

Variability of the proton-to-electron mass ratio on cosmological scales

Quasar absorption line spectroscopy

Martin Wendt – Hamburger Sternwarte



Osservatorio Astronomico di Trieste, November 28th

Overview

- Short introduction of theory behind variation
- How is variation reflected in observations?
- Molecular Hydrogen H_2
- Methods involved
- Analysis
- Summary and Outlook

Theory of shilly-shally constants

Theory of shilly-shally constants

Kaluza-Klein theories

- Theodor Kaluza: generalisation of Einstein's GTR and the Maxwell equations in five dimensions in 1921

Theory of shilly-shally constants

Kaluza-Klein theories

- Theodor Kaluza: generalisation of Einstein's GTR and the Maxwell equations in five dimensions in 1921
- Oscar Klein: fifth dimension possibly curled up

Theory of shilly-shally constants

Kaluza-Klein theories

- Theodor Kaluza: generalisation of Einstein's GTR and the Maxwell equations in five dimensions in 1921
- Oscar Klein: fifth dimension possibly curled up
- fundamental constants only constant in e.g. five-dimensional space

Theory of shilly-shally constants

Kaluza-Klein theories

- Theodor Kaluza: generalisation of Einstein's GTR and the Maxwell equations in five dimensions in 1921
- Oscar Klein: fifth dimension possibly curled up
- fundamental constants only constant in e.g. five-dimensional space
- observed constants a mere projection

Theory of shilly-shally constants

Kaluza-Klein theories

- Theodor Kaluza: generalisation of Einstein's GTR and the Maxwell equations in five dimensions in 1921
- Oscar Klein: fifth dimension possibly curled up
- fundamental constants only constant in e.g. five-dimensional space
- observed constants a mere projection
- another byproduct: scalar field as possible source for acceleration

How to measure variation?

- possible variation of the finestructure constant $\alpha = e^2/4\pi\hbar c \approx 1/137$.
Results under heavy debate.

How to measure variation?

- possible variation of the finestructure constant $\alpha = e^2/4\pi\hbar c \approx 1/137$.
Results under heavy debate.
- variation of the gravitational constant G_N .
Recent paper last week:
 $\dot{G}_N/G_N \lesssim 10^{-17} \text{yr}^{-1}$.

How to measure variation?

How to measure variation?

- possible variation of the proton-to-electron mass ratio $\mu = \frac{m_p}{m_e}$.

$$\mu_0 = 1836.15267261(85) \text{ (Mohr \& Taylor 2000)}$$

How to measure variation?

- possible variation of the proton-to-electron mass ratio $\mu = \frac{m_p}{m_e}$.

$$\mu_0 = 1836.15267261(85) \text{ (Mohr \& Taylor 2000)}$$

- laboratory experiments not yet very accurate (5 years)

How to measure variation?

- possible variation of the proton-to-electron mass ratio $\mu = \frac{m_p}{m_e}$.

$$\mu_0 = 1836.15267261(85) \text{ (Mohr \& Taylor 2000)}$$

- laboratory experiments not yet very accurate (5 years)
- measure possible variation on cosmological scales

How to measure variation?

molecular hydrogen H₂ - energy levels

- $E_{\text{total}} = E_{\text{electronic}} + E_{\text{vibration}} + E_{\text{rotation}}$ (BOA)

How to measure variation?

molecular hydrogen H₂ - energy levels

- $E_{\text{total}} = E_{\text{electronic}} + E_{\text{vibration}} + E_{\text{rotation}}$ (BOA)
- vibrational: $E_v = \left(v + \frac{1}{2}\right) \bar{\omega}_{\text{osc}}$

How to measure variation?

molecular hydrogen H₂ - energy levels

- $E_{\text{total}} = E_{\text{electronic}} + E_{\text{vibration}} + E_{\text{rotation}}$ (BOA)
- vibrational: $E_v = \left(v + \frac{1}{2}\right) \bar{\omega}_{\text{osc}}$
- with $\omega_{\text{osc}} = \frac{1}{2\pi c} \cdot \sqrt{\frac{k}{\mu}}$

