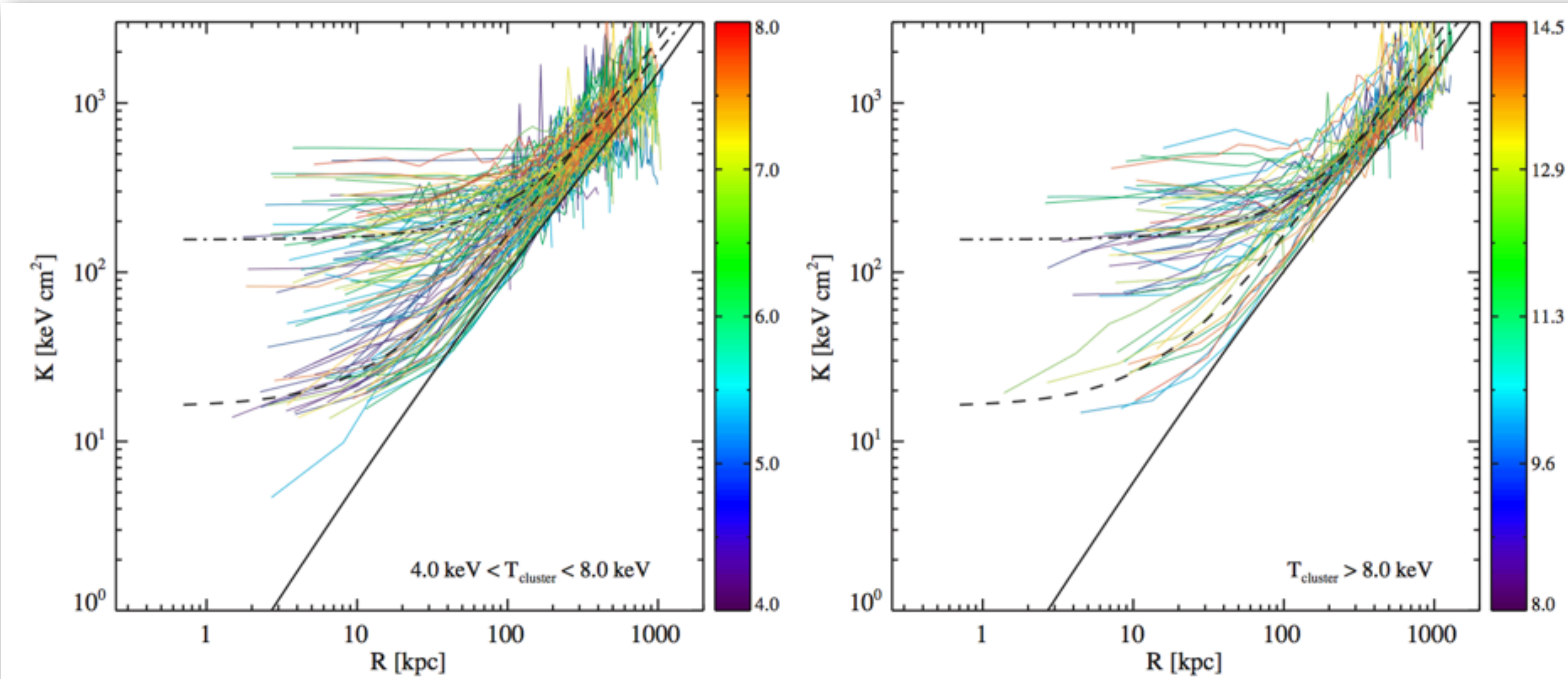
Durret et al 2013



wide range of core entropies signals complex core astrophysics

Cavagnolo et al. 2009

archival Chandra data of ACCEPT cluster sample



low entropy core clusters have BCG's that

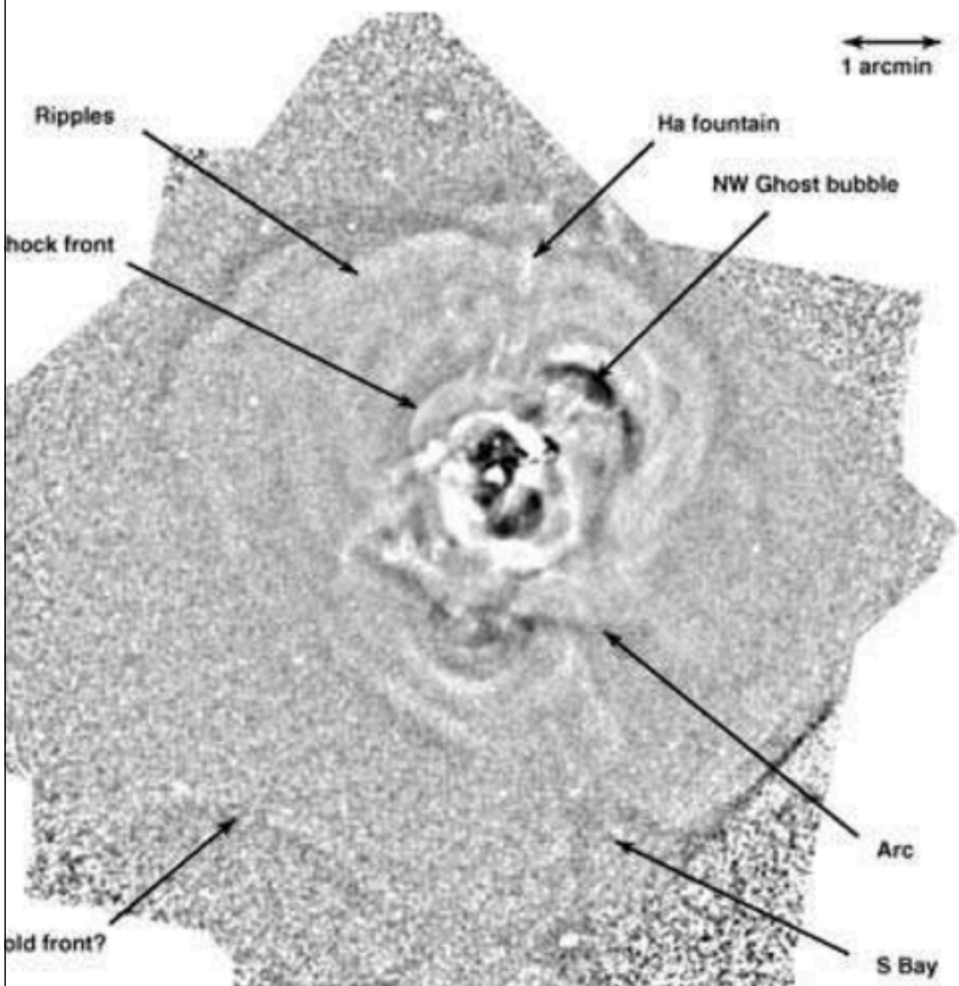
- are centered closely with X-ray peak,
- show more evidence of recent star formation

Hoffer et al. 2012

the central galaxies of clusters are sometimes quite active

Fabian et al. 2005

~Msec Chandra observation of Perseus cluster core



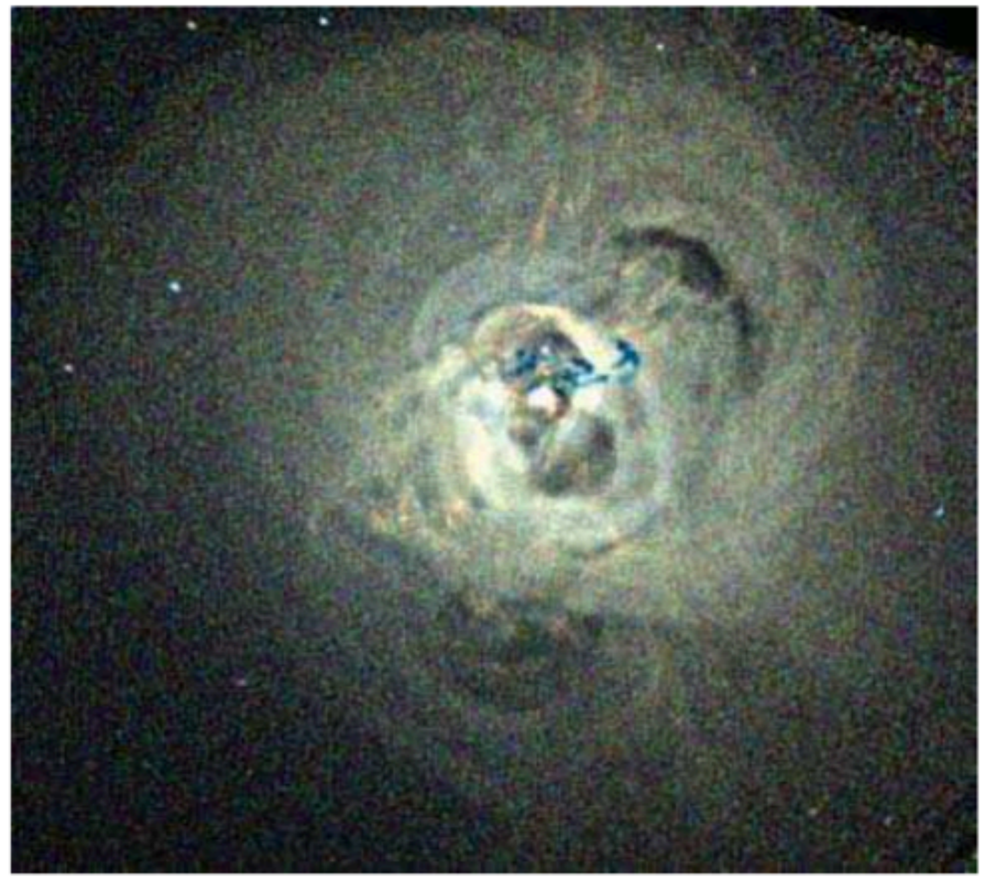
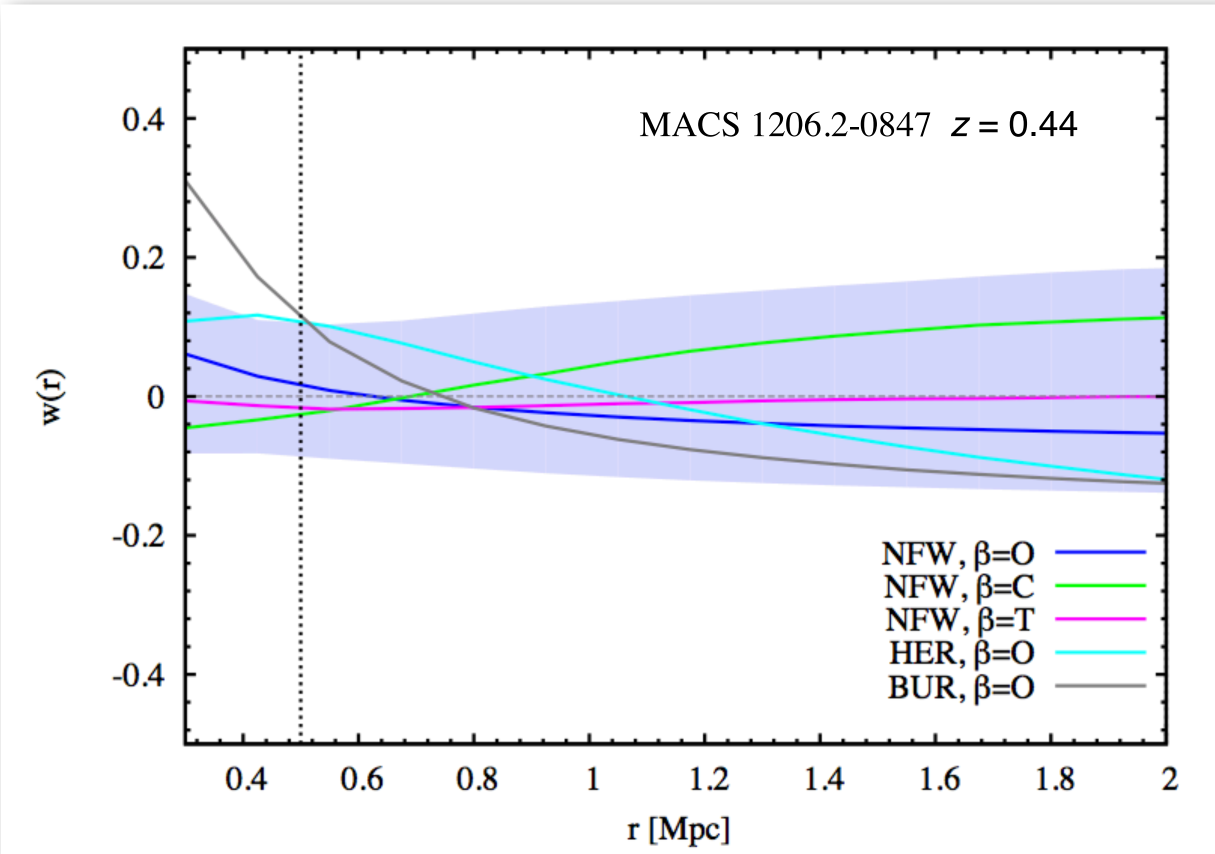


Figure 3. Colour image made from the 0.3–1.2 (red), 1.2–2 (green) and 2–7 keV (blue) bands. A 10-arcsec smoothed image has been scaled to 80 per cent of its intensity and then subtracted in order to bring out fainter features lost in the high-intensity range of raw images. The blue structure to the N of the nucleus is caused by absorption in the infalling high-velocity system,

massive halos can probe fundamental physics

Sartoris et al. 2014

comparing lensing mass to kinematic mass measures
equation of state of the dark matter fluid



$$w(r) = \frac{p_r(r) + 2\,p_t(r)}{3\,c^2\rho(r)}$$

consistent with "dust-like" expectation for weakly interacting dark matter (w=0)

the N-body era

humble beginnings in pre-internet times

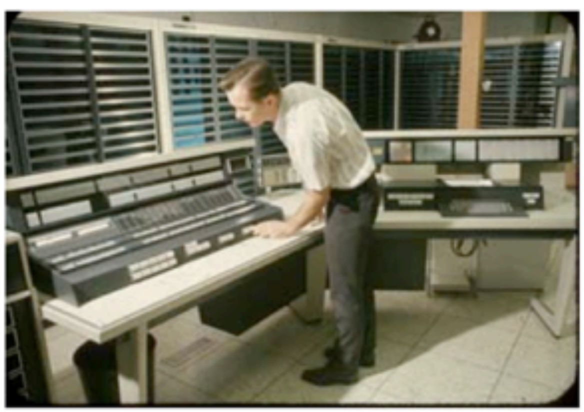
THE ASTRONOMICAL JOURNAL VOLUME 75, NUMBER 1

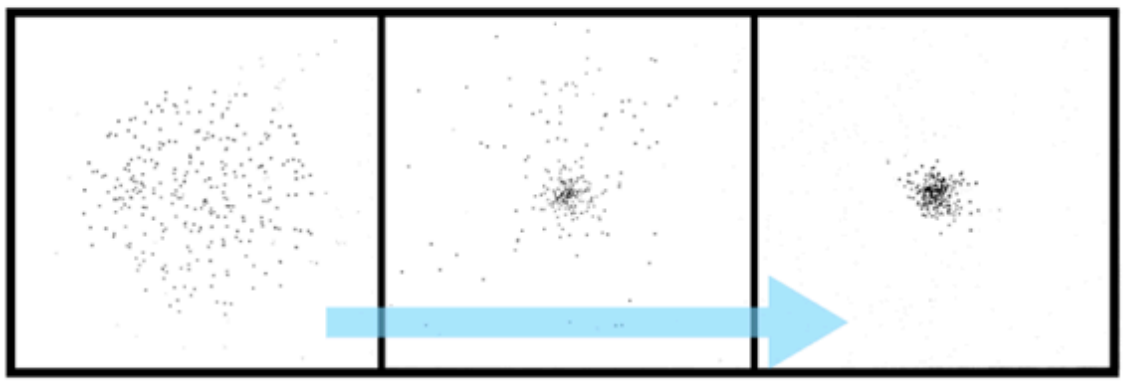
Structure of the Coma Cluster of Galaxies*

P. J. E. Peebles†
Palmer Physical Laboratory, Princeton University, Princeton, New Jersey
(Received 7 October 1969)

In some cosmologies, a cluster of galaxies is imagined to be a gravitationally bound system which, in analogy with the formation of the Galaxy, originated as a collapsing protocluster. It is shown that a numerical model based on this picture is consistent with the observed features of the Coma Cluster of galaxies. The cluster mass derived from this model agrees with previous values; however, an analysis of the observational uncertainty within the framework of the model shows that the derived mass could be consistent with the estimated total mass provided by the galaxies in the cluster.

- **Suite of 300 (and less) particle simulations**
- **Run on a CDC 3600, ~1Mflops, 32KB+ at LANL**
- **Is nine orders of magnitude improvement in both performance and memory good enough for precision cosmology?**





“The Universe is far too complicated a structure to be studied deductively, starting from initial conditions and solving the equations of motion.”

Robert Dicke (Jayne Lectures, 1969)

(courtesy S. Habib, 2013 APS mtg.)

cosmological simulations: algorithmic evolution

1960's+70's - direct (NxN) force summation

studies of galaxy encounters and stellar clusters

1980's - particle-mesh (FFT's) and Tree algorithms for large-scale gravity

studies of `cosmic web' topology from initial random noise field

1990's - parallelization on Beowulf clusters, special purpose chips (GRAPE)

detailed studies of clustering statistics, cosmological dependence

- first multi-fluid codes to model coupled dark matter and baryons

initial studies of galaxy formation

>2000's - massive parallelization on large-scale supercomputers

precise calibration of large-scale structure statistics

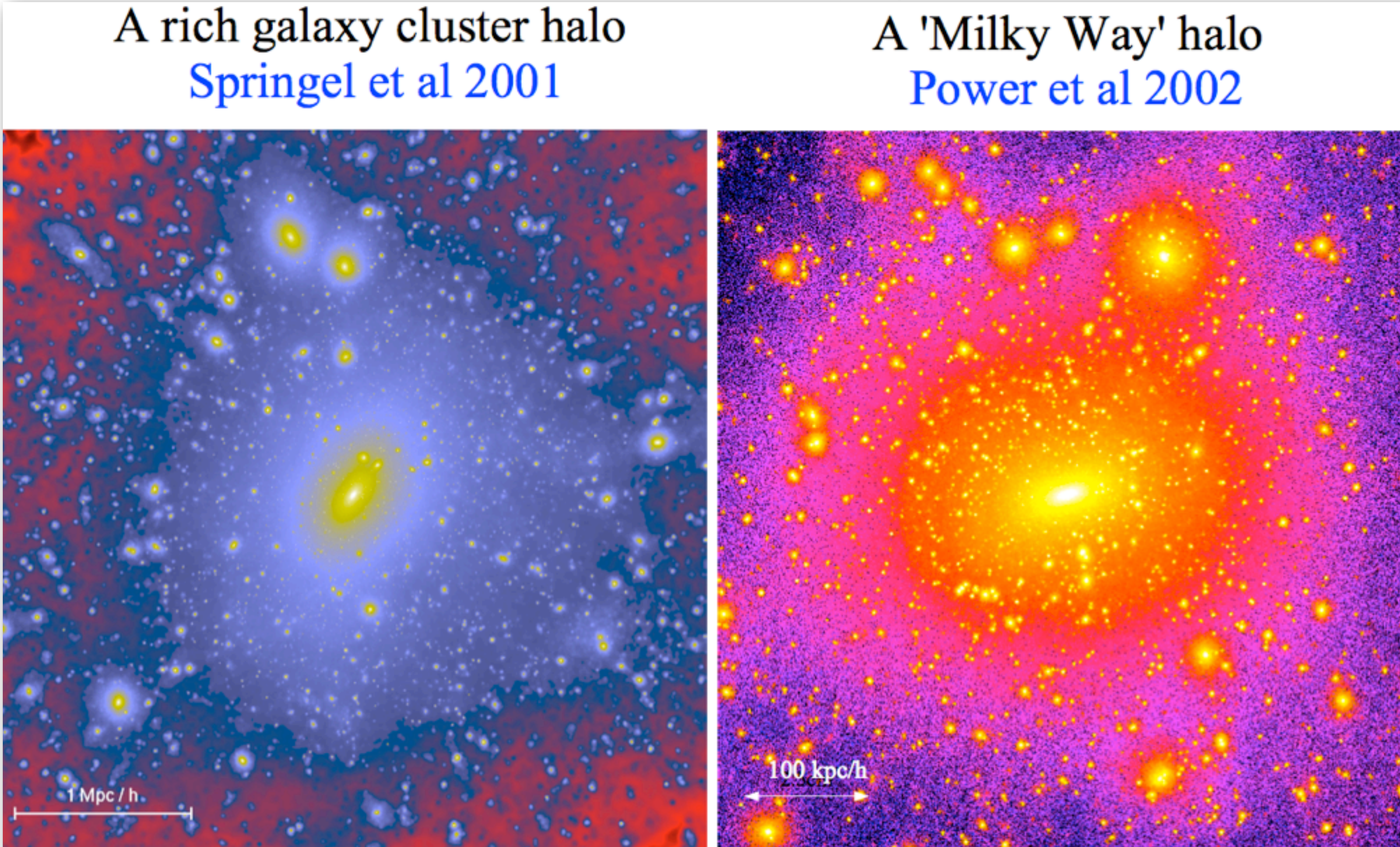
- multi-fluid codes with hydro + approx. radiation transfer, MHD

SMBH formation, feedback effects, high-rez galaxy formation, ...

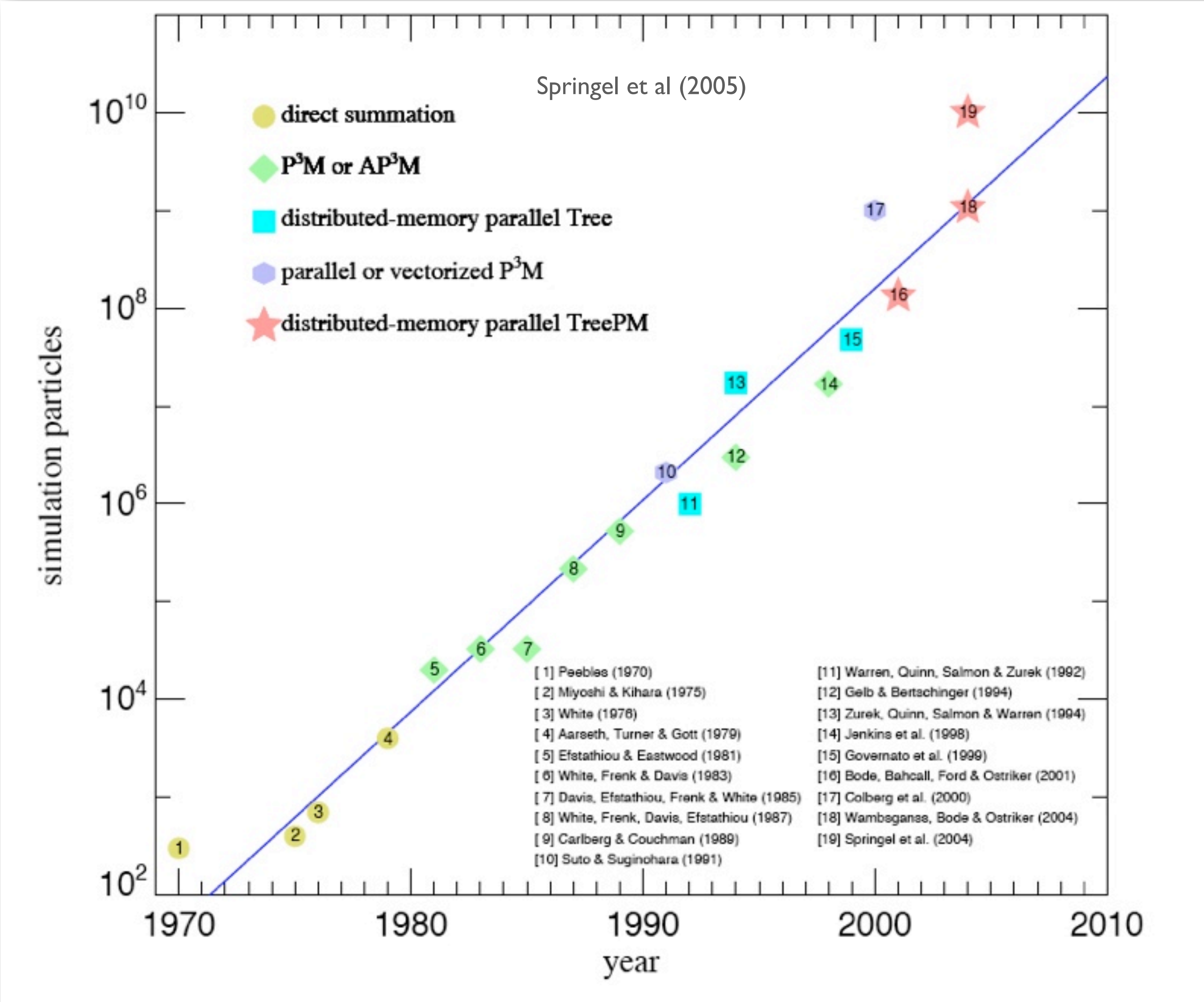
http://www.mpa-garching.mpg.de/aquarius/

similarity of internal halo structure, from dwarf galaxy to cluster scales

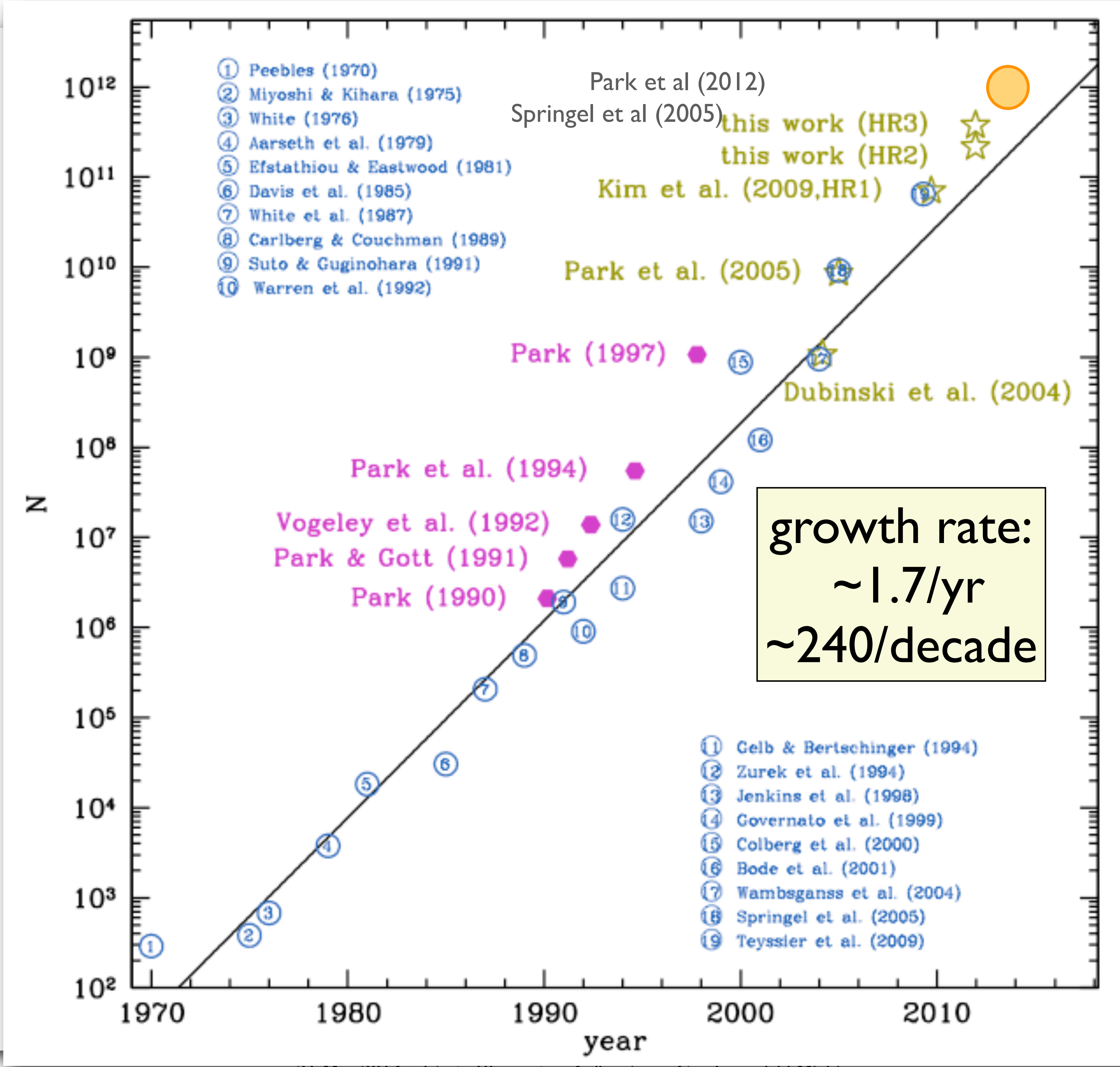
$M_{\text{cluster}} \sim 1000\ M_{\text{galaxy}}$



courtesy S.D.M. White, CATB2009



Springel et al (2005)
simulation particles
year
10^{2}
10^{4}
10^{6}
10^{8}
10^{10}
1970
1980
1990
2000
2010
direct summation
P³M or AP³M
distributed-memory parallel Tree
parallel or vectorized P³M
distributed-memory parallel TreePM
[1] Peebles (1970)
[2] Miyoshi & Kihara (1975)
[3] White (1976)
[4] Aarseth, Turner & Gott (1979)
[5] Efstathiou & Eastwood (1981)
[6] White, Frenk & Davis (1983)
[7] Davis, Efstathiou, Frenk & White (1985)
[8] White, Frenk, Davis, Efstathiou (1987)
[9] Carlberg & Couchman (1989)
[10] Suto & Suginohara (1991)
[11] Warren, Quinn, Salmon & Zurek (1992)
[12] Gelb & Bertschinger (1994)
[13] Zurek, Quinn, Salmon & Warren (1994)
[14] Jenkins et al. (1998)
[15] Governato et al. (1999)
[16] Bode, Bahcall, Ford & Ostriker (2001)
[17] Colberg et al. (2000)
[18] Wambsganss, Bode & Ostriker (2004)
[19] Springel et al. (2004)



Peebles (1970)
Miyoshi & Kihara (1975)
White (1976)
Aarseth et al. (1979)
Efstathiou & Eastwood (1981)
Davis et al. (1985)
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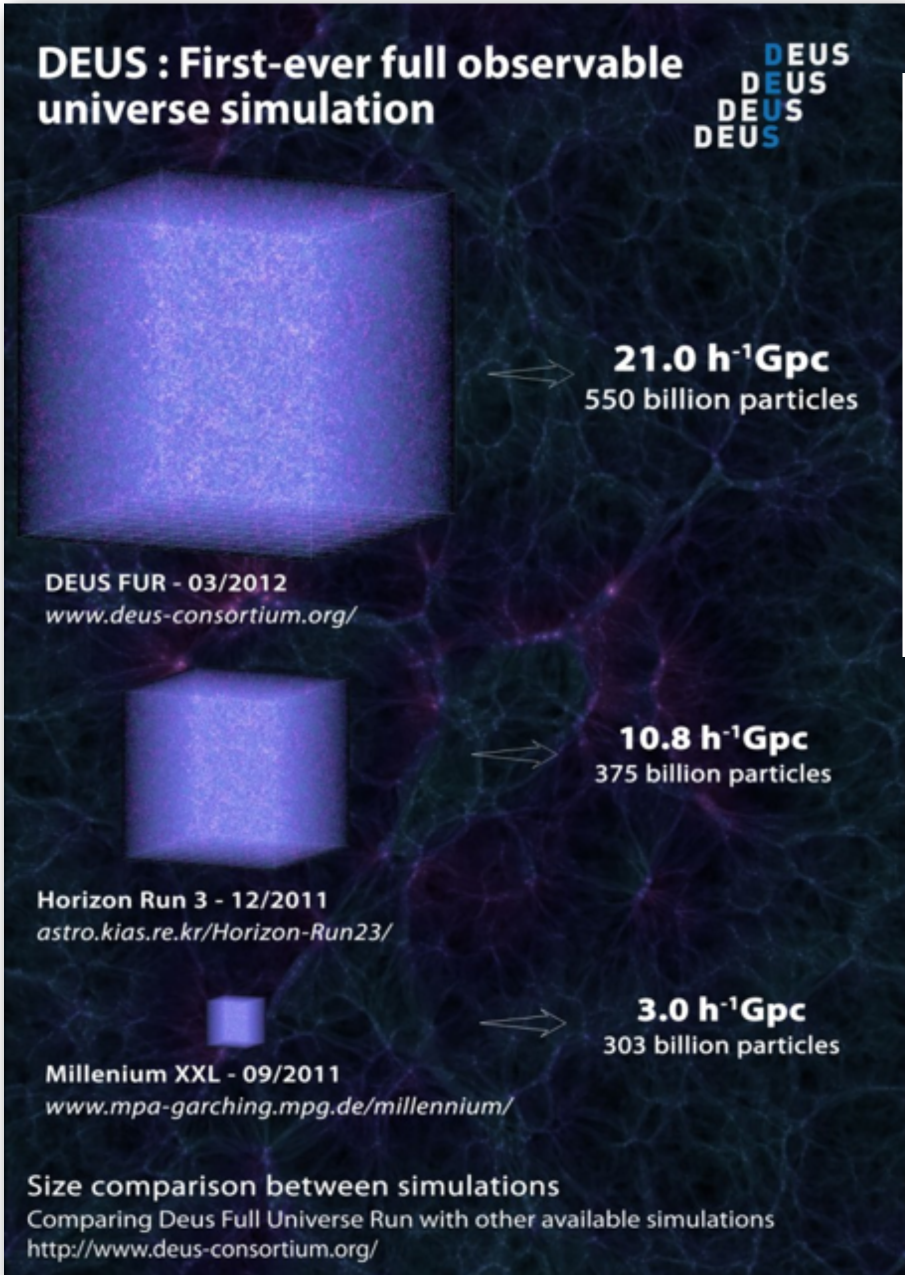
Carlberg & Couchman (1989)
Suto & Suginohara (1991)
Warren et al. (1992)
Park et al (2012)
Springel et al (2005)
this work (HR3)
this work (HR2)
Kim et al. (2009,HR1)
Park et al. (2005)
Park (1997)
Dubinski et al. (2004)
Park et al. (1994)
Vogeley et al. (1992)
Park & Gott (1991)
Park (1990)
growth rate:
~1.7/yr
~240/decade
Gelb & Bertschinger (1994)
Zurek et al. (1994)
Jenkins et al. (1998)
Governato et al. (1999)
Colberg et al. (2000)
Bode et al. (2001)
Wambsganss et al. (2004)
Springel et al. (2005)
Teyssier et al. (2009)
N
year

DEUS simulation effort by Obs. de Paris Meudon group (~0.5 teraparticle)



http://www.deus-consortium.org/

DEUS FUR Characteristics

DEUS FUR is currently the largest and most performing dark matter simulation of the entire cosmos ever realized probing scales from 40 kpc/h to 21 Gpc/h for three dar energy models; the concordance ΛCDM model a quintessence model and a fantom model.

DEUS FUR is an acronym for **« Dark Energy Universe Simulation – Full Universe Run ».**

The simulation has followed the self-gravitational evolution of **8192^3** (**549 755 813 888**) particles in cubic volume of (**21 Gpc)^3** .

This simulation has required **5** million cpu hour on **76032** cores of the Curie Supercomputer at TGCC

It has used **304 128** Go of memory.

150 Po data have been generated during the run. Using an optimized chain of post-processing programs we have been able to reduce these data to about **1.5 Po**.

Energy (power+cooling) cost:

5M core-hours x ~50 W/core
= 250,000 kW-hours
x $0.1/kW-hour
=> $25,000 x (PUE) ~ $50k

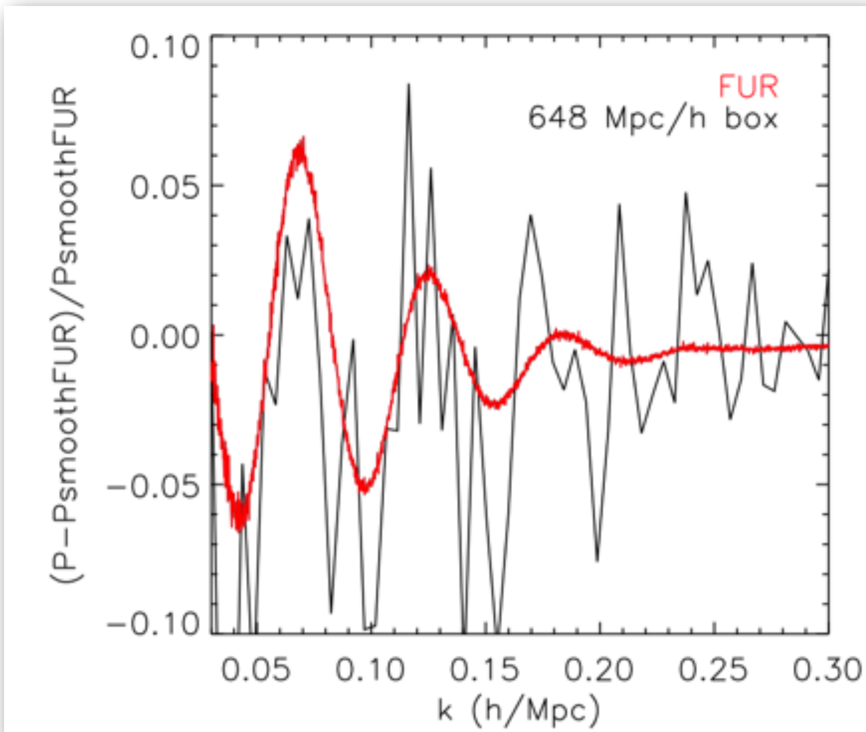
~ $10k / Mcpu-hr

DEUS-FUR detailed study of BAO spectral features

http://www.deus-consortium.org/

Cosmic variance limited Baryon Acoustic Oscillations from the DEUS-FUR ΛCDM simulation

Y. Rasera*, P.-S. Corasaniti, J.-M. Alimi, V. Bouillot, V. Reverdy, and I. Balmes
CNRS, Laboratoire Univers et Théories (LUTh), UMR 8102 CNRS, Observatoire de Paris, Université Paris Diderot ; 5 Place Jules Janssen, 92190 Meudon, France



DEUS-FUR detailed study of BAO spectral features

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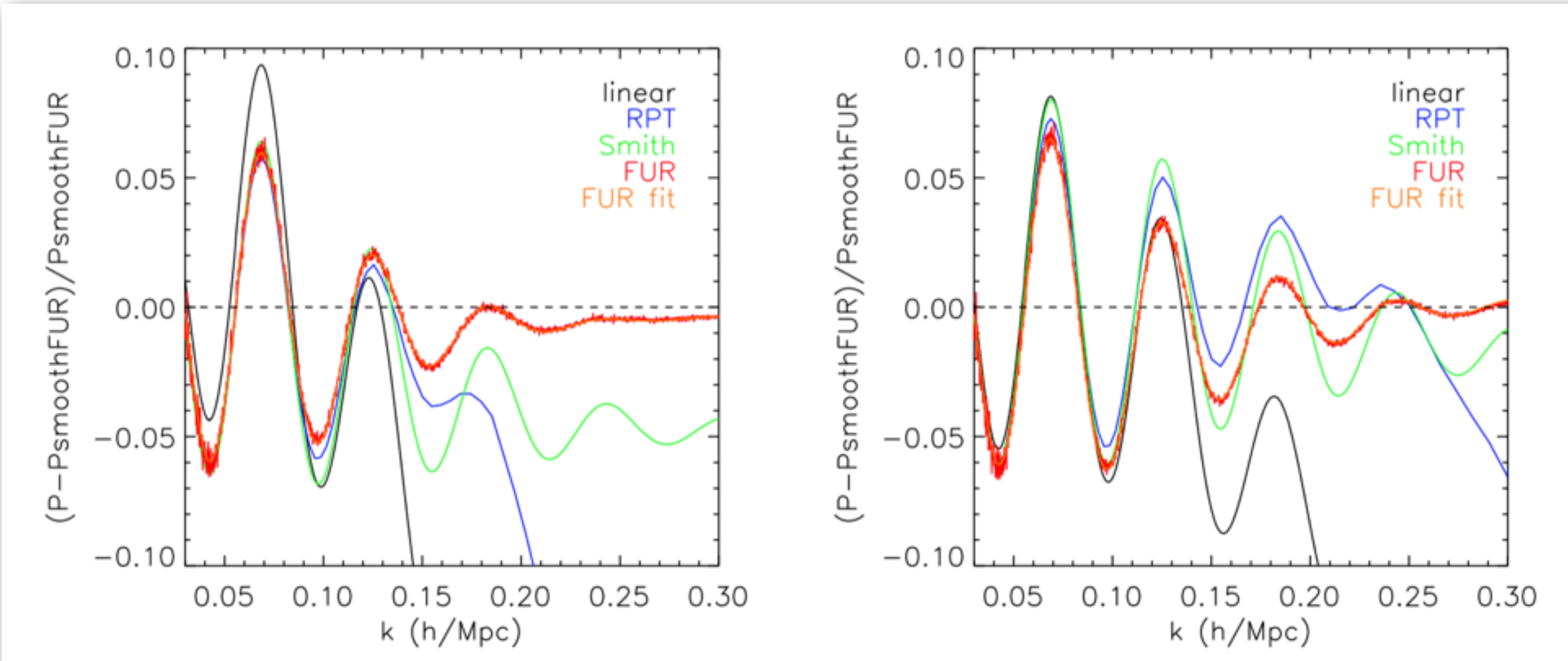
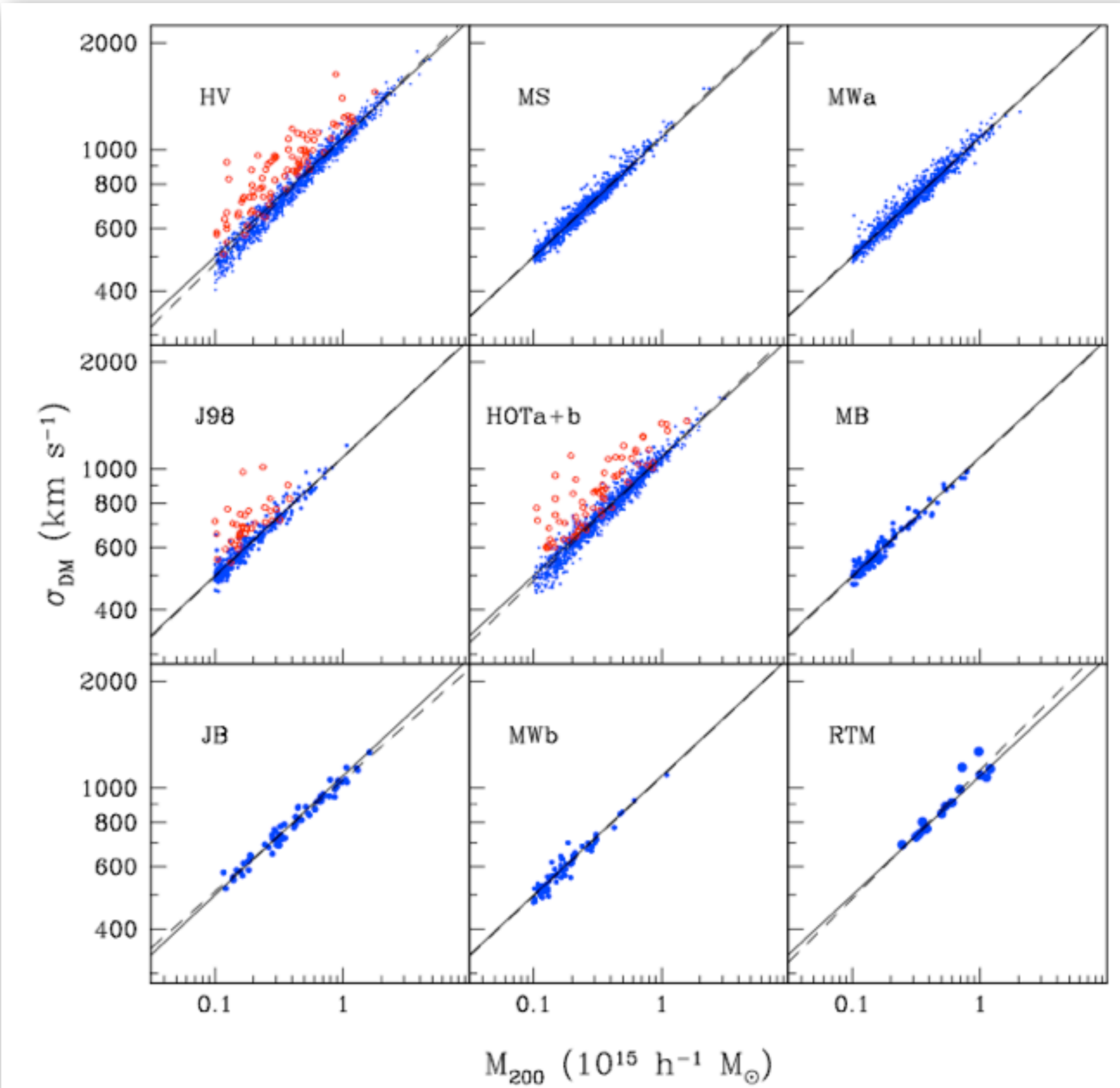


multi-code, precise calibration of DM halo virial scaling relation

Evrard et al (2008)