How to measure variation?

molecular hydrogen H₂ - energy levels

- $E_{\text{total}} = E_{\text{electronic}} + E_{\text{vibration}} + E_{\text{rotation}}$ (BOA)
- vibrational: $E_v = \left(v + \frac{1}{2}\right) \bar{\omega}_{\text{osc}}$
- with $\omega_{\text{osc}} = \frac{1}{2\pi c} \cdot \sqrt{\frac{k}{\mu}}$
- the classical oscillation frequency dependent on the *reduced mass* as $\mu^{-\frac{1}{2}}$.

How to measure variation?

molecular hydrogen H₂ - energy levels

- rotational: $I = \frac{m_1 m_2}{m_1 + m_2} r_0^2 = \mu r_0^2$

How to measure variation?

molecular hydrogen H₂ - energy levels

- rotational: $I = \frac{m_1 m_2}{m_1 + m_2} r_0^2 = \mu r_0^2$
- $E_J = BJ(J + 1); J = 0, 1, 2, 3, \dots$

How to measure variation?

molecular hydrogen H₂ - energy levels

- rotational: $I = \frac{m_1 m_2}{m_1 + m_2} r_0^2 = \mu r_0^2$
- $E_J = BJ(J + 1); J = 0, 1, 2, 3, \dots$
- $B = \frac{h}{8\pi^2 I c}$

How to measure variation?

molecular hydrogen H₂ - energy levels

- rotational: $I = \frac{m_1 m_2}{m_1 + m_2} r_0^2 = \mu r_0^2$
- $E_J = BJ(J + 1); J = 0, 1, 2, 3, \dots$
- $B = \frac{h}{8\pi^2 I c}$
- rotational transitions are proportional to μ^{-1}

How to measure variation?

- homonuclear molecule - no detectable mere vibrational or rotational spectrum

How to measure variation?

- homonuclear molecule - no detectable mere vibrational or rotational spectrum
- only observable in combinations with electronic transitions (UV-Band)

How to measure variation?

- homonuclear molecule - no detectable mere vibrational or rotational spectrum
- only observable in combinations with electronic transitions (UV-Band)
- UV radiation is a very efficient dissociator of H_2 , so any H_2 that survived would presumably be located inside very dense interstellar clouds.

How to measure variation?

- homonuclear molecule - no detectable mere vibrational or rotational spectrum
- only observable in combinations with electronic transitions (UV-Band)
- UV radiation is a very efficient dissociator of H_2 , so any H_2 that survived would presumably be located inside very dense interstellar clouds.
- So far observations have borne out this supposition.

How to measure variation?

molecular hydrogen H_2

How to measure variation?

molecular hydrogen H_2

- electron-vibro-rotational transitions depend on reduced mass of molecule
- different dependence for different transitions

How to measure variation?

molecular hydrogen H₂

- electron-vibro-rotational transitions depend on reduced mass of molecule
- different dependence for different transitions
- distinguish cosmological redshift of a line from the shift caused by possible variation of μ

How to measure variation?

molecular hydrogen H₂

- electron-vibro-rotational transitions depend on reduced mass of molecule
- different dependence for different transitions
- distinguish cosmological redshift of a line from the shift caused by possible variation of μ

- $$\lambda_{\text{obs}} = \lambda_{\text{rest}} \times (1 + z_{\text{abs}}) \left(1 + K_i \frac{\Delta\mu}{\mu}\right)$$

How to measure variation?

molecular hydrogen H₂

- electron-vibro-rotational transitions depend on reduced mass of molecule
- different dependence for different transitions
- distinguish cosmological redshift of a line from the shift caused by possible variation of μ

- $$\lambda_{\text{obs}} = \lambda_{\text{rest}} \times (1 + z_{\text{abs}}) \left(1 + K_i \frac{\Delta\mu}{\mu} \right)$$

How to measure variation?