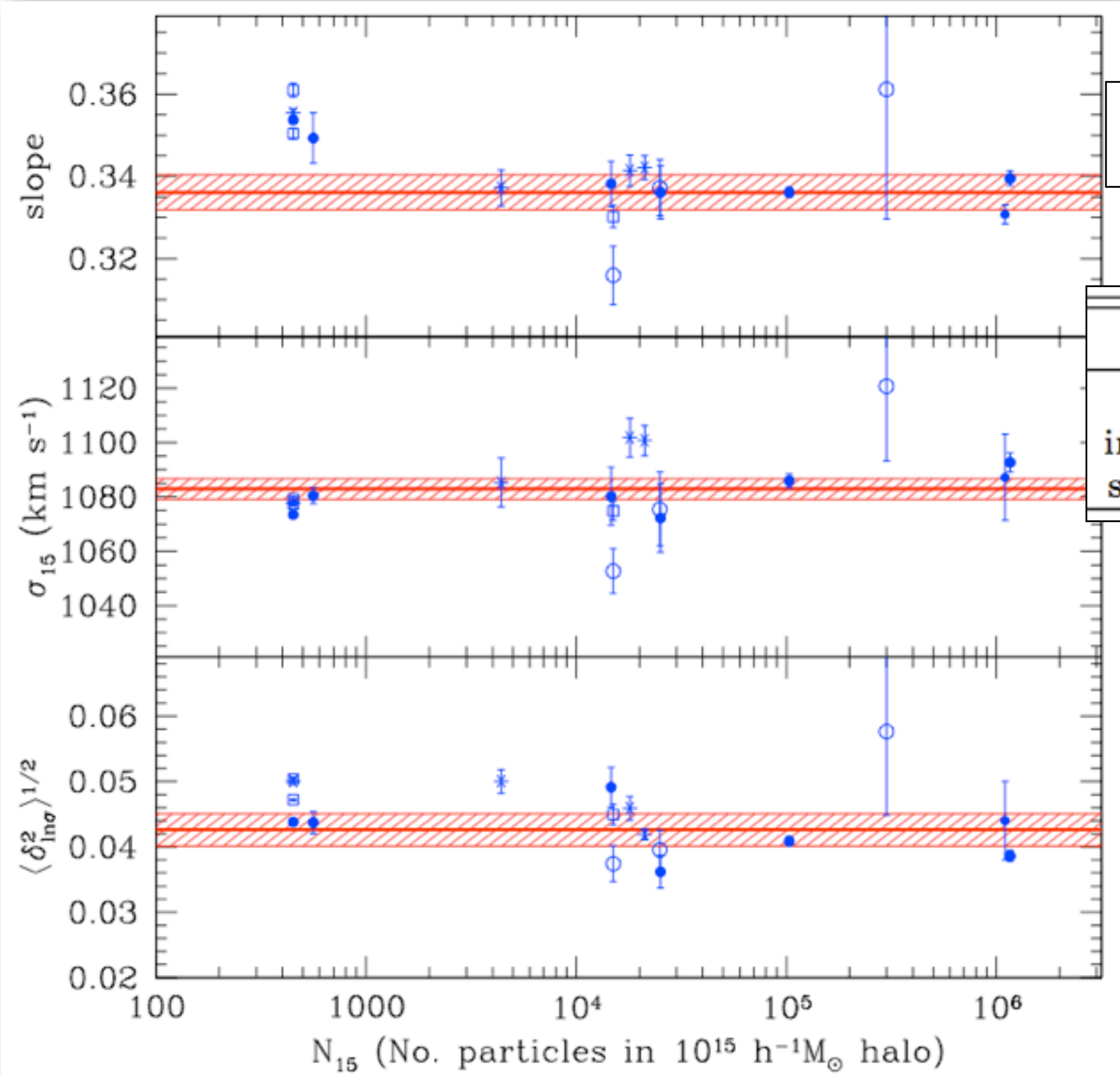


$$\sigma_{\rm DM}{}^2 = \frac{1}{3N_p}\sum_{i=1}^{N_p}\sum_{j=1}^{3}|v_{i,j}-\bar{v_j}|^2$$

* results from six different N-body codes

* evidence that mergers are dynamically soft (in the bulk)

multi-code, precise calibration of DM halo virial scaling relation



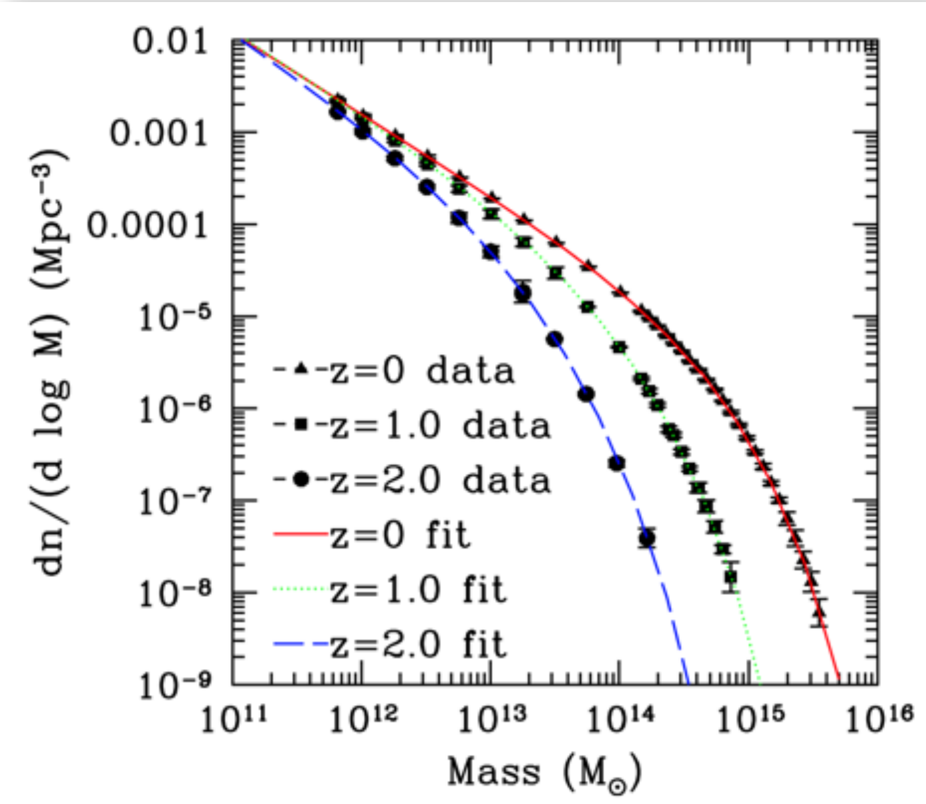
Evrard et al (2008)

$$\sigma_{\rm DM}(M,z) \;=\; \sigma_{\rm DM,15}\left(\frac{h(z)M_{200}}{10^{15}\ {\rm M}_\odot}\right)^{\alpha}$$

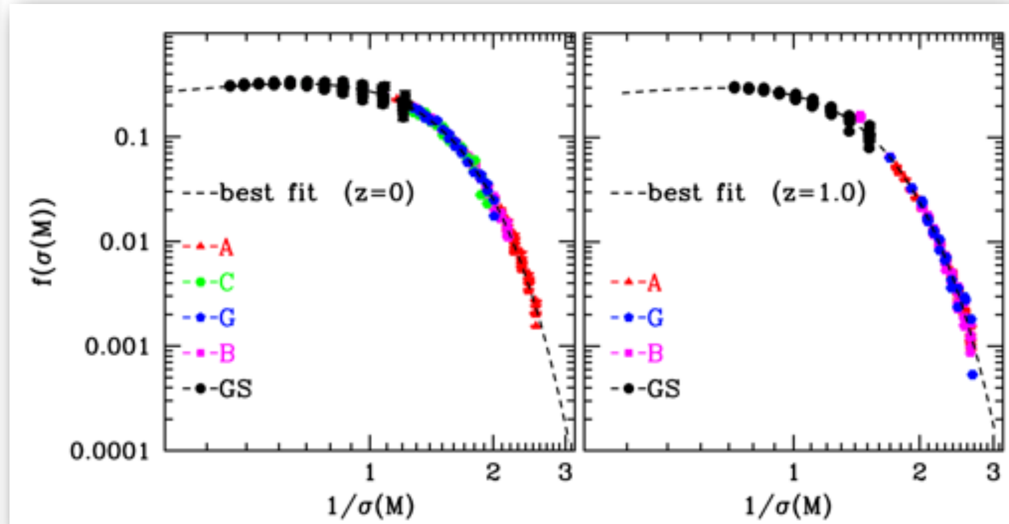
Parameter	Value
slope, α	0.3361 ± 0.0026
intercept, $\sigma_{\rm DM,15}$	1082.9 ± 4.0 km s^{-1}
scatter, $\langle\delta^2_{\ln\sigma}\rangle^{1/2}$	0.0426 ± 0.0015

halo space density calibration in *w*CDM cosmologies

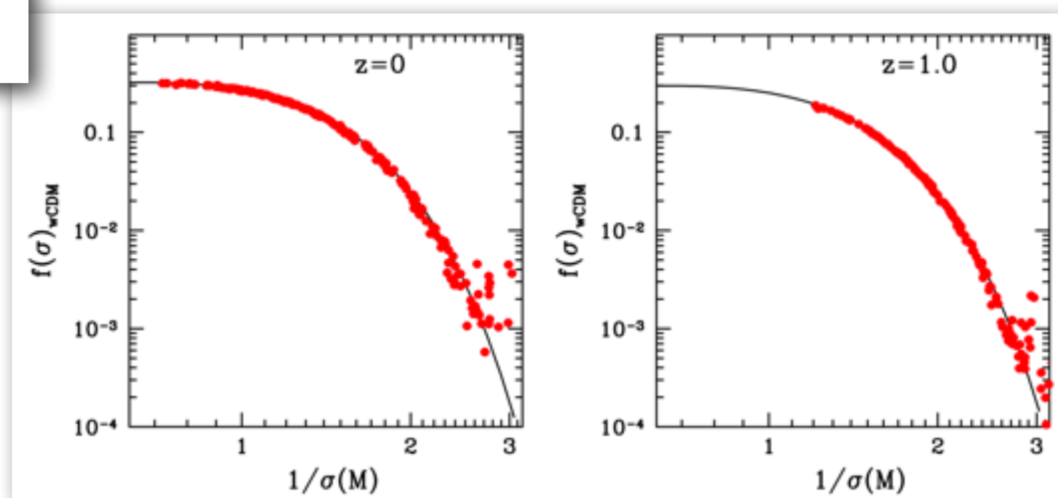
Bhattacharya et al (2011)



~2% rms fit for ΛCDM



~10% for *w*CDM



universe with ***only***
collisionless dark matter

N-body simulations inform an increasingly precise **halo model** of LSS

* general aspects
  – halos are dynamically evolving systems: close to virial equilibrium but frustrated by mergers and continual accretion
  – ellipsoidal in shape (tending prolate) with 2:1 axis ratios common aligned with surrounding filaments

* internal structure
  – relaxation to similar density + velocity radial profiles
  – surviving sub-halo structures contain a small percentage of total mass
  – hierarchical nesting of sub-structure families reflect accretion history

* low-order spatial distribution
  – functional forms for *mass function*, n(M,z), and *bias function*, b(M,z), precisely calibrated via similarity variable, σ(M) (mainly *w*CDM)
  – different, one-parameter mass assignment methods (FOF, SO) exist
    good: flexibility, reflects edge complexity **bad**: literature confusing

multi-fluid models:
N-body + baryons
(+ chemistry + radiation + B-fields + ...)

multi-fluid systems: N-body + basic baryonic gas dynamics

Bertschinger 1998

* baryon fluid coupled via gravity to DM
* solve Euler equation in comoving coordinates
* energy or entropy equation
* requires shock treatment

In comoving coordinates, the cosmological fluid equations are

$$\frac{\partial}{\partial t}\left(\frac{\rho_b}{\bar{\rho}_b}\right) + \frac{1}{a}\vec{\nabla}\cdot\vec{v}_b = 0,$$

$$\frac{\partial \vec{v}_b}{\partial t} + \frac{1}{a}\vec{v}_b\cdot\vec{\nabla}\vec{v}_b + H\vec{v}_b = -\frac{1}{a\rho_b}\vec{\nabla}p + \vec{g}, \quad (3)$$

where ρ_b, $\bar{\rho}_b$, $\vec{v}_b$, and p are the (baryonic) mass density, mean mass density, peculiar velocity, and pressure, respectively, and $\vec{g}$ is the gravitational field (Equation 1). These must be supplemented by either an energy or entropy equation. Outside of shocks, these take the form

$$\frac{\partial u}{\partial t} + \frac{1}{a}\vec{v}_b\cdot\vec{\nabla}u = -\frac{p}{a\rho_b}\vec{\nabla}\cdot\vec{v}_b + \frac{1}{\rho_b}(\Gamma - \Lambda),$$

$$\frac{\partial S}{\partial t} + \frac{1}{a}\vec{v}_b\cdot\vec{\nabla}S = \frac{1}{p}(\Gamma - \Lambda). \quad (4)$$

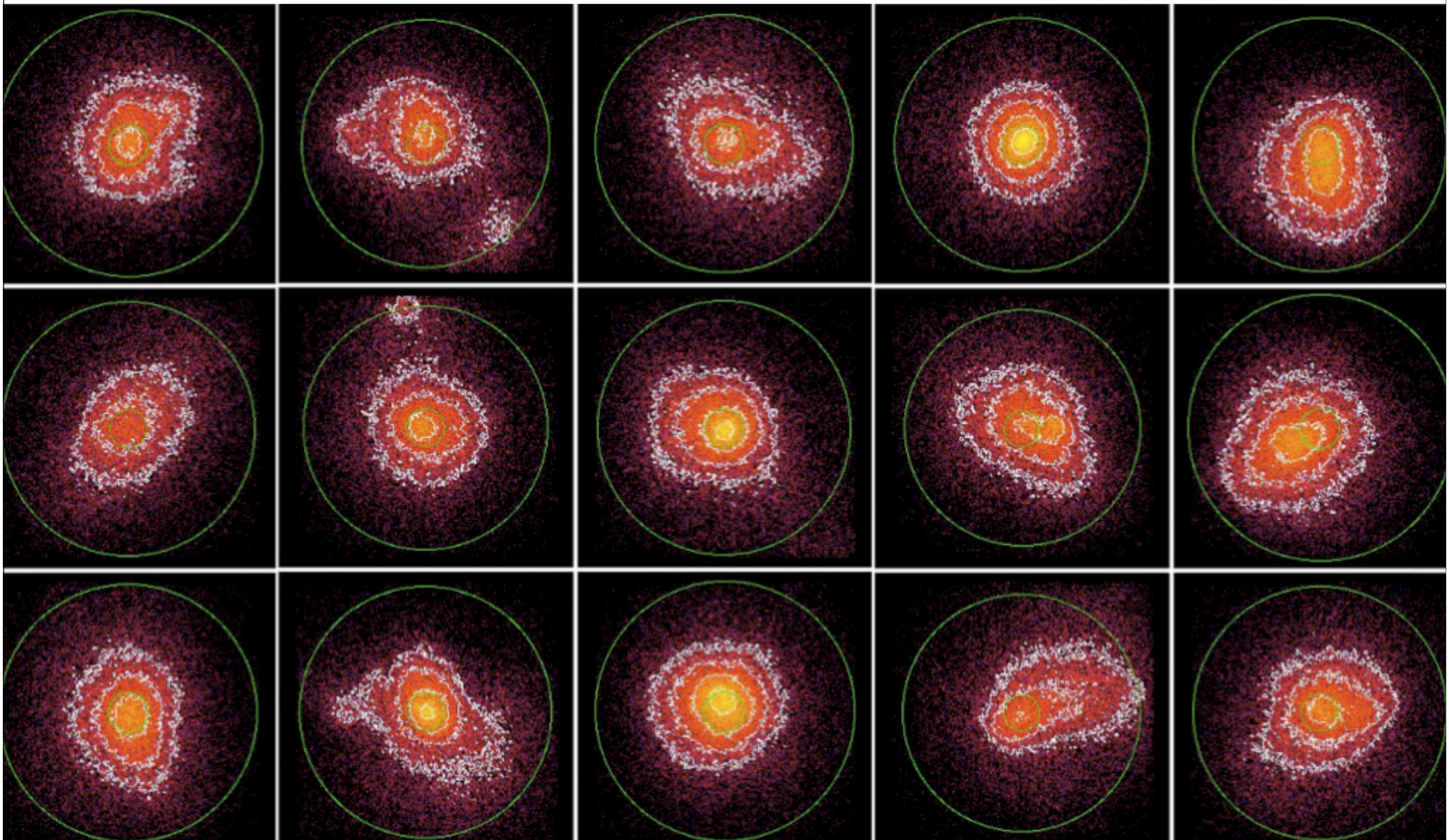
For a perfect gas with ratio of specific heats γ, the thermal energy and entropy per unit mass are $u = p/[(\gamma-1)\rho_b]$ and $S = (\gamma-1)^{-1}\ln(p\rho_b^{-\gamma})$, respectively. Artificial viscosity is often added to Equation 4 to generate the entropy needed across shock waves. In nonadiabatic calculations, heating and cooling rates per unit volume Γ and Λ and all they depend on, such as ionization and chemistry rate equations, radiative transfer, etc, must be included.

hydro solution methods: various flavors

method	character	advantages	disadvantages	examples
Lagrangian (particle)	•solve fluid eq'n along streamlines •local kernel density estimates	•simple, fast •good dynamic range w/ variable kernel scale	•approx. shock treatment •poor error control (no grid)	smoothed particle hydro (SPH) • gadget • gasoline
Eulerian fixed mesh	•uniform (cubic) spatial grid	•simple, fast •good (trunc.) error control •shocks	•limited spatial resolution	• c.f., Kang et al (1994)
Eulerian Adaptive Mesh Refi. (AMR)	•grid cells refined (sub-divided) in target regions	•improved spatial and mass resol'n •wider dynamic range	•memory intensive •complex post-processing	• ART • Enzo • RAMSES • FLASH
Moving Mesh	•hybrid Lagr./Eul. •deformable, moveable grid cells (up to max.)	•best of breed?	•complex to code	•Arepo

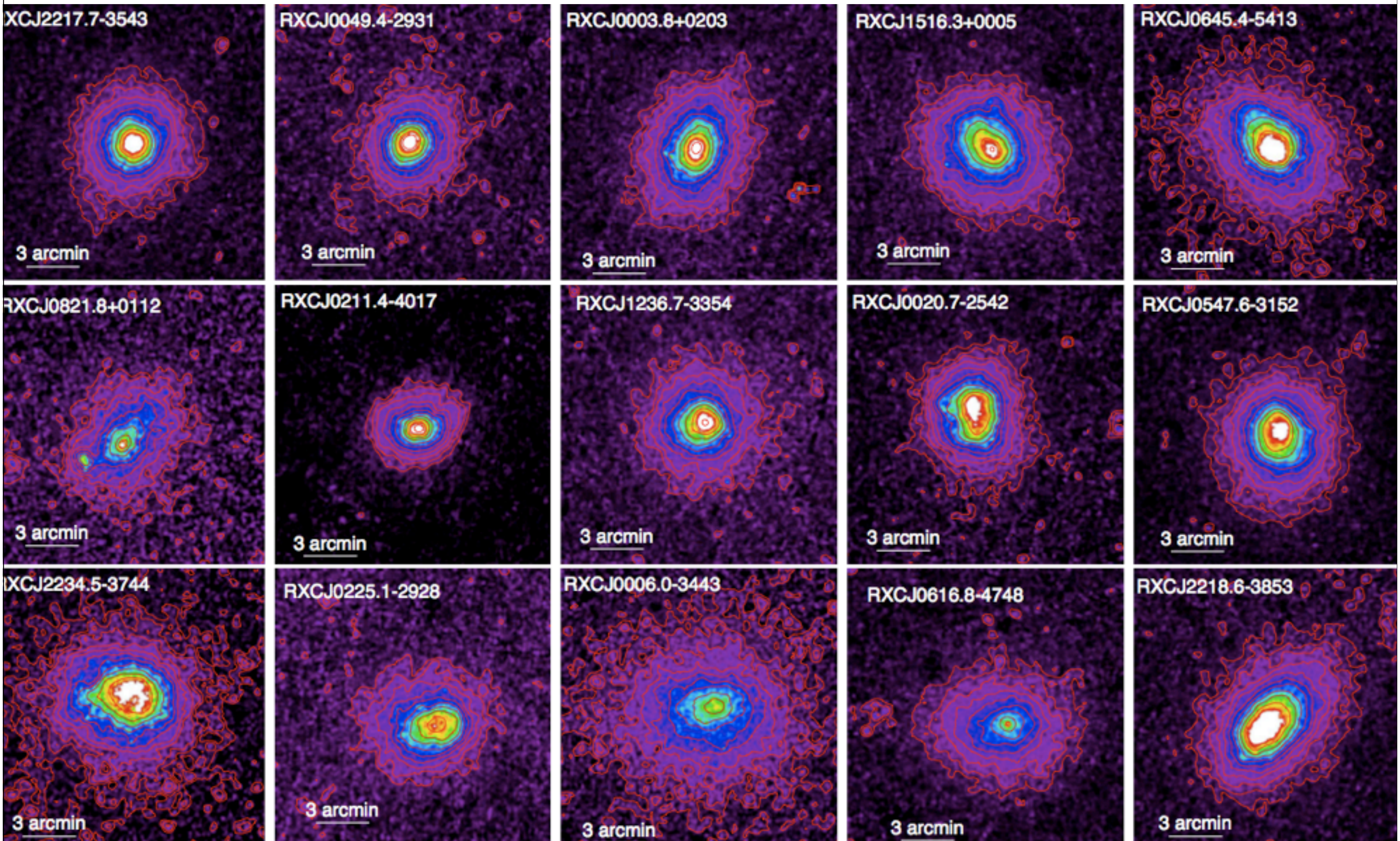
synthetic XMM images of Millennium Gas simulations (Gadget SPH)

courtesy E. Rasia



XMM images of REXCESS cluster sample

Pratt et al. (2009)