molecular hydrogen H₂

- electron-vibro-rotational transitions depend on reduced mass of molecule
- different dependence for different transitions
- distinguish cosmological redshift of a line from the shift caused by possible variation of μ

- $$\lambda_{\text{obs}} = \lambda_{\text{rest}} \times (1 + z_{\text{abs}}) \left(1 + K_i \frac{\Delta\mu}{\mu}\right)$$

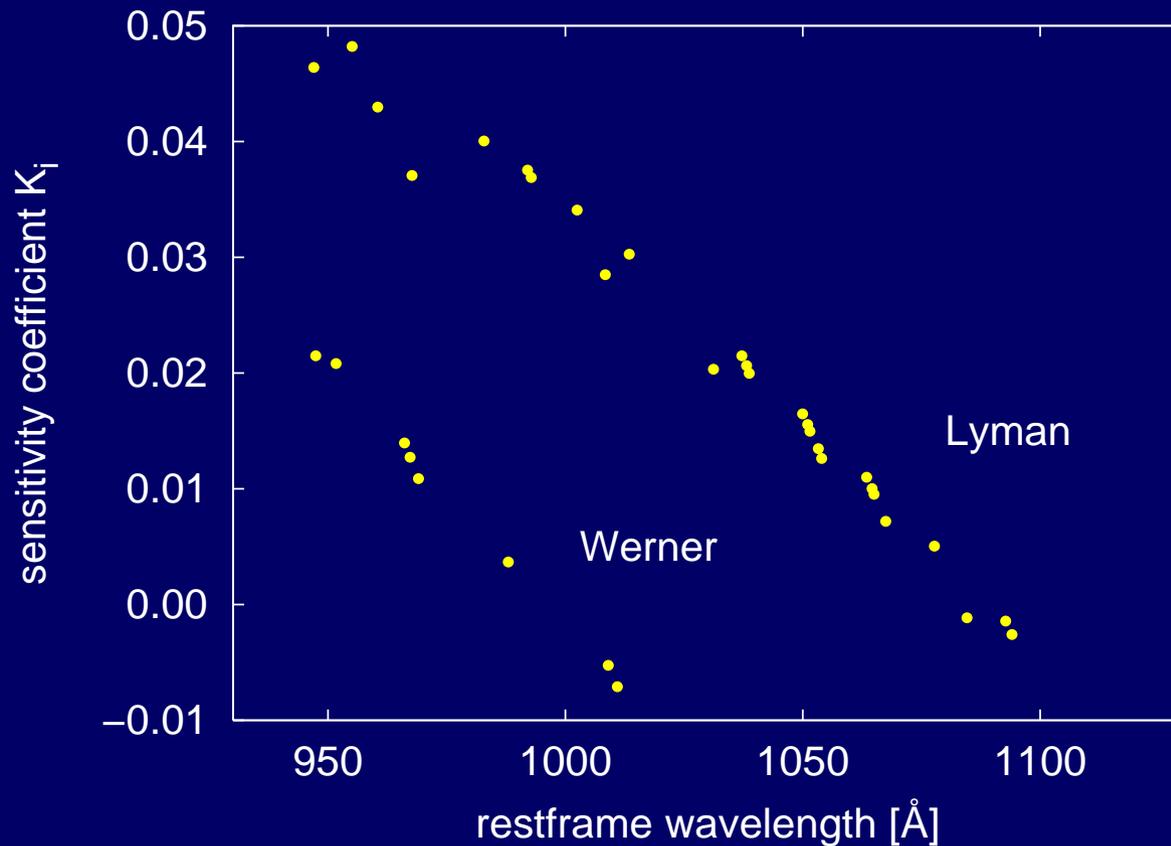
(Varshalovich & Levshakov 1993)

How to measure variation?

sensitivity coefficient $K_i = \frac{d \ln \lambda_i^0}{d \ln \mu}$

How to measure variation?

sensitivity coefficient $K_i = \frac{d \ln \lambda_i^0}{d \ln \mu}$



(Reinhold et al. 2006)

extragalactic H_2

- local observations with UV-satellite
COPERNICUS

extragalactic H₂

- local observations with UV-satellite
COPERNICUS
- the most abundant molecule in space

extragalactic H₂

- local observations with UV-satellite
COPERNICUS
- the most abundant molecule in space
- formation on dust grains

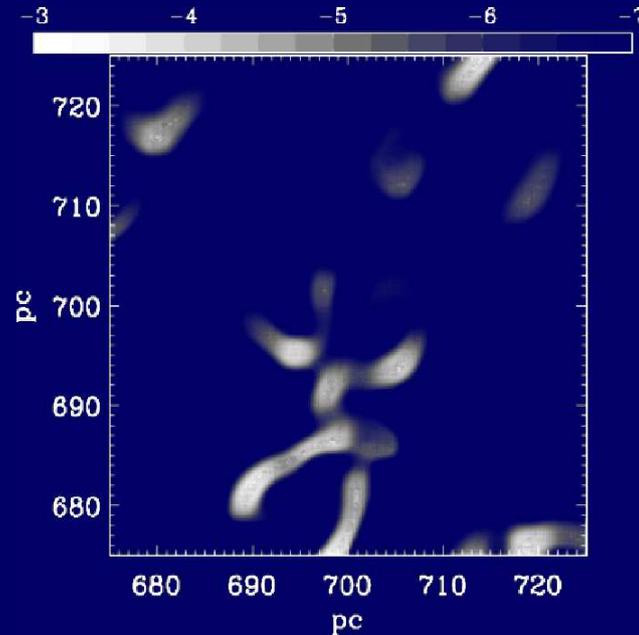
extragalactic H₂

- local observations with UV-satellite
COPERNICUS
- the most abundant molecule in space
- formation on dust grains
- shielding from UVB vs. dust obscuration

extragalactic H₂

- local observations with UV-satellite
COPERNICUS
- the most abundant molecule in space
- formation on dust grains
- shielding from UVB vs. dust obscuration
- transitions in UV (restframe) – redshifted into
visual band

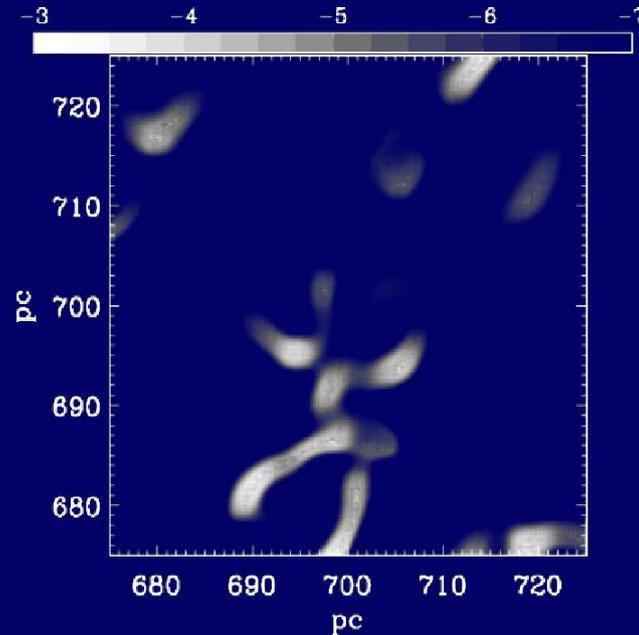
extragalactic H₂



- highly inhomogeneous, clumpy distribution [1]