`designer' simulation: SPH 2-body model of X-ray emission from A3376

Machado & Lima Neto 2013

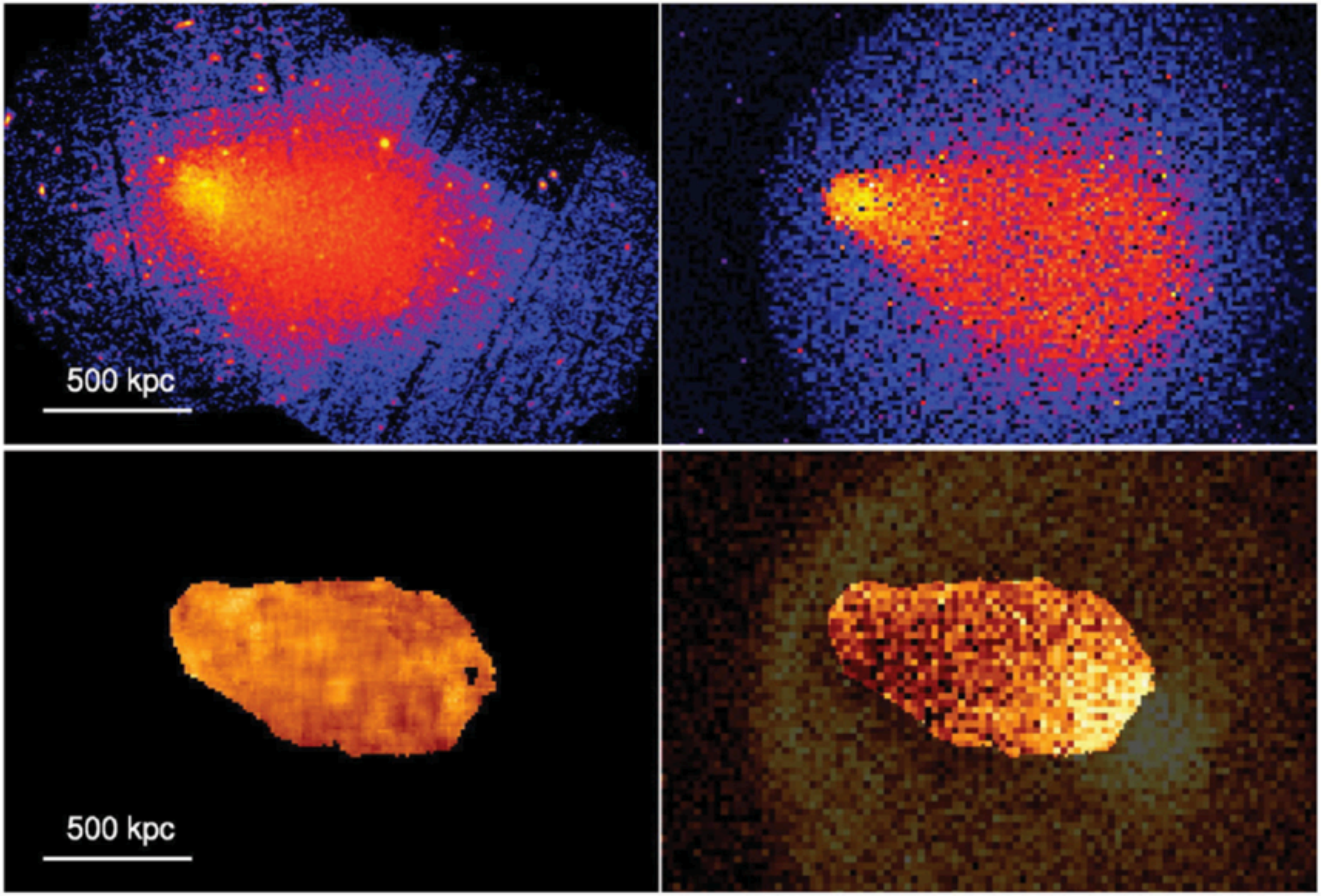
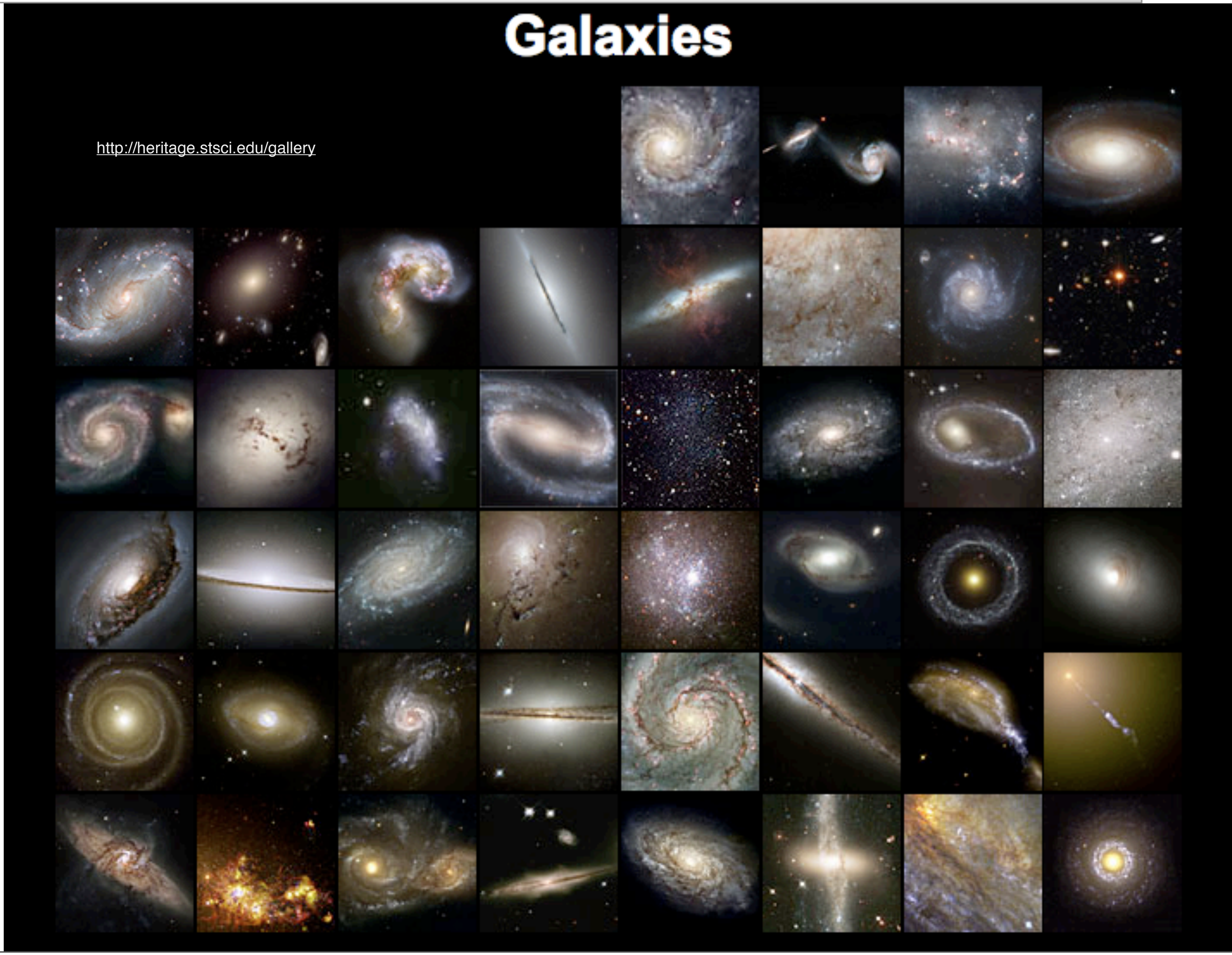
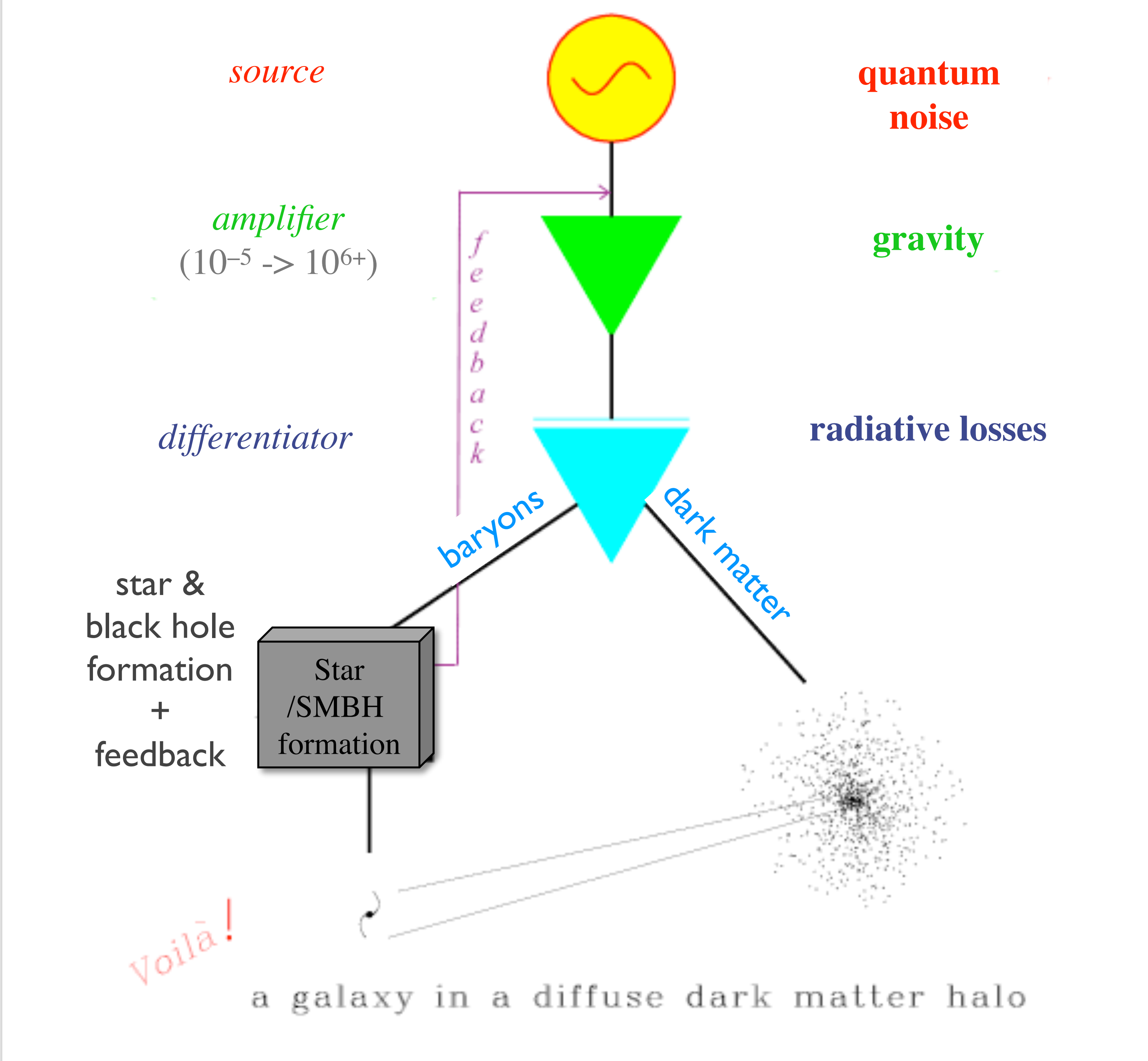


Figure 3. Comparison between observations of A3376 (left) and model 233

challenge: halos (and sub-halos) should contain baryonic objects like this!



cosmic engineering: a flowchart for galaxy formation



early galaxy formation results with P3MSPH

- 16 Mpc cube in Ω_m=1 universe (aka, SCDM)
- 2 x 64^3 particles on CRAY Y-MP (@SDSC)
- DM $m_p \approx$1e9 Msun, baryon $m_p \approx$1e8 Msun, soft≈10kpc
- shock heating + radiative cooling only

first cosmological simulation
to naturally form disk galaxies

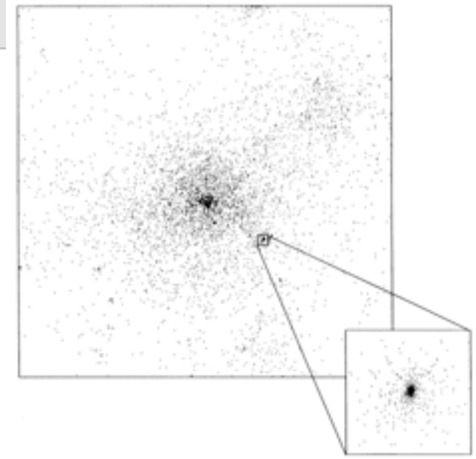


Fig. 17a

Evrard, Summers & Davis 1994

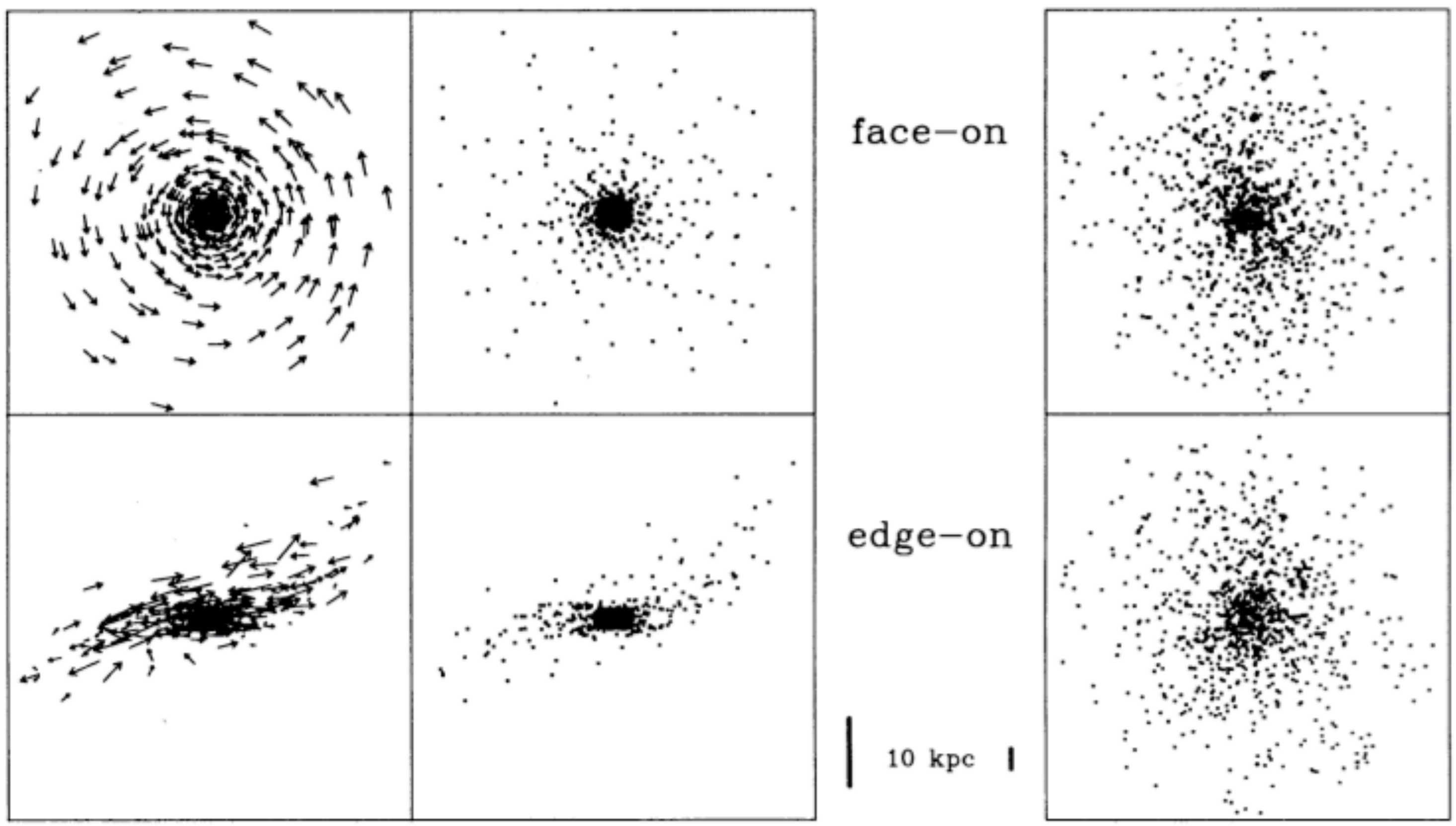


Fig. 17b

early results on galaxy formation with P3MSPH

Evrard, Summers & Davis 1994

first estimate of the galaxy Halo Occupation Distribution (HOD)

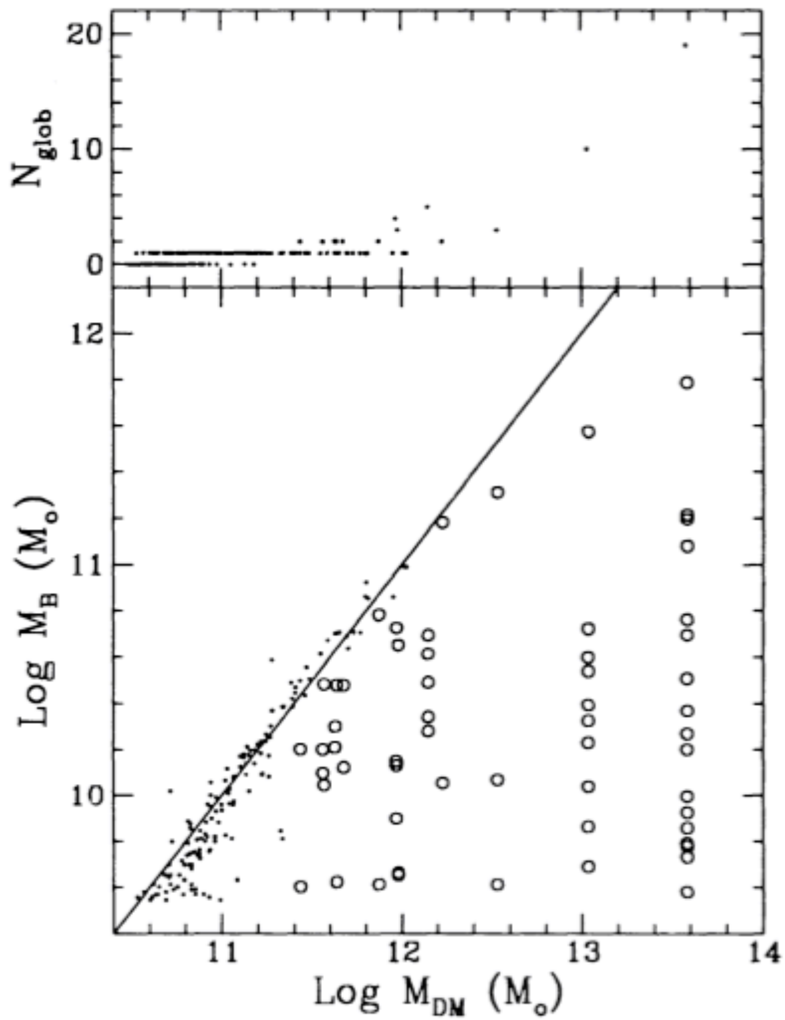


FIG. 11.—Halo occupation number N_{glob} and glob mass within each halo as a function of halo mass. Circles in the lower panel indicate halos containing multiple globs. The line in the lower panel is $M_B = \Omega_b M_{DM}$.

\`GLOBs'
= Galaxy-Like OBjects

baryon physics in current codes is now much more sophisticated

Benson (2010)

Table 1: A survey of physical processes included in several of the major hydrodynamical codes. The primary reference is indicated next to the name of the code. Where implementations of major physical processes are described elsewhere the reference is given next to the entry in the relevant row.

Feature	GADGET-3[1]	GASOLINE[2]	HART[3]	ENZO(ZEUS)[4]	FLASH[5]
Gravity	Tree	Tree	AMR[6] PM[7]	AMR[6] PM[7]	Multi-grid
Hydrodynamics	SPH[8]	SPH[8]	AMR[6]	AMR[6]	AMR[6]
→ Multiphase subgrid model[9]	✓[10]	×	N/A	N/A	N/A
Radiative Cooling	✓	✓	✓	✓	✓[11]
→ Metal dependent	✓[12]	×	✓[13]	✓[14]	✓[11]
→ Molecular chemistry	✓[15]	×	✓[13] [16]	✓[17]	×
Thermal Conduction	✓[18]	×	×	×	✓
Star formation	✓[19]	✓[20]	✓[13]	✓[21]	×
→ SNe feedback	✓[19]	✓[20]	✓[13]	✓[21]	×
→ Chemical enrichment	✓[19]	✓[20]	✓[13]	✓[21]	×
Black hole formation	✓[22]	×	×	×	✓[23]
→ AGN feedback	✓[22]	×	×	×	×
Radiative transfer	OTVET[24,25]	×	OTVET[24]	✓[26]	✓[27]
Magnetic fields	✓[28]	×	×	✓[29]	✓[30]

Notes

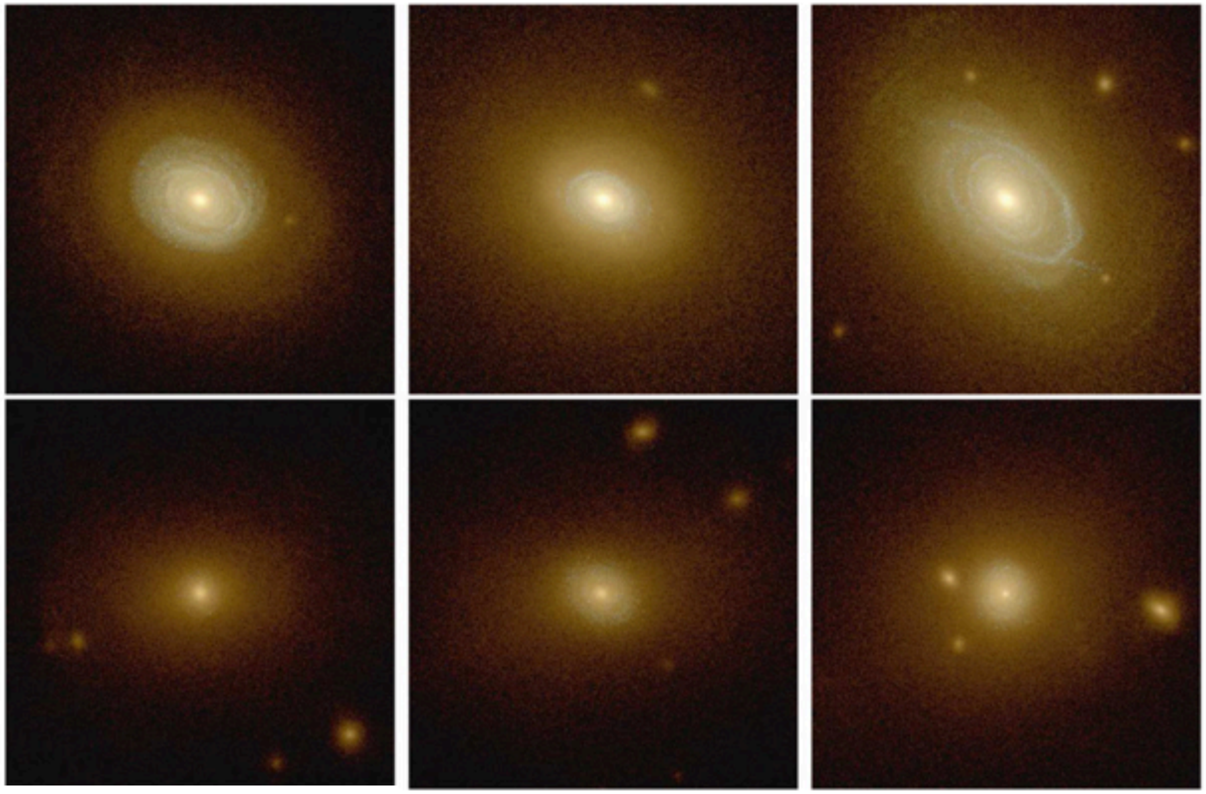
1 "GAlaxies with Dark matter and Gas intEracT" (Springel, 2005);
2 Wadsley et al. (2004);
3 Hydrodynamic Adaptive Refinement Tree (Kravtsov et al., 2002);
4 O'Shea et al. (2004);
5 http://flash.uchicago.edu (Fryxell et al., 2000);
6 Adaptive Mesh Refinement;
7 Particle–mesh;
8 Smoothed Particle Hydrodynamics;
9 Applicable only to SPH codes—used correctly, AMR codes naturally resolve multiphase media;
10 Scannapieco et al. (2006a);
11 Banerjee et al. (2006);
12 Scannapieco et al. (2005);
13 Tassis et al. (2008);
14 Smith et al. (2009);
15 Yoshida et al. (2003);
16 Equilibrium only;
17 Turk (2009);
18 Jubelgas et al. (2004);
19 Scannapieco et al. (2005);
20 Governato et al. (2007);
21 Tasker and Bryan (2008);
22 Matteo et al. (2005);
23 Federrath et al. (2010);
24 Optically Thin Variable Eddington Tensor;
25 Petkova and Springel (2009);
26 Flux-limited diffusion approximation (Norman et al. 2009; see also Wise and Abel 2008b);
27 Rijkhorst et al. (2006); Peters et al. (2010);
28 Dolag and Stasyszyn (2008);
29 Collins et al. (2009; see also Wang and Abel 2009);