[1] H.Hirashita, A.Ferrara, K.Wada, P.Richter,P.2003, MNRAS, 341, L18

extragalactic H₂

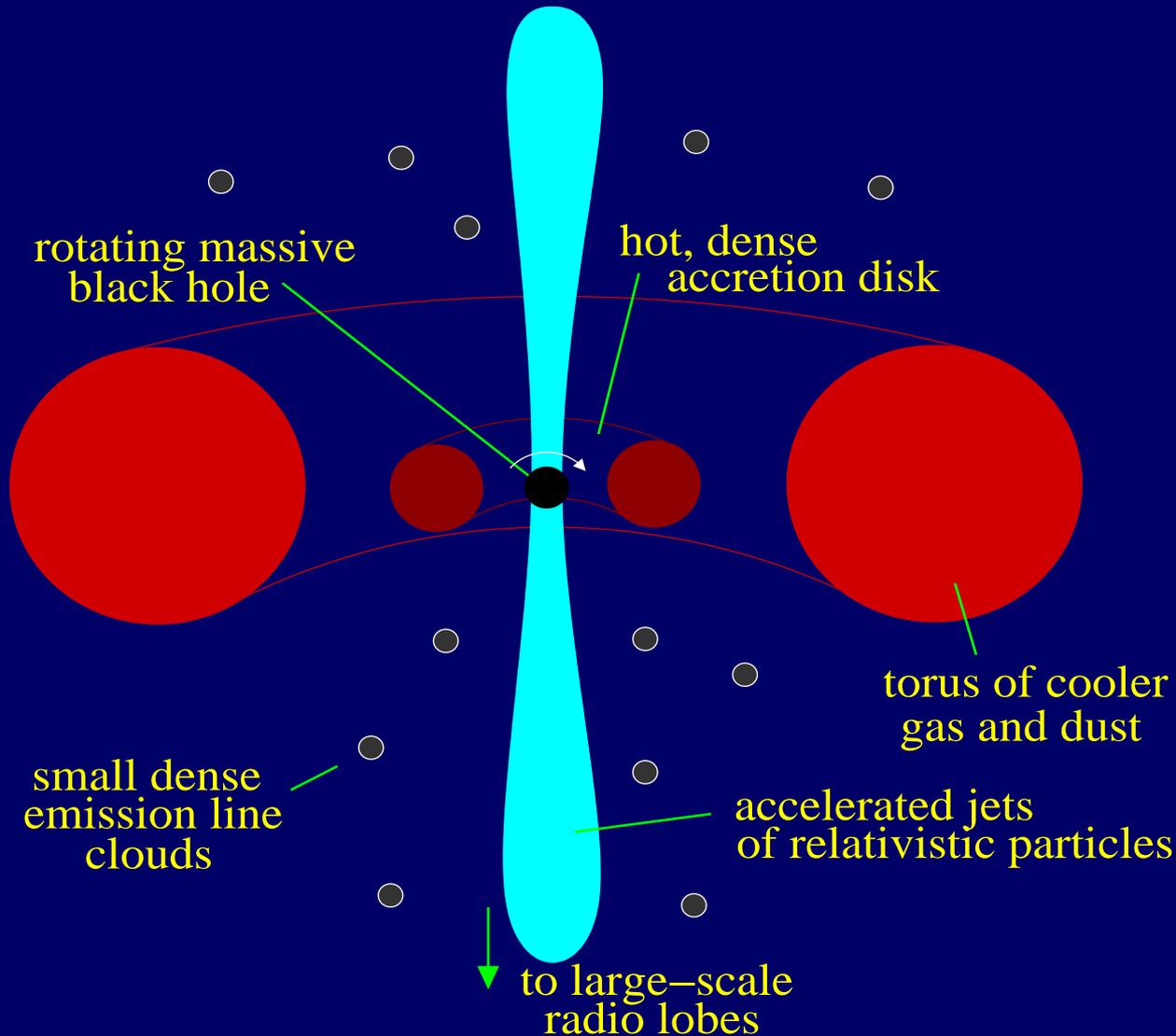


- highly inhomogeneous, clumpy distribution [1]
- observable only in dense systems

[1] H.Hirashita, A.Ferrara, K.Wada, P.Richter, P.2003, MNRAS, 341, L18

Quasar absorption line spectroscopy

Quasar absorption line spectroscopy



Quasar absorption line spectroscopy

Quasar absorption line spectroscopy

- cosmological redshift z due to expansion of space.