RAMSES simulations w/ AGN feedback

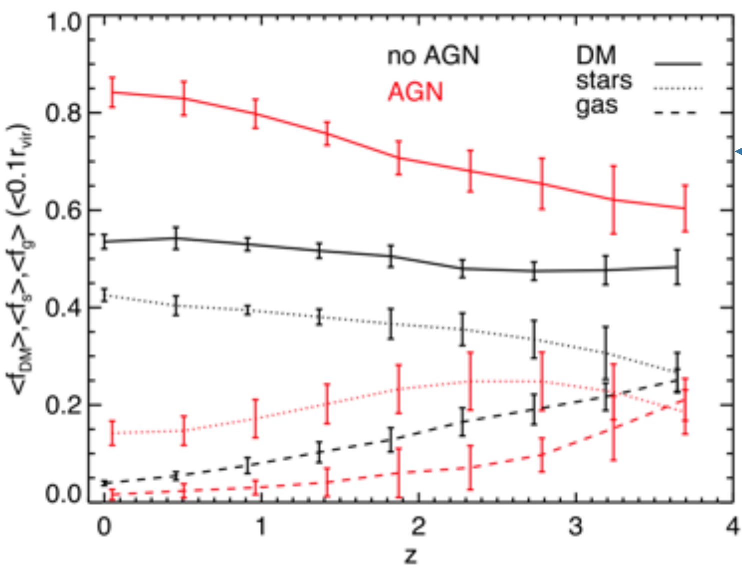
Dubois et al 2013

no AGN →

AGN →

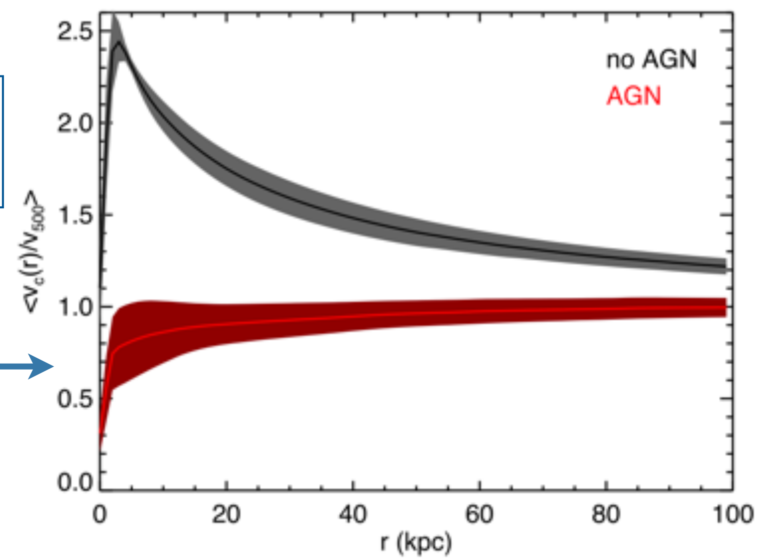


central galaxies in group-scale halos ~10^{13} M_{sun}



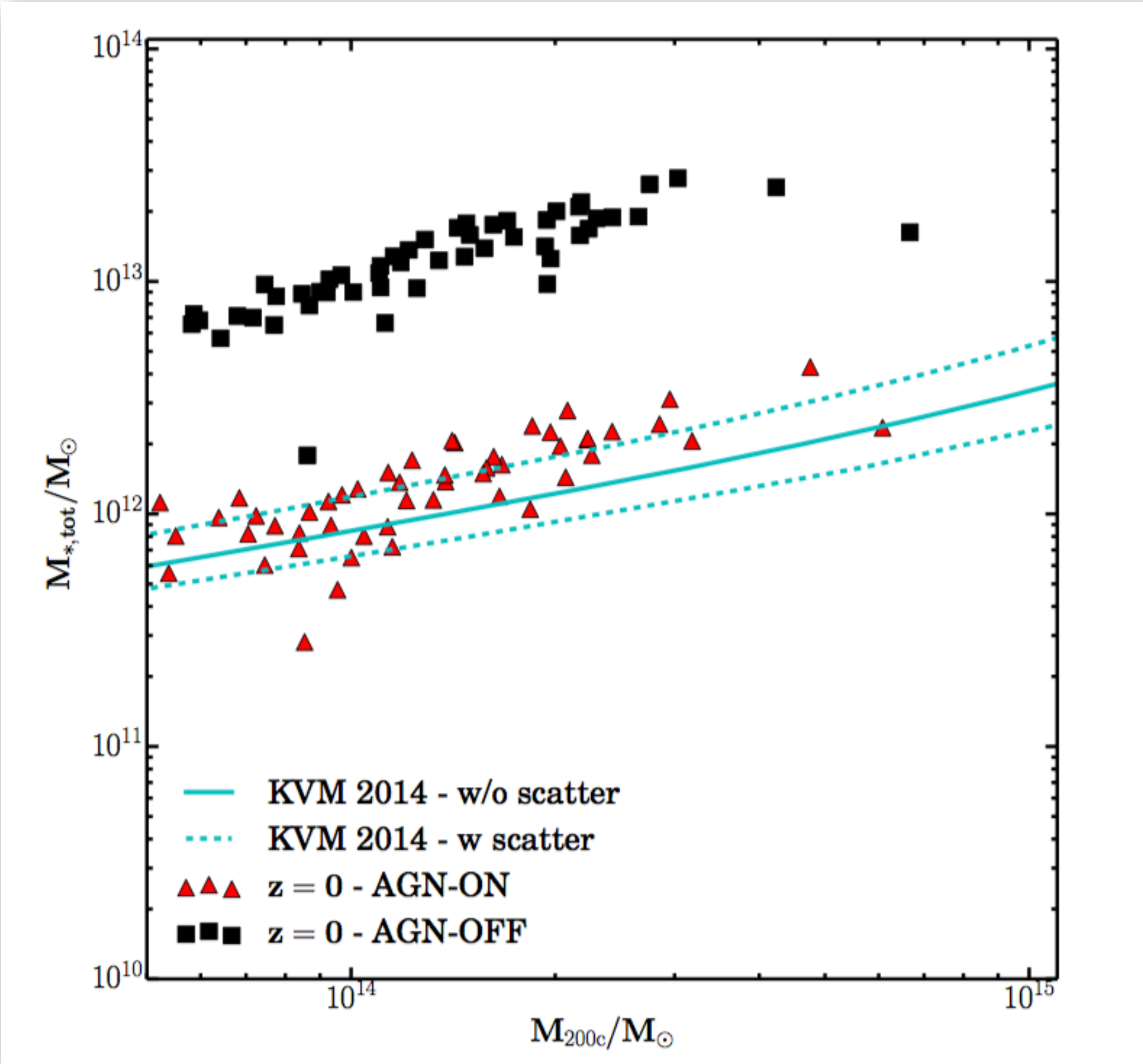
reduces fraction of baryons in stars

reduces core density; better rotation curves



RAMSES simulations w/ AGN feedback

Martizzi et al 2014



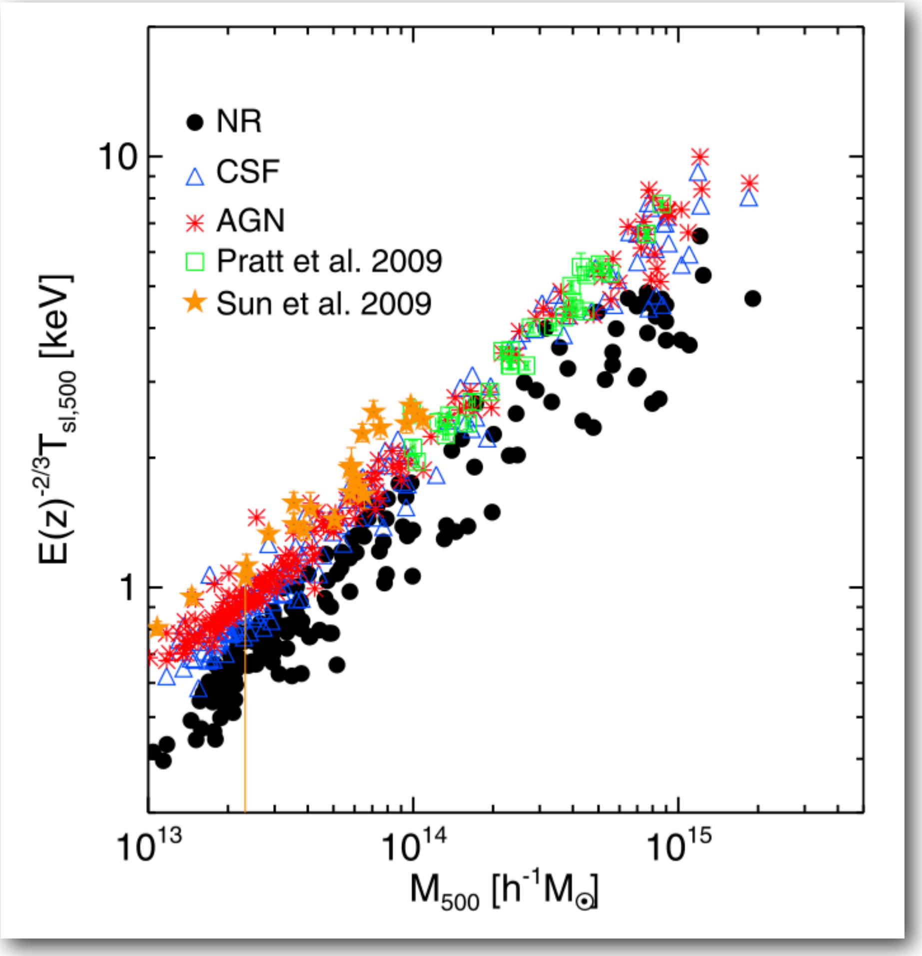
AGN feedback reduces stellar mass content of massive halos

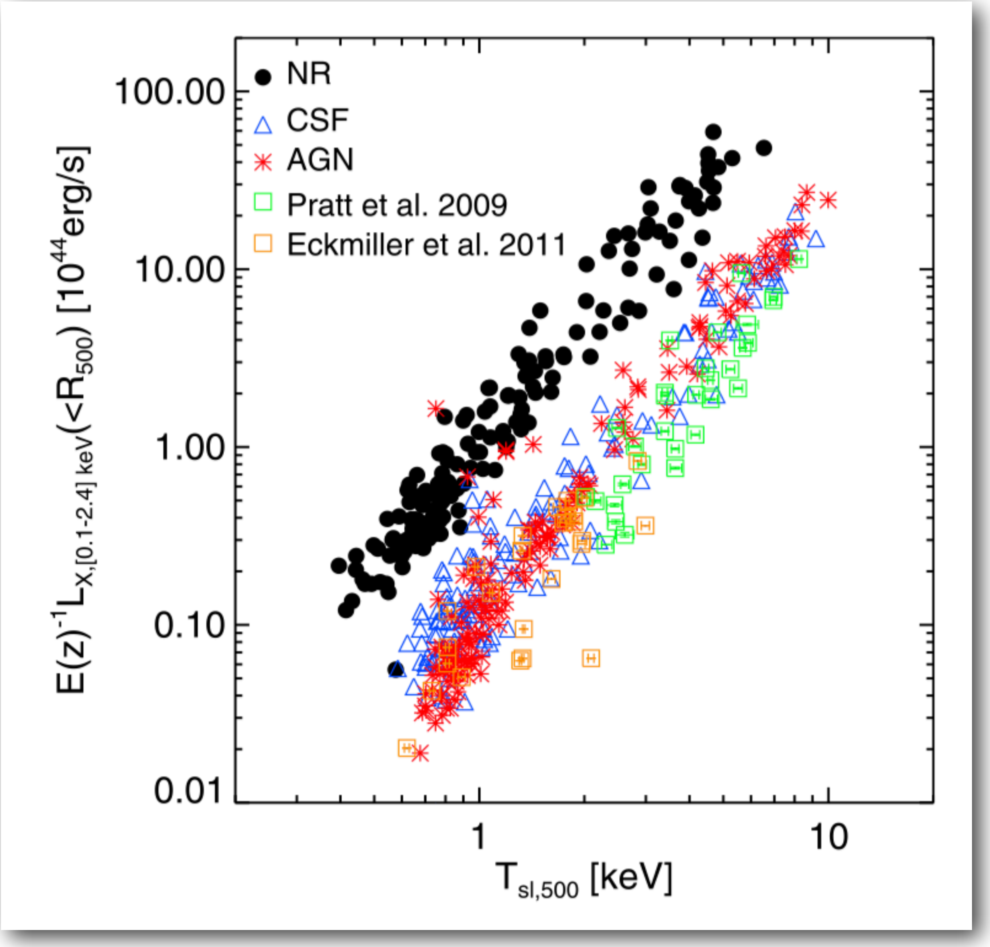
dramatic reduction relative to star formation feedback only models

Gadget simulations w/ AGN feedback

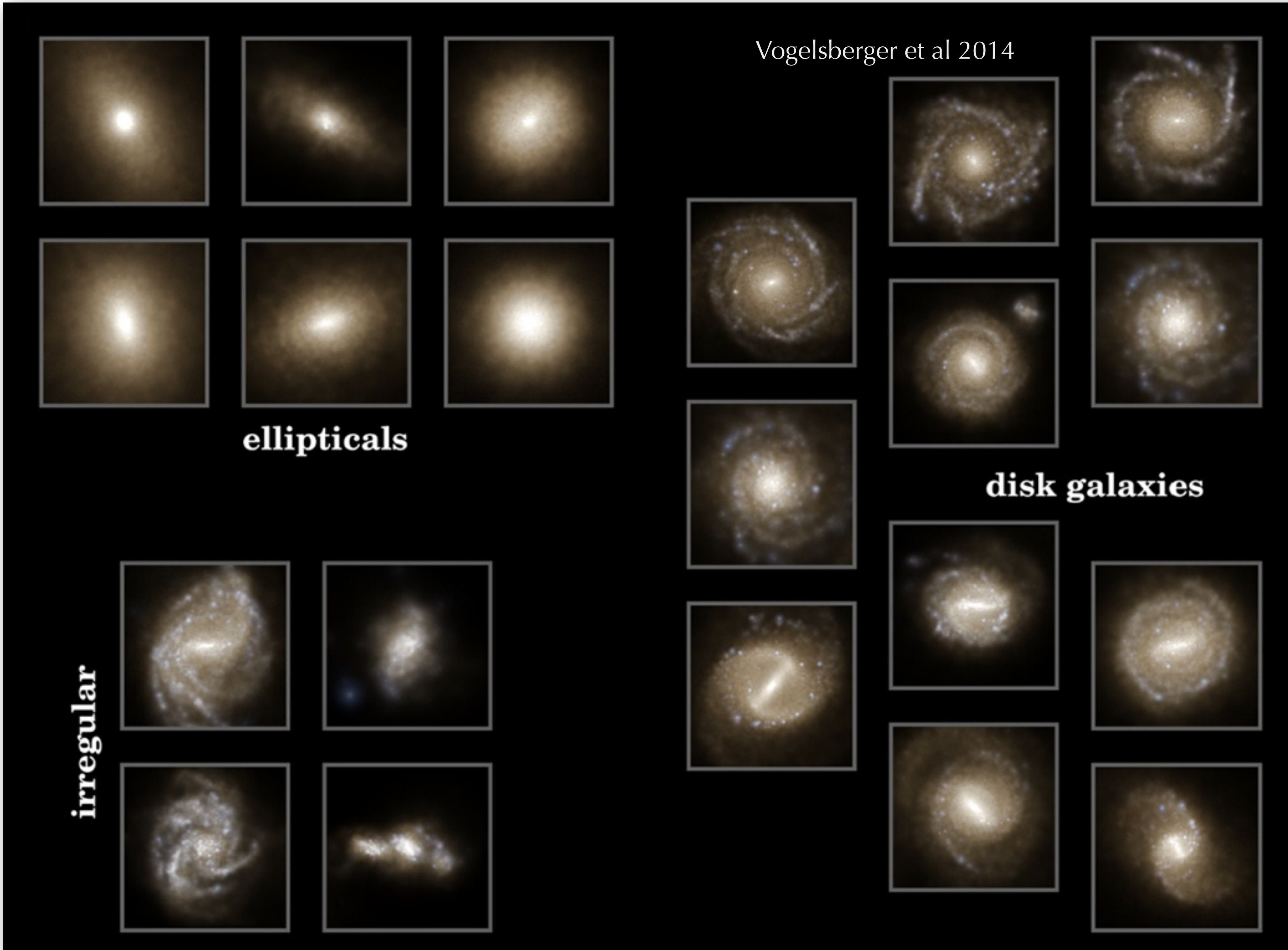
Planelles et al 2013

hot gas (intracluster medium) scaling relations with AGN offer better match to observations





Arepo simulation of 100 Mpc region w/ ~0.1kpc resolution : ***illustris***



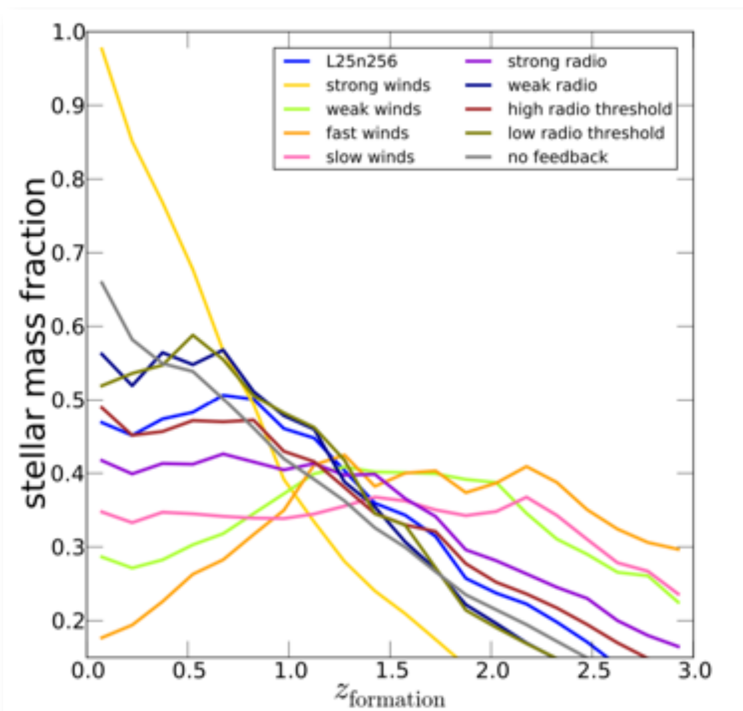
Arepo simulation of 100 Mpc region w/ ~0.1kpc resolution : ***illustris***

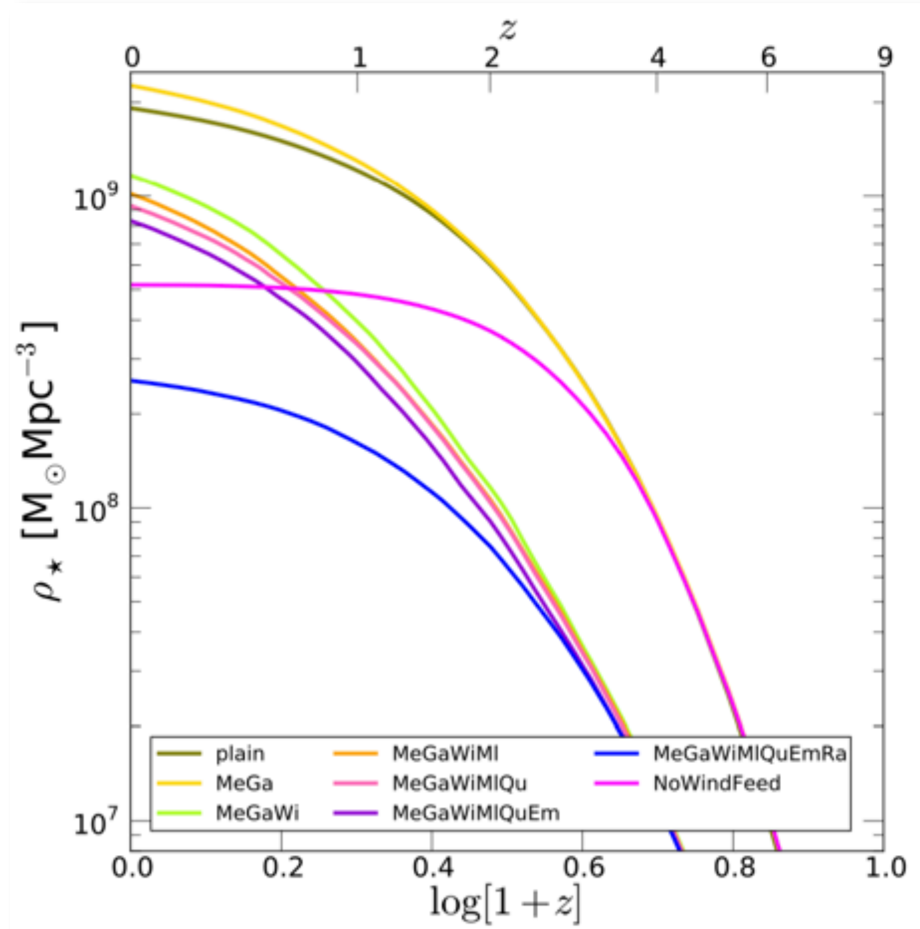
Movie?

AREPO models with layered physics, up to ***3 modes of AGN feedback***

Vogelsburger et al 2013

name	physics
plain	same as in Vogelsberger et al. (2012) (except for IMF, softer eEOS, self-shielding)
MeGa	+ met. line cool., gas recycl. = "no feedback"
MeGaWi	+ stellar winds
MeGaWiMl	+ separate metal mass loading of winds
MeGaWiMlQu	+ quasar-mode AGN feedback
MeGaWiMlQuEm	+ electro-magnetic AGN feedback
MeGaWiMlQuEmRa	+ radio-mode AGN = L25n256
NoWindFeed	L25n256 without stellar feedback





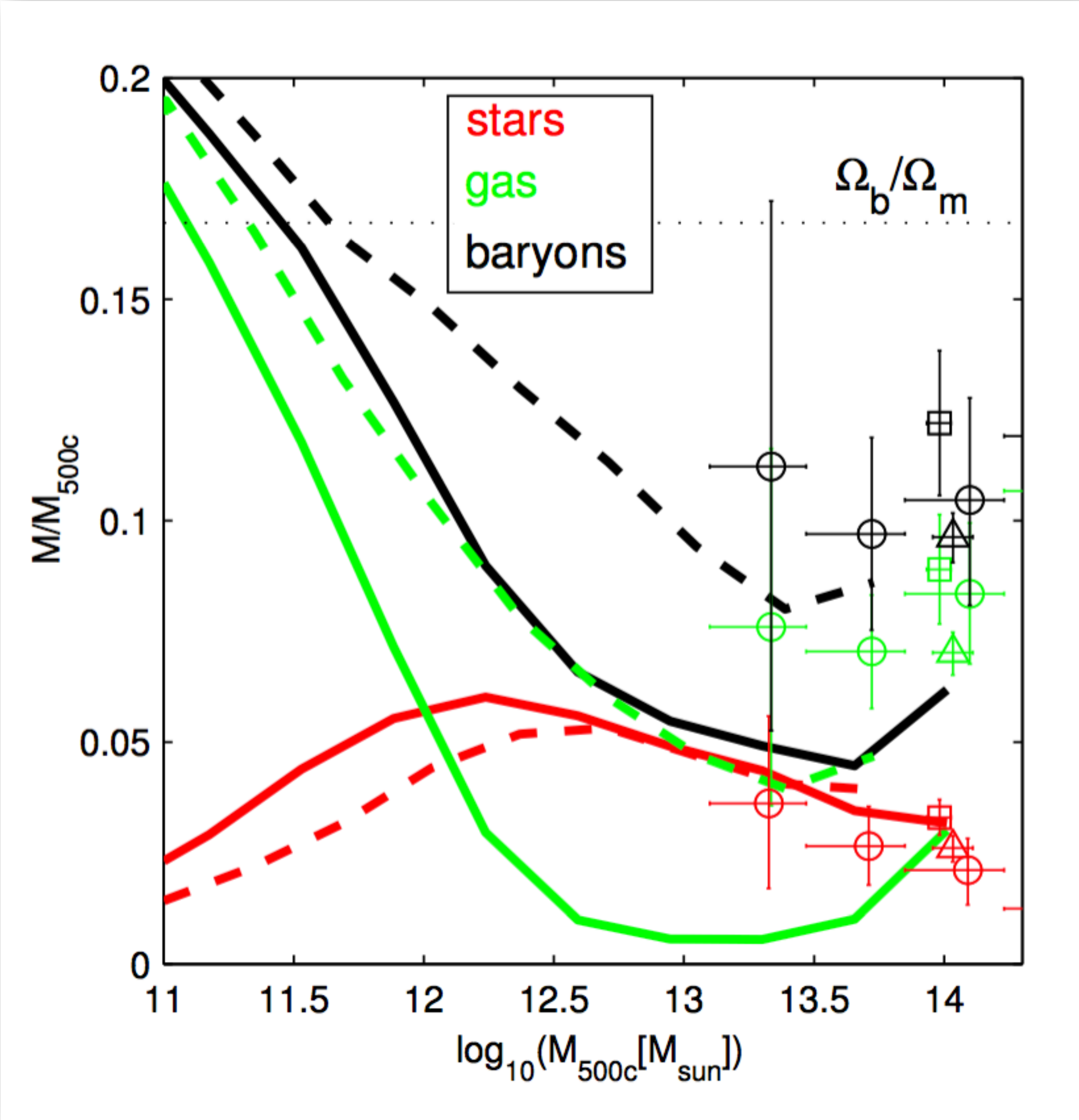
Illustris simulation: tunable parameter

Vogelsberger et al 2014

variable	fiducial value	description
stellar feedback (non-local energy-driven)		
κ_w	3.7	wind velocity relative to local DM 1D velocity dispersion
$\mathrm{egy_w}/\mathrm{egy_w^0}$	3.0	available SNII energy per formed stellar mass in units of $[\mathrm{egy_w^0}]$
$\mathrm{mom_w}\,/\,\mathrm{km\,s^{-1}}$	0	specific wind momentum in units of $[\,\mathrm{km\,s^{-1}}]$
AGN feedback (quasar-mode)		
ϵ_f	0.05	quasar mode feedback energy fraction
ϵ_r	0.2	radiative efficiency for Bondi accretion
AGN feedback (radio-mode)		
χ_radio	0.05	accretion rate threshold for radio-mode in units of $[\dot{M}_\mathrm{Edd}]$
δ_BH	1.15	duty cycle of radio-mode
ϵ_m	0.35	radio-mode feedback energy fraction
AGN feedback (electro-magnetic)		
ω_1, ω_2	0.3, 0.07	AGN obscuration parameterisation
wind metal loading		
γ_w	0.4	metal loading of wind particles

Illustris simulation doesn't solve everything

Genel et al 2014



hot gas fraction at z=0 (green solid) in galaxy groups lies far below observations

reflections on simulation-observation interplay: the ***trust*** issue

Vogelsberger et al 2014

> analytic models or hydrodynamical simulations. Special care needs to be taken when invoking new physical mechanisms for feedback or star formation to accommodate certain aspects of observational data if systematic errors have to be considered. Reproducing observational data perfectly does not imply a correct and physically meaningful model.

!!

fundamental issue: **uniqueness** in the presence of process **complexity**

* modeling star formation in direct gas dynamic simulations requires
  - shocks
  - cooling in a plasma heated by multiple processes (non-LTE...)
  - magnetic fields + cosmic ray heating
  - species chemistry and dust
  - mass loading and metal pollution by SN blastwaves
  - effects of jet heating from central BH (AGN activity)
  - +...

  All of this entails ***many 'sub-grid' control parameters***, effects of which can often be degenerate.

  Do purely *DETERMINISTIC* sub-grid solutions make sense?

  Are resolved solutions of local astrophysical conditions sufficient?

  Does `missing' small-scale physics add stochasticity?

* semi-analytic models already have >~100 input parameters :(

  (e.g., galacticus.org)

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– +...

All of this entails ***many ‘sub-grid’ control parameters***, effects of which can often be degenerate.

How do we know we’ve reached *THE* solution of nature?

Do purely *DETERMINISTIC* sub-grid solutions make sense?

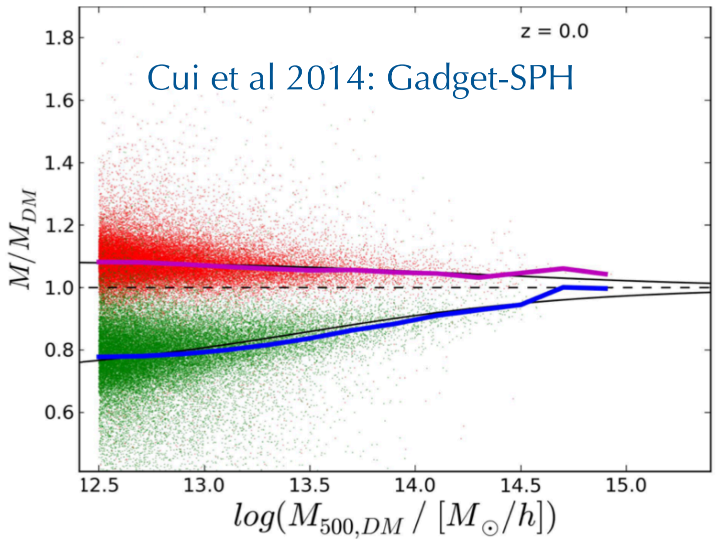
Are resolved solutions of local astrophysical conditions sufficient?

Does `missing’ small-scale physics add stochasticity?

* semi-analytic models already have >~100 input parameters :(

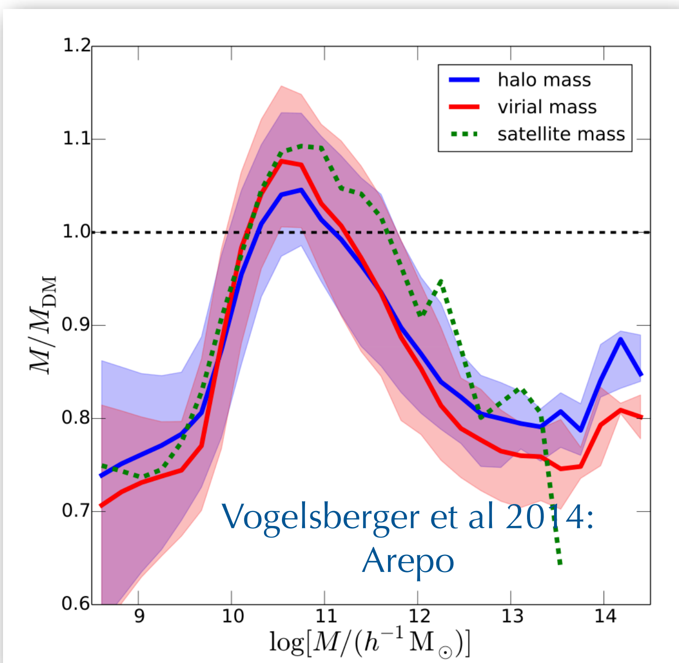
(e.g., galacticus.org)

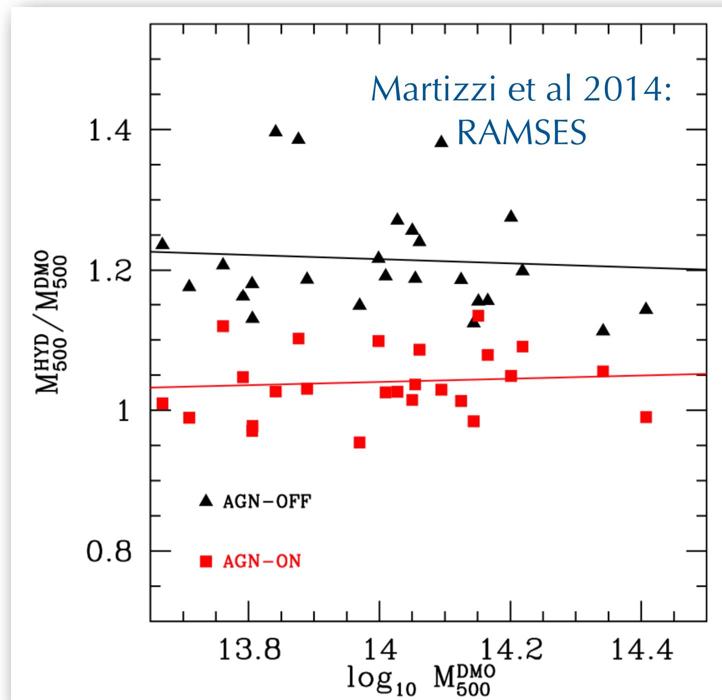
complexity example: the effect of baryons (w/ AGN) on final halo mass



← Halo mass is **suppressed** in SPH and Arepo

Halo mass is **enhanced** in RAMSES ↓





synthetic skies
in support of
dark energy studies

Dark Energy Survey

An NSF+DOE-funded study of dark energy using four techniques

1) Galaxy cluster surveys (with SPT)
2) Galaxy angular power spectrum
3) Weak lensing/cosmic shear
4) SN Ia distances

Two linked, multiband optical surveys

5000 deg^2 *g r i z Y* bands to ~24^{th} mag in r

Repeated observations of 40 deg^2

Development and schedule

Construction: 2007-2012

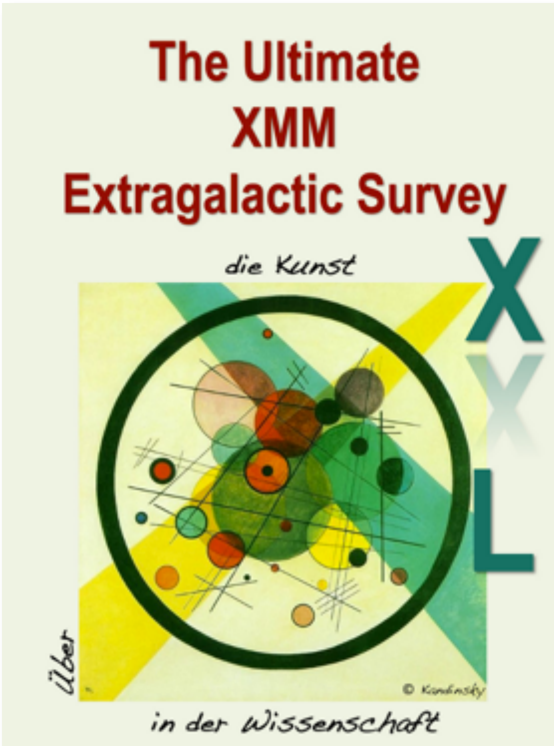
- New 3 deg^2 camera on Blanco 4m, Cerro Tololo
- Data management system at NCSA

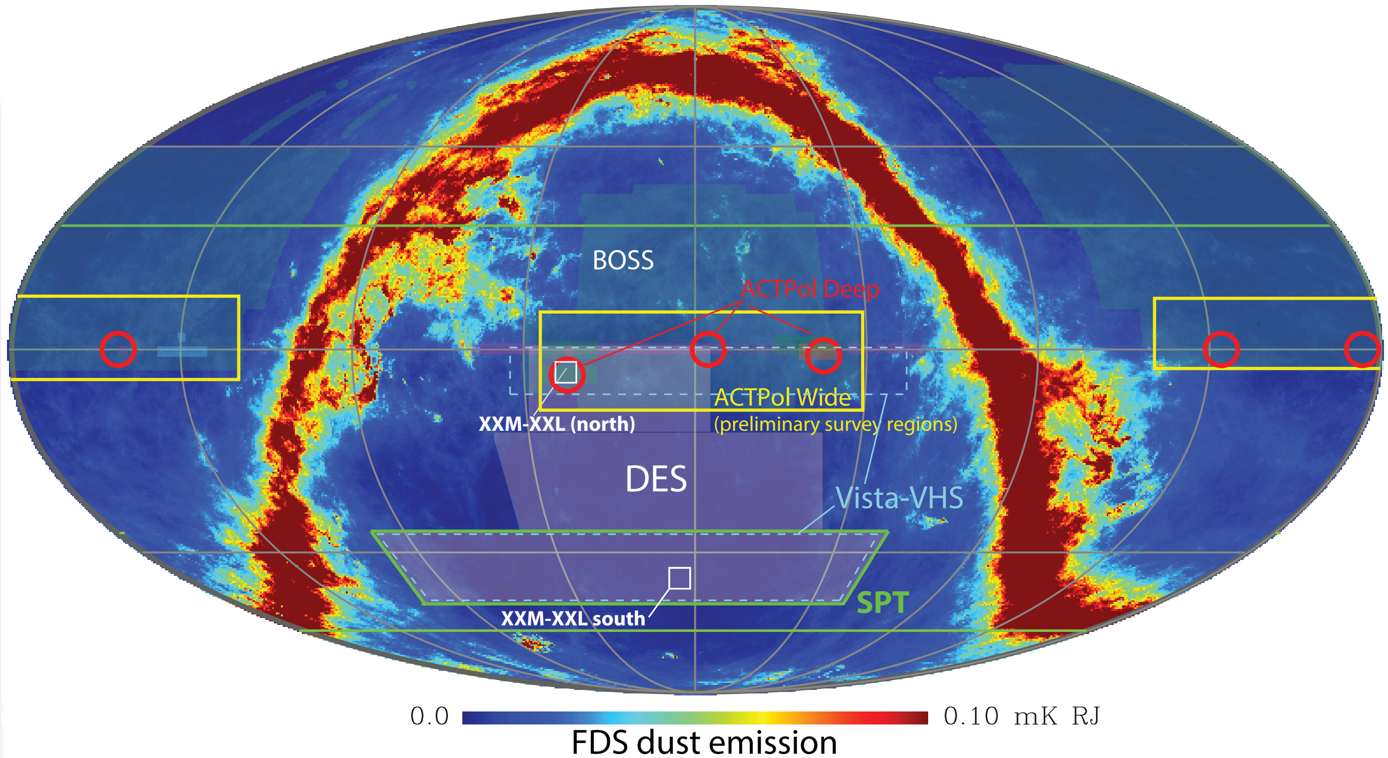
Survey Operations: 2013-2017

- 510 nights of telescope time over 5 years

XMM-XXL survey will use X-ray detected cluster counts to test DE

PI: M. Pierre (Saclay)

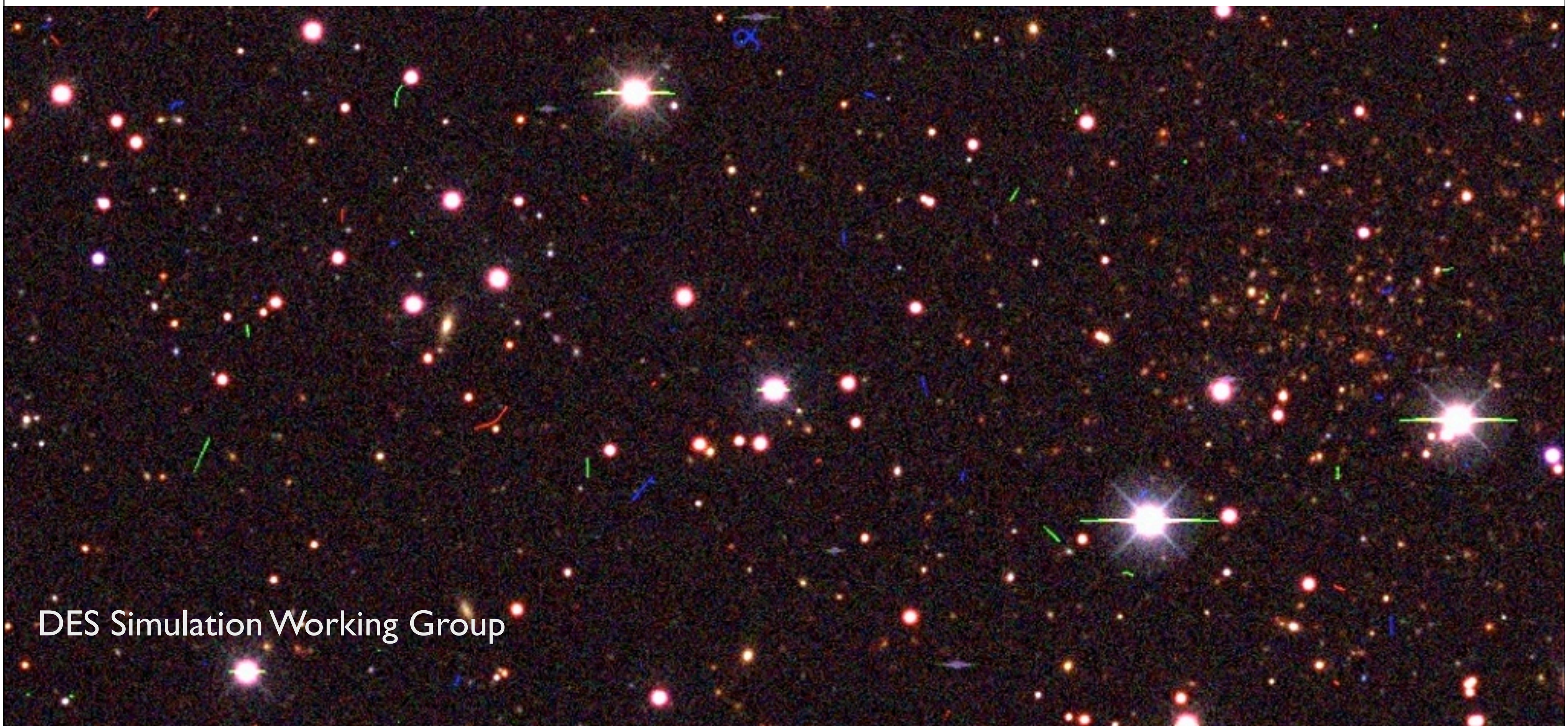




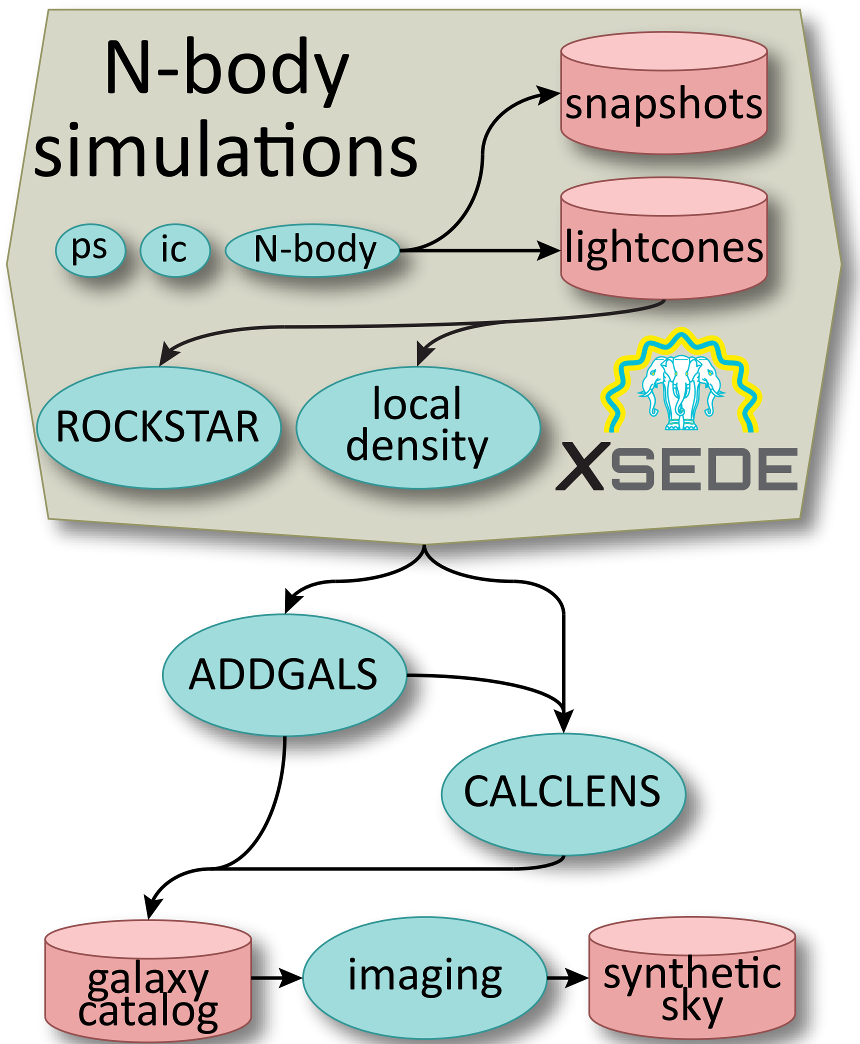
6 Msec XMM allocation, 10 ksec exposures
tiling 2 x 25 sq deg regions (S field overlaps DES)
+ multi-wavelength campaigns
expect ~500 groups and clusters to z~2