Quasar absorption line spectroscopy

- cosmological redshift z due to expansion of space.

$$\lambda_{\text{obs}} = \lambda_{\text{rest}} \times (1 + z_{\text{abs}})$$

Quasar absorption line spectroscopy

- cosmological redshift z due to expansion of space.

$$\lambda_{\text{obs}} = \lambda_{\text{rest}} \times (1 + z_{\text{abs}})$$

- Quasars as bright distant background sources against which intervening gas can be observed.

Quasar absorption line spectroscopy

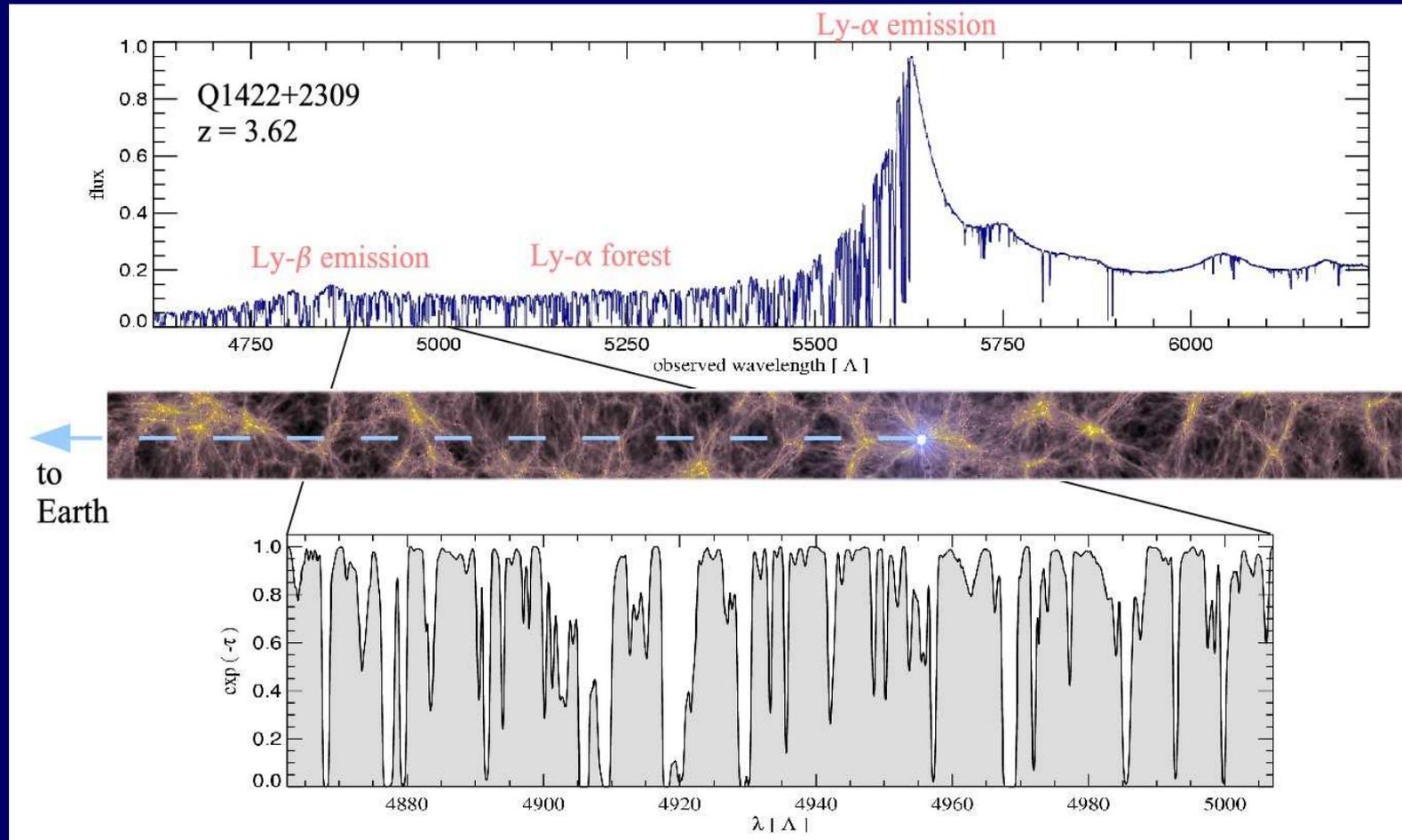
- cosmological redshift z due to expansion of space.