Survey-specific simulations enable **key capabilities:**

* test signal extraction from survey data (create *truth catalogs*)
* predict (range of) observable features for a cosmology
* calculate expected signal covariance



simulation workflow to support Dark Energy Survey (DES) science analysis



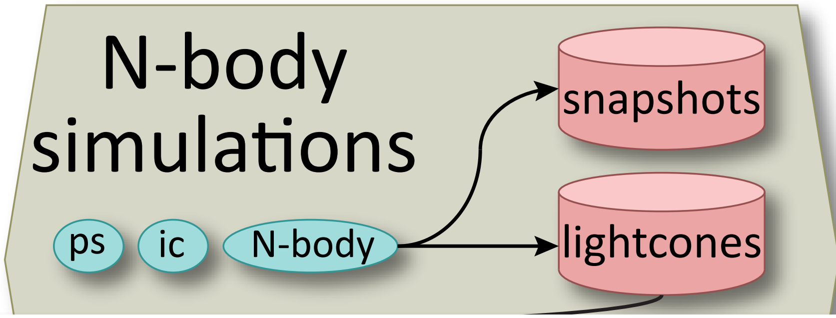
<u>Catalog Simulations</u>

M. Becker (Stanford)
M. Busha (Stanford)
B. Erickson (Michigan)
A. Evrard (Michigan)
A. Kravtsov (Chicago)
R. Wechsler (Stanford)

<u>Image Simulations</u>

H. Lin (Fermilab)
Nikolai Kuropatkin (Fermilab)
+ DES Data Management

simulation workflow to support Dark Energy Survey (DES) science analysis



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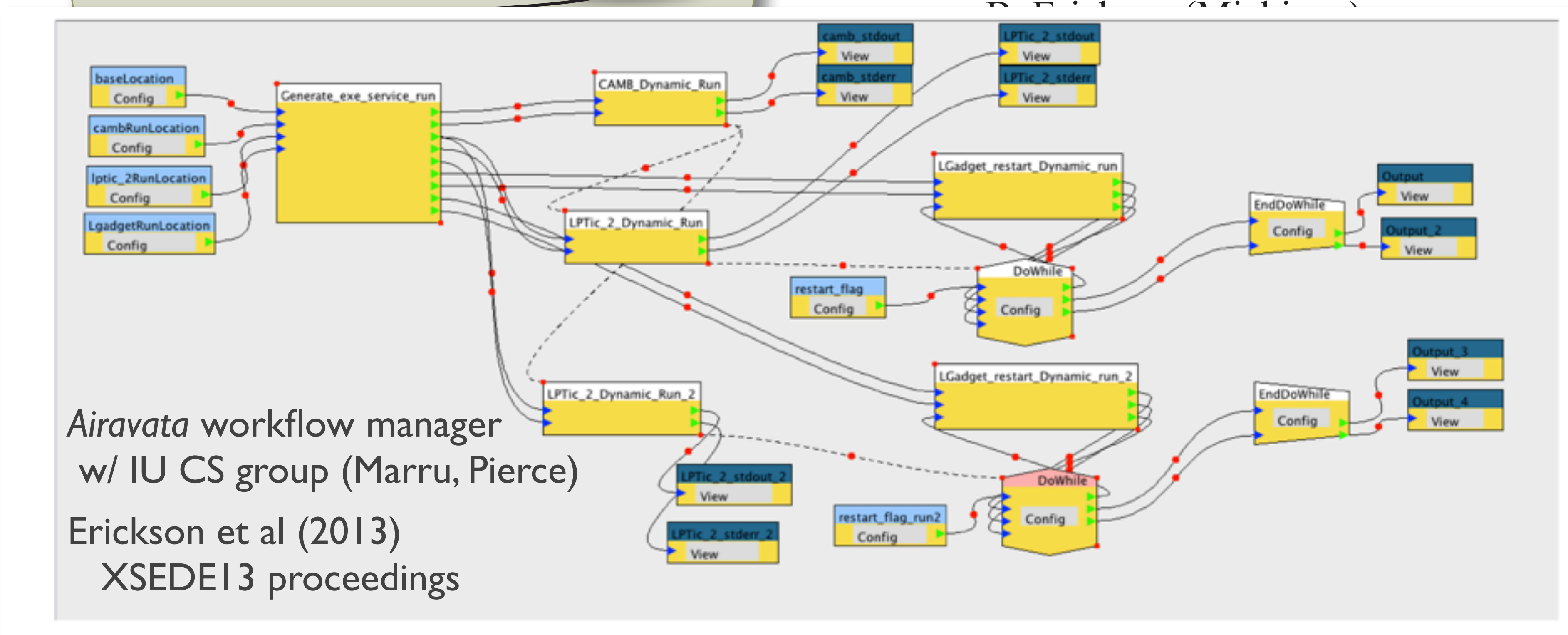
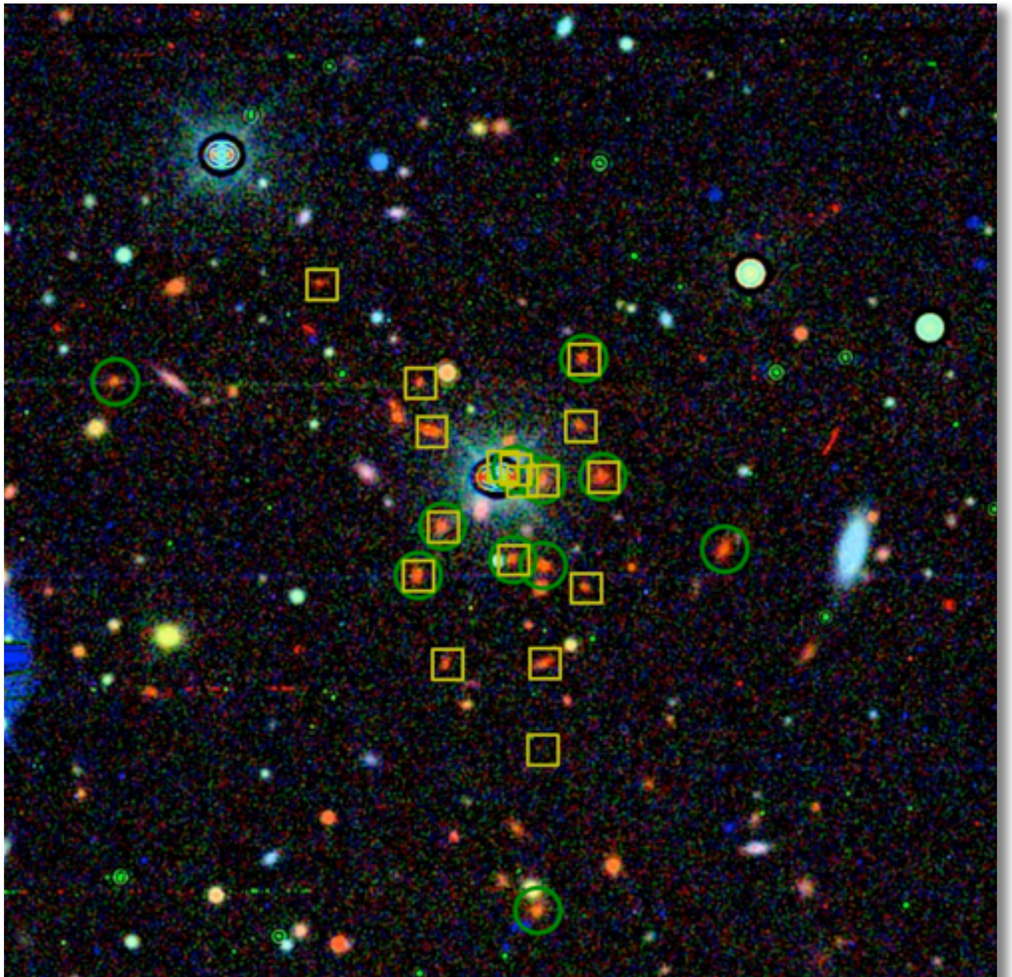
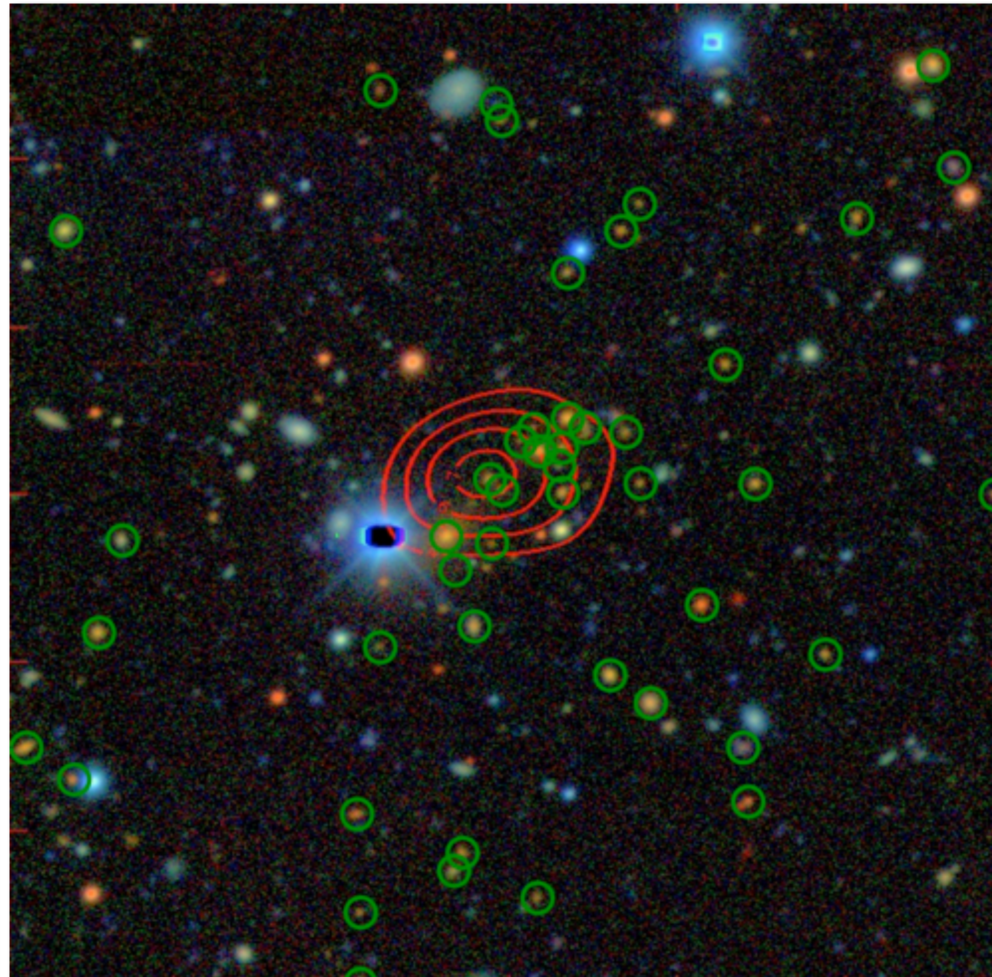


Airavata workflow manager
w/ IU CS group (Marru, Pierce)

Erickson et al (2013)
XSEDE13 proceedings

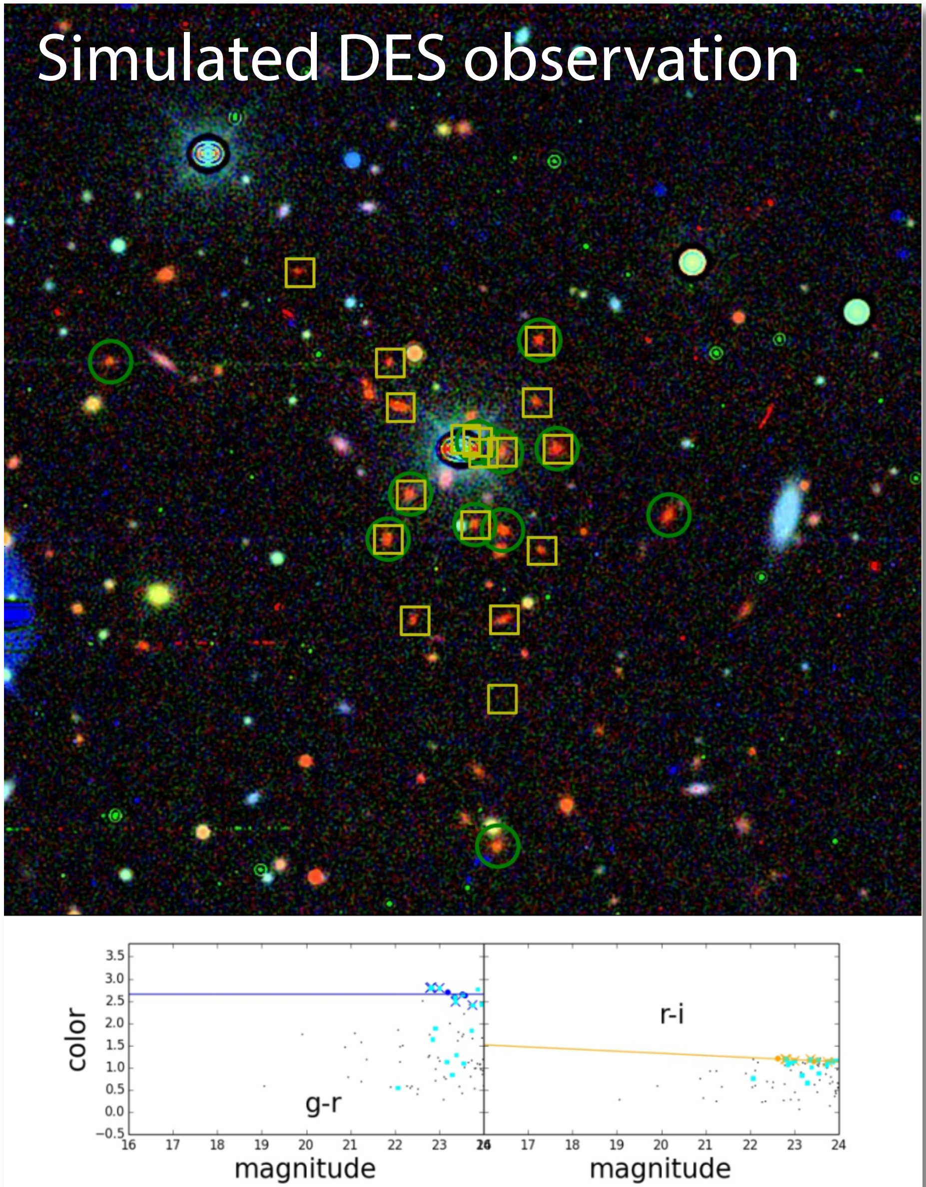
one real and one synthetic galaxy cluster image from DECam: *which is which?*

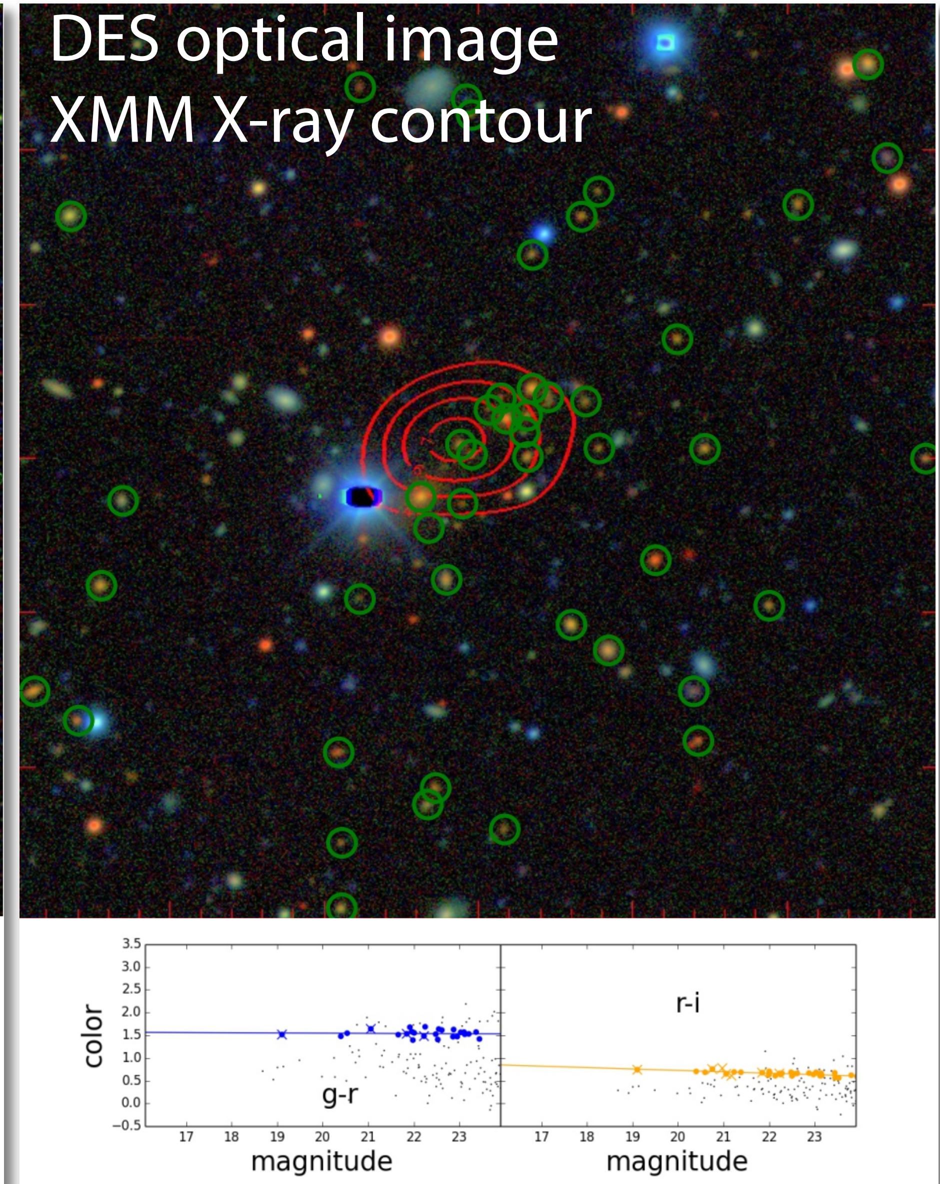




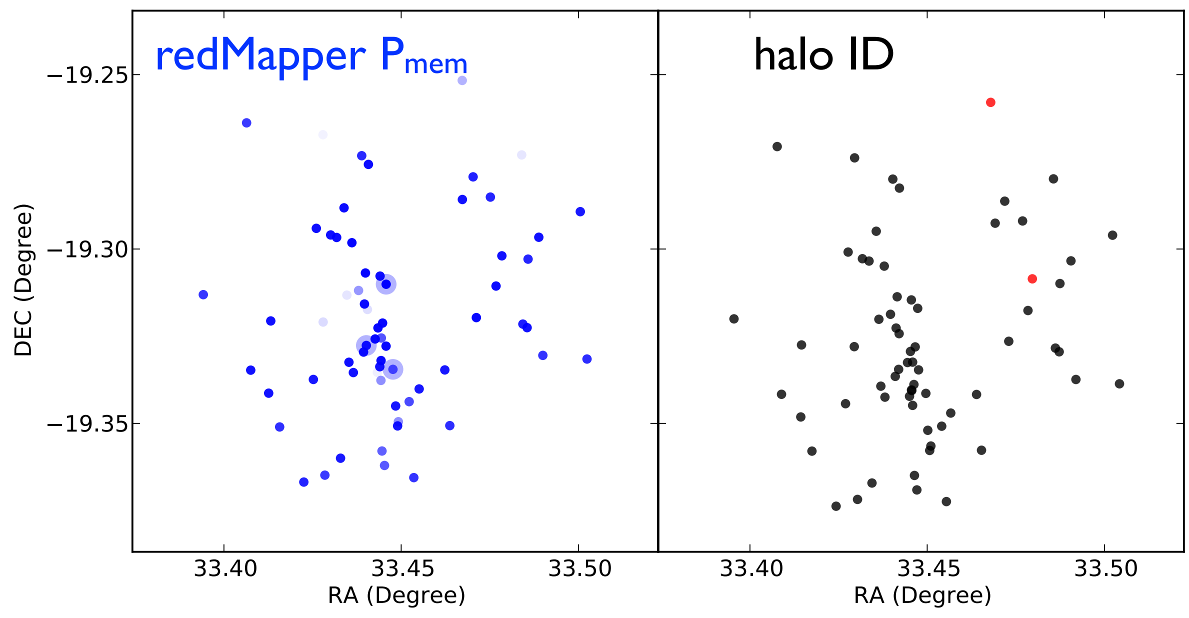
Courtesy Chris Miller (Michigan)

synthetic and real galaxy cluster images from DECam





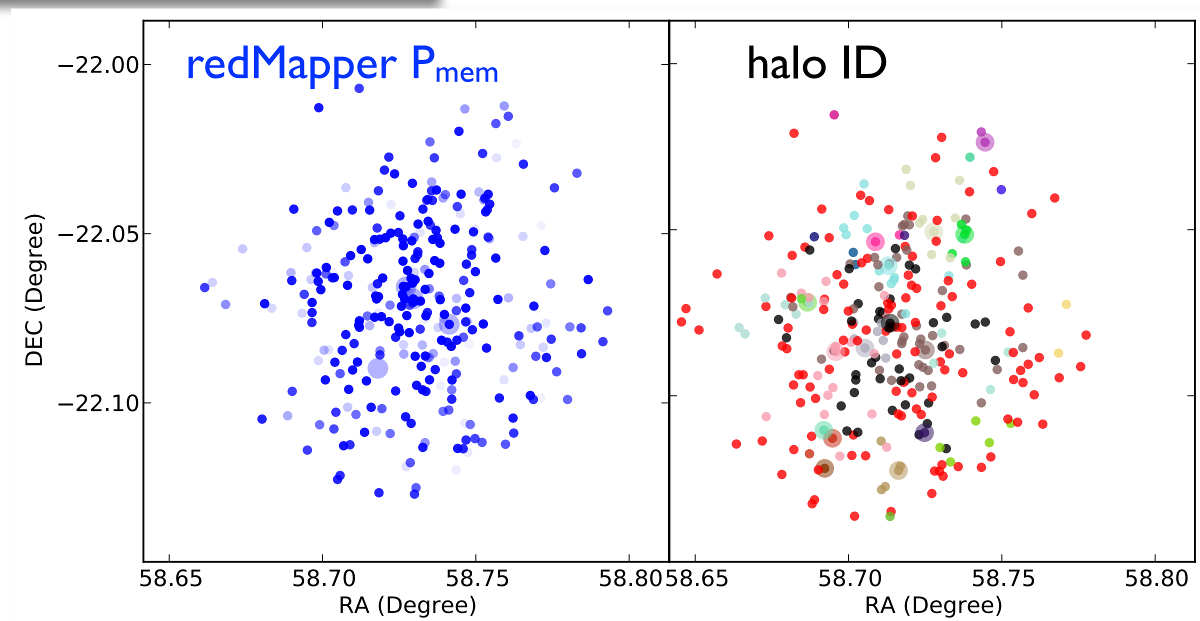
connecting clusters to halos through galaxy membership



Farahi et al, in prep.