$$\lambda_{\text{obs}} = \lambda_{\text{rest}} \times (1 + z_{\text{abs}})$$

- Quasars as bright distant background sources against which intervening gas can be observed.

e.g., the Ly α transition at $\lambda_{\text{rest}} = 1215.67 \text{ \AA}$

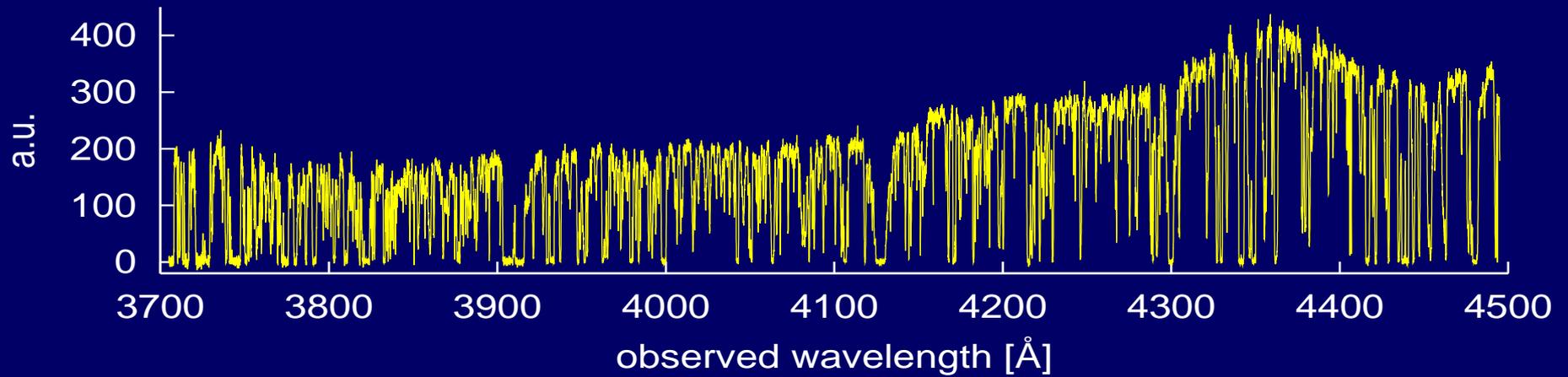
Quasar absorption line spectroscopy



(Springel et. al 2006)