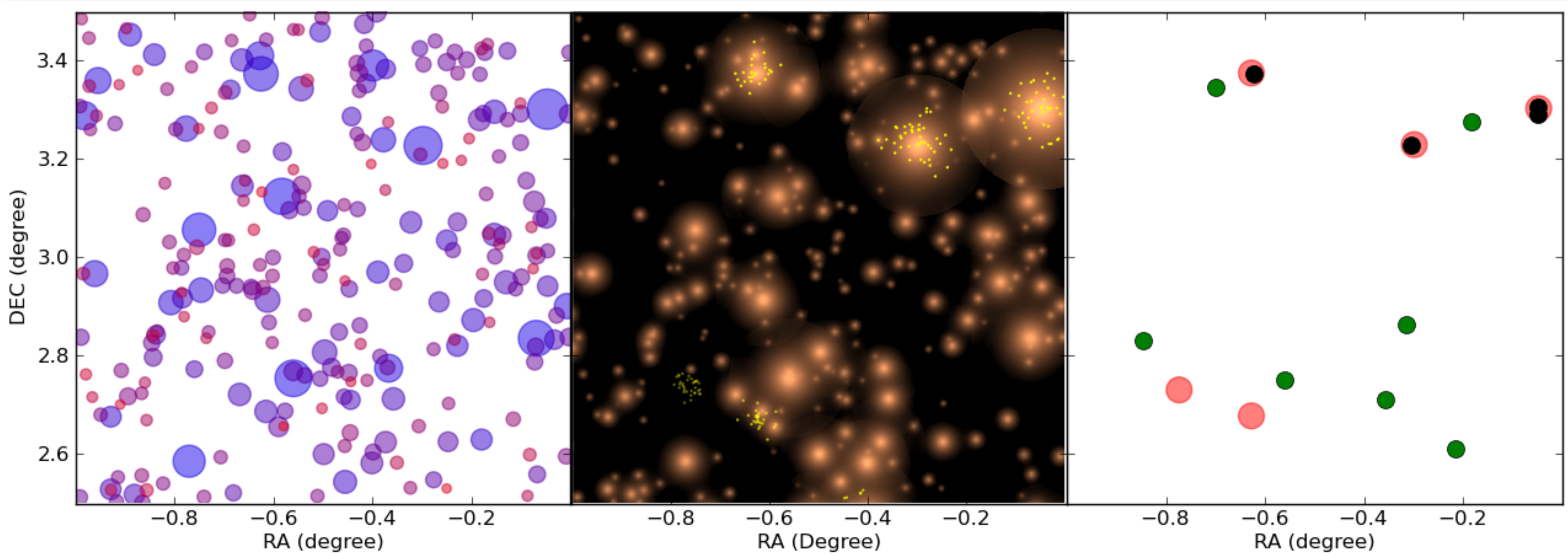
"clean" cluster-halo match

"dirty" cluster-halo match



testing methods to find clusters using X-ray of optical signatures

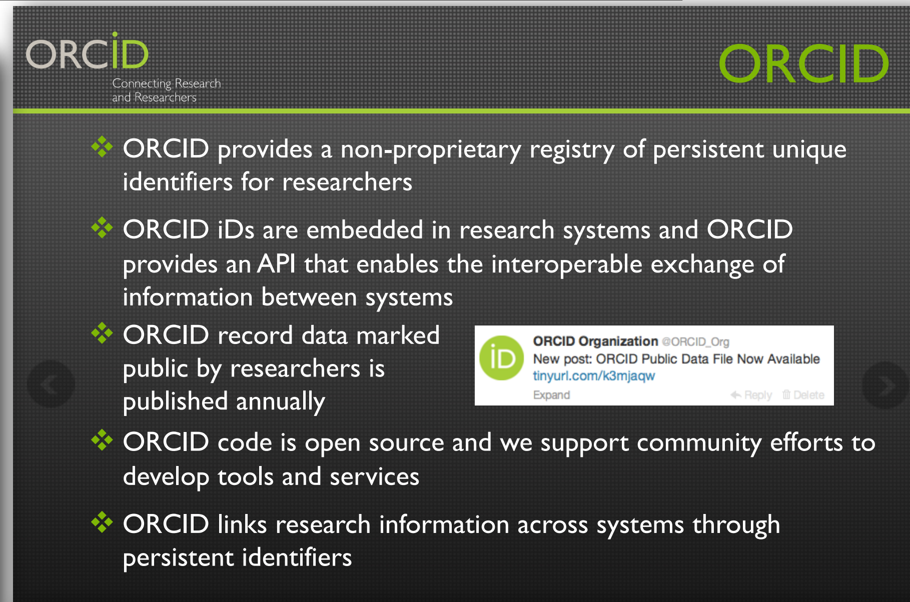
Farahi et al, in prep.

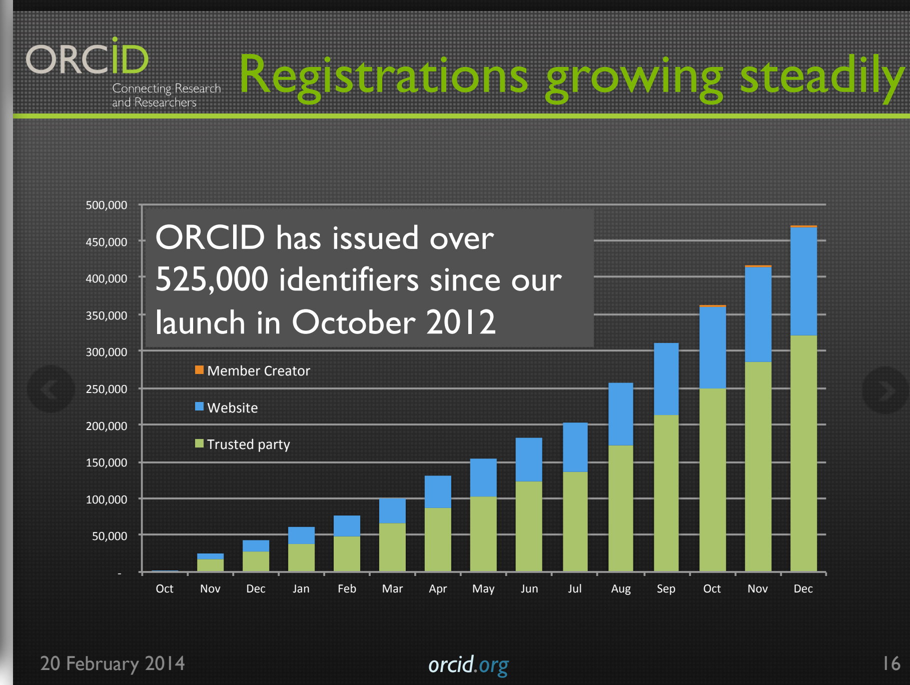


halos (angular size; color = redshift)	X-ray emission (greyscale) + optical galaxies in r *redMaPPer* clusters (points)	X-ray cluster detections (**black**/**green** = hi/lo S/N) + optically-detected *redMaPPer* clusters (**red**)

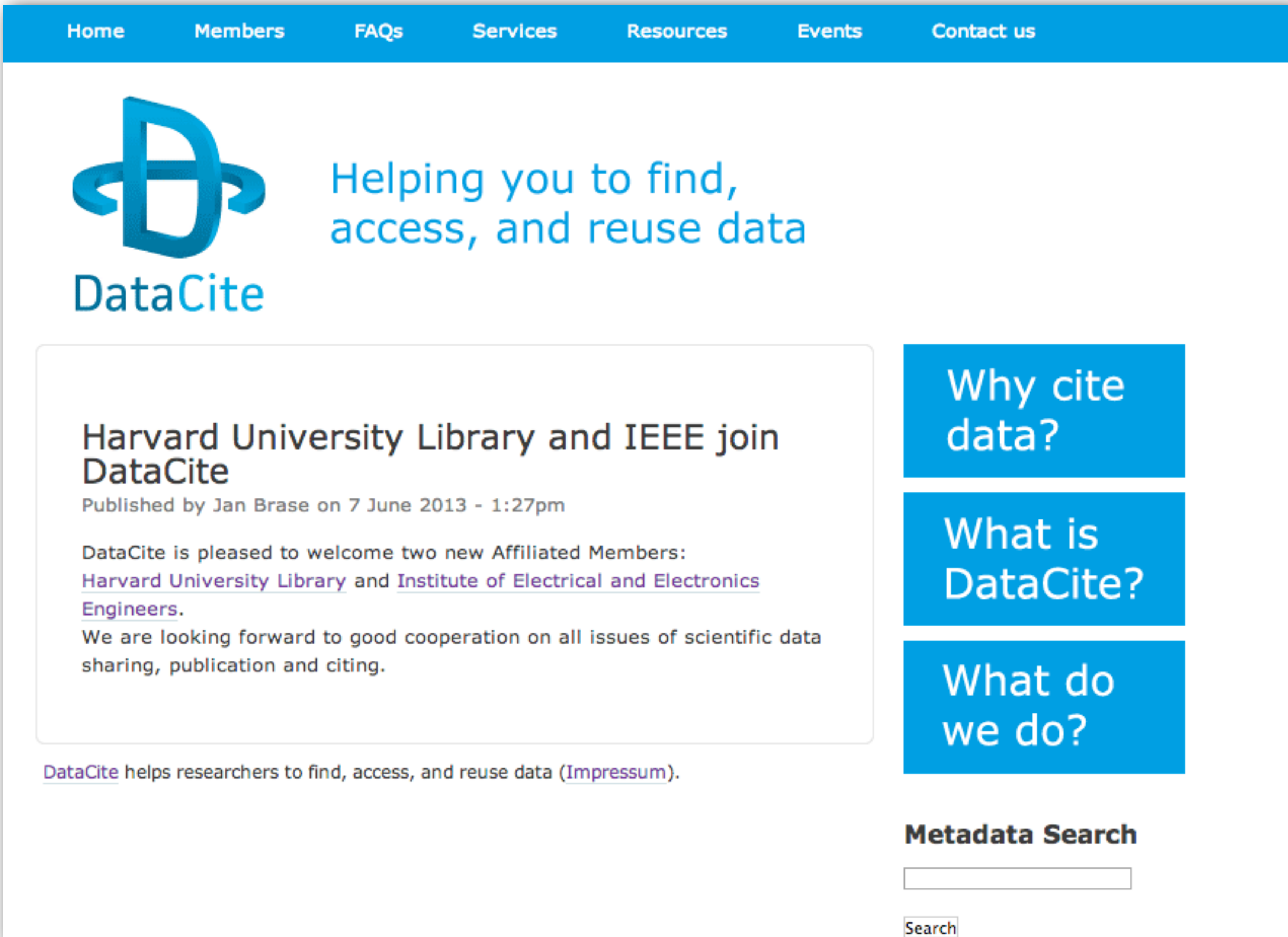
thoughts on
post-adolescent
computational cosmology

personal identity: a foundation for scholarly reputation

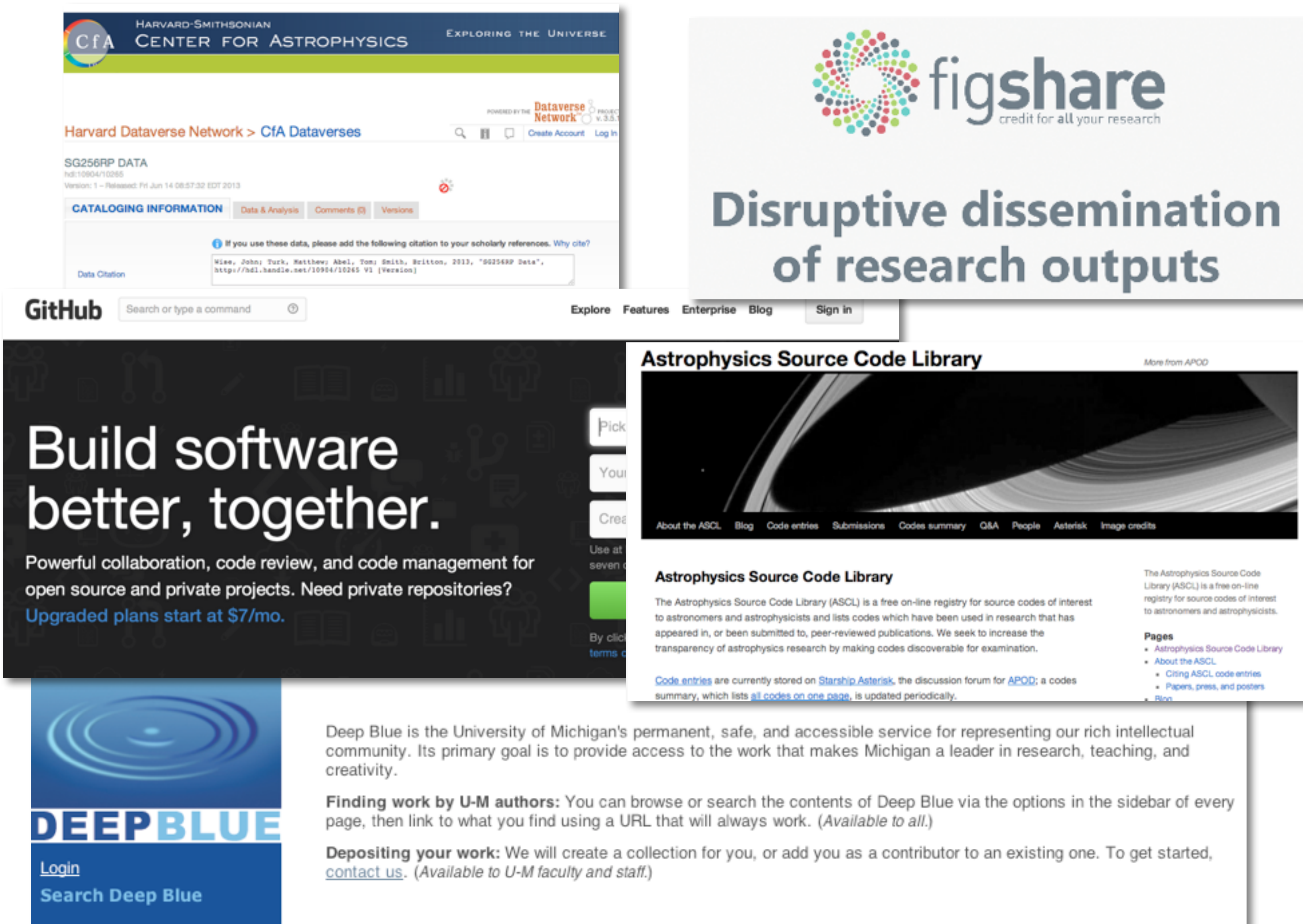




data identity: a foundation for re-use and reproducibility



data/code repositories: emerging public/private solutions



a notional stack of production/access layers

ORCID / DataCite / alt-metrics

- provisioning (condo-of-condo's?, ...)
- portal (hub-zero, ...)
- workflow (galaxy, pegasus, Airavata, ...)
- analysis, post-proc. (yt, RT, ADDGALS, XMAS, ...)
- cosmo codes (Enzo, RAMSES, gadget, ART, ...)
- physics modules (TBD)

Goal:
Improve community coordination to maximize science return of exascale simulations.

a notional stack of production/access layers

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Goal: Improve community coordination to maximize science return of exascale simulations.

How to build across institutions?
How to incentivize community alignment?

positive developments in community structure

nIFTy Cosmology: NUMERICAL SIMULATIONS FOR LARGE SURVEYS

SOC:
Alexander Knebe (UAM)
Frazer Pearce (Nottingham)
Juan Garcia-Bellido (UAM/IFT-CSIC)
Chris Power (Western Australia)
Richard Bower (Durham)

The AGORA High-resolution Galaxy Simulations Comparison

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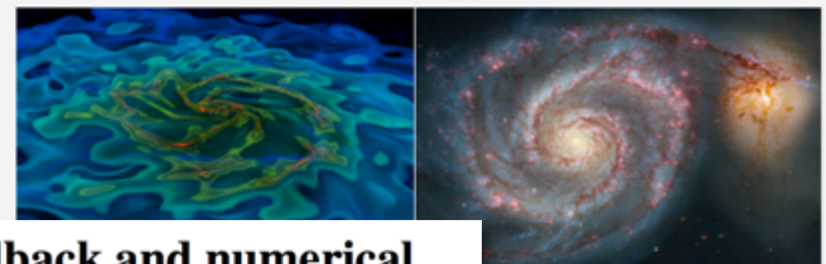
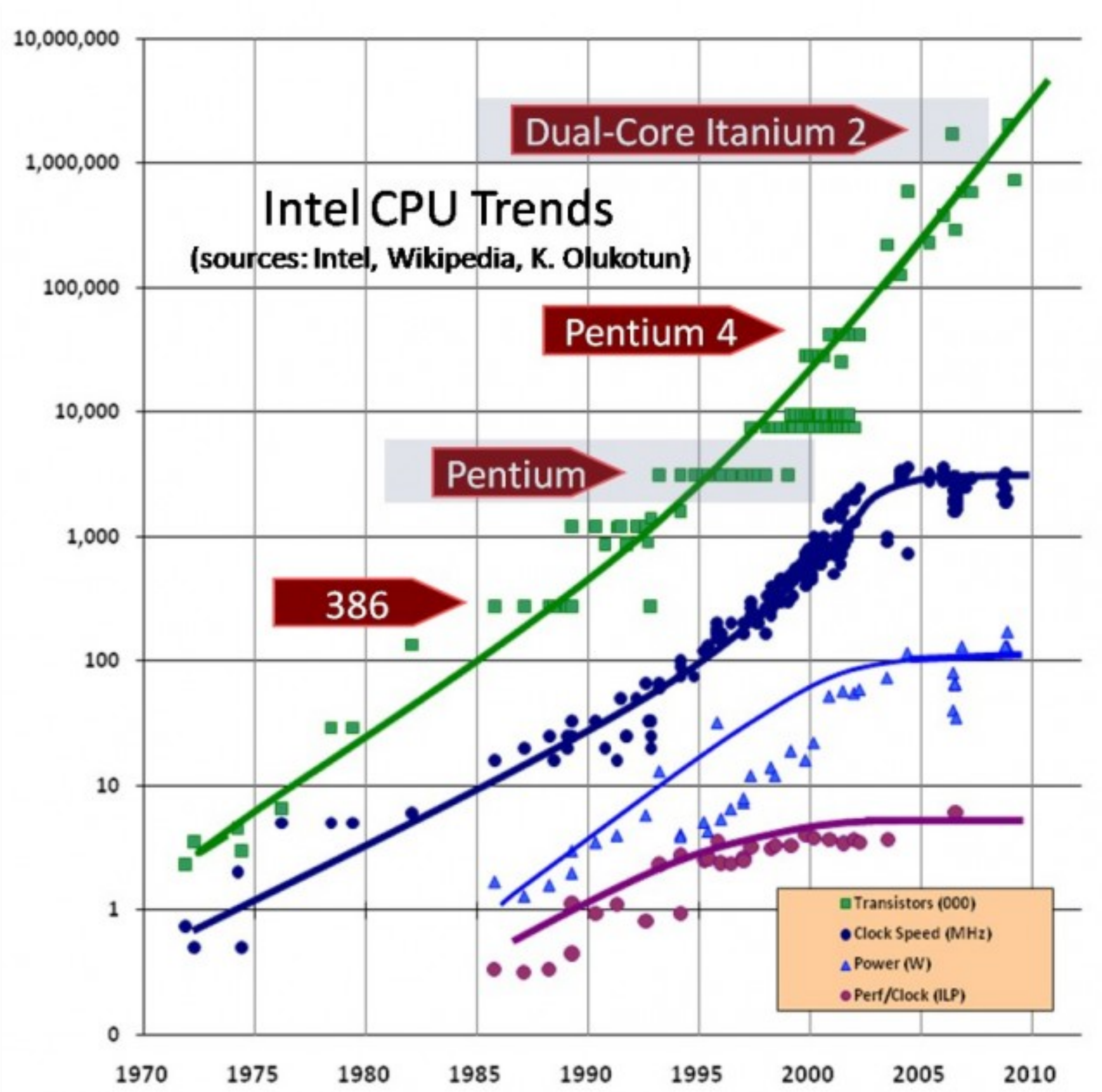


The Aquila comparison project: the effects of feedback and numerical methods on simulations of galaxy formation

C. Scannapieco,[1★] M. Wadepuhl,[2] O. H. Parry,[3,4] J. F. Navarro,[5] A. Jenkins,[3] V. Springel,[6,7] R. Teyssier,[8,9] E. Carlson,[10] H. M. P. Couchman,[11] R. A. Crain,[12,13] C. Dalla Vecchia,[14] C. S. Frenk,[3] C. Kobayashi,[15,16] P. Monaco,[17,18] G. Murante,[17,19] T. Okamoto,[20] T. Quinn,[10] J. Schaye,[13] G. S. Stinson,[21] T. Theuns,[3,22] J. Wadsley,[11] S. D. M. White[2] and R. Woods[11]

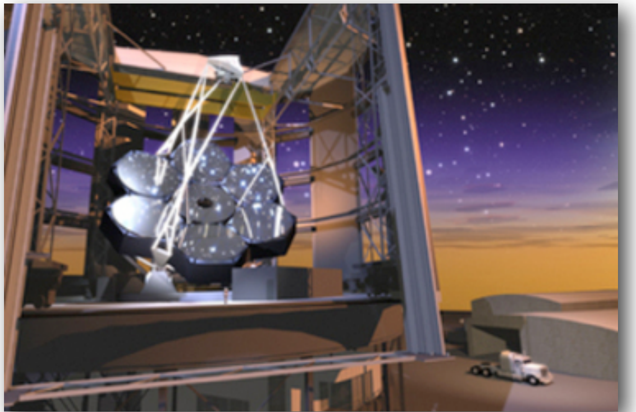
Moore's law has been true for almost four decades - how much longer?



Bob Colwell,
formerly Intel,
now DARPA
@ Hot Chips 2013

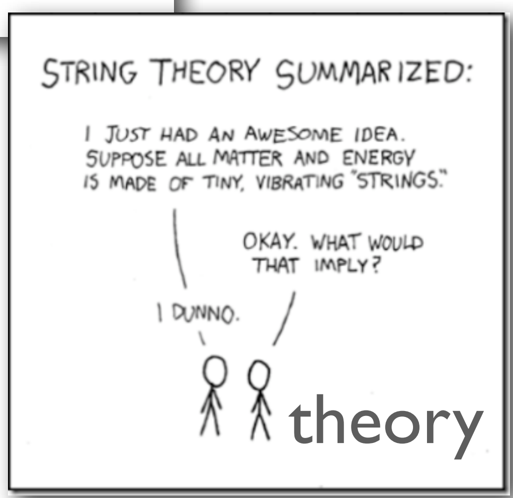
"For planning horizons, I pick 2020 as the earliest date we could call [Moore's law] dead," Colwell said. "You could talk me into 2022, but whether it will come at 7 or 5nm, it's a big deal."

http://www.extremetech.com/computing/165331-intels-former-chief-architect-moores-law-will-be-dead-within-a-decade



observation

simulation



theory

a maturing future:

- simulations as an *integral element* of large survey projects
  - planning + funding issues
- synthetic multi-wavelength skies available to perform cross-survey science analysis
  - (cyber-)infrastructure issues

GOAL:
gain physical understanding by

precision ***measurement***

and

accurate ***modeling***

The End