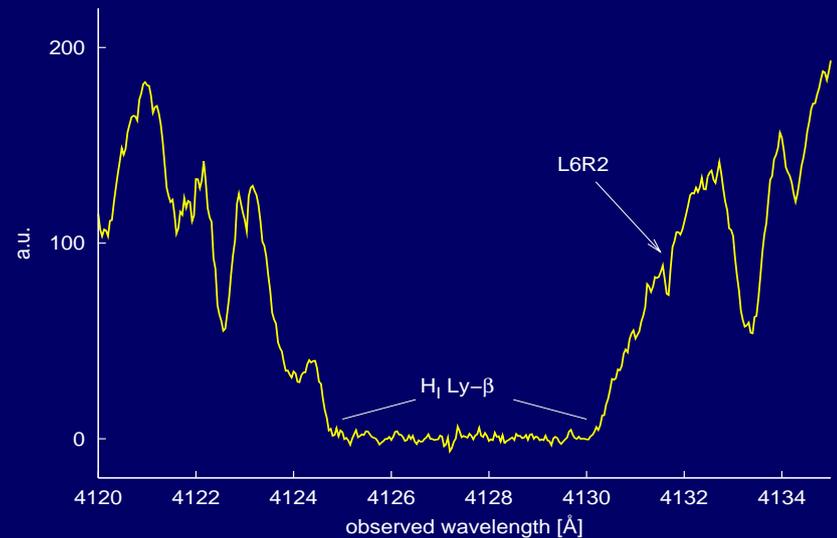
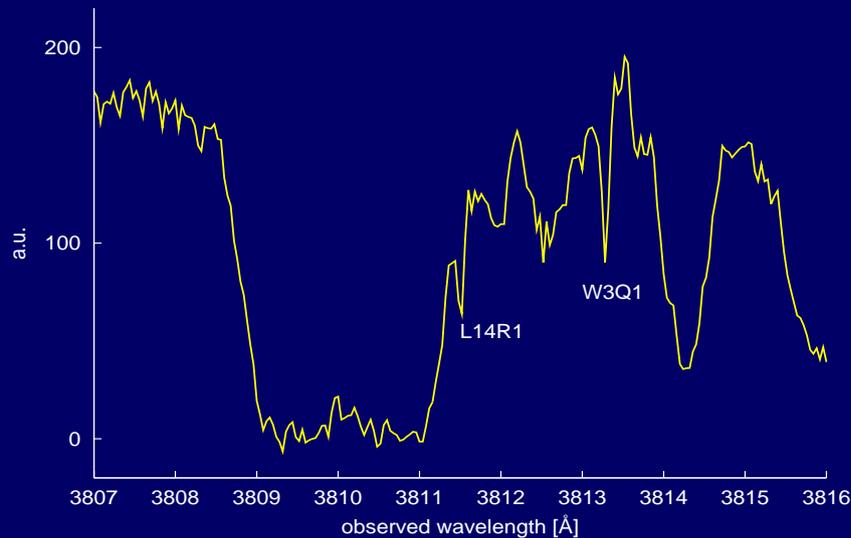
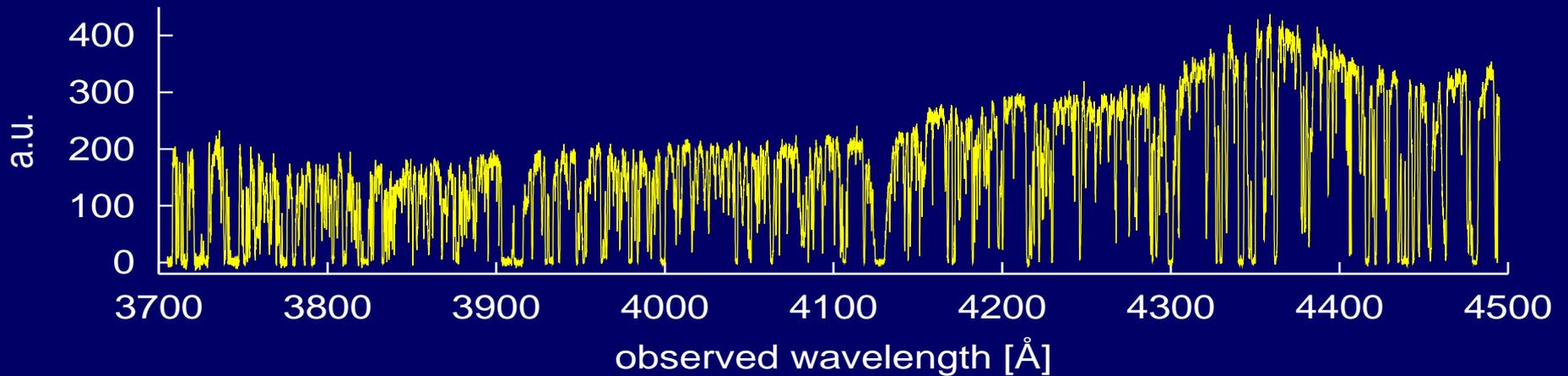
Quasar absorption line spectra - probing the universe

Q 0347-383



Quasar absorption line spectra - probing the universe

Q 0347-383



Quasar absorption line spectra - probing the universe

Quasar absorption line spectra - probing the universe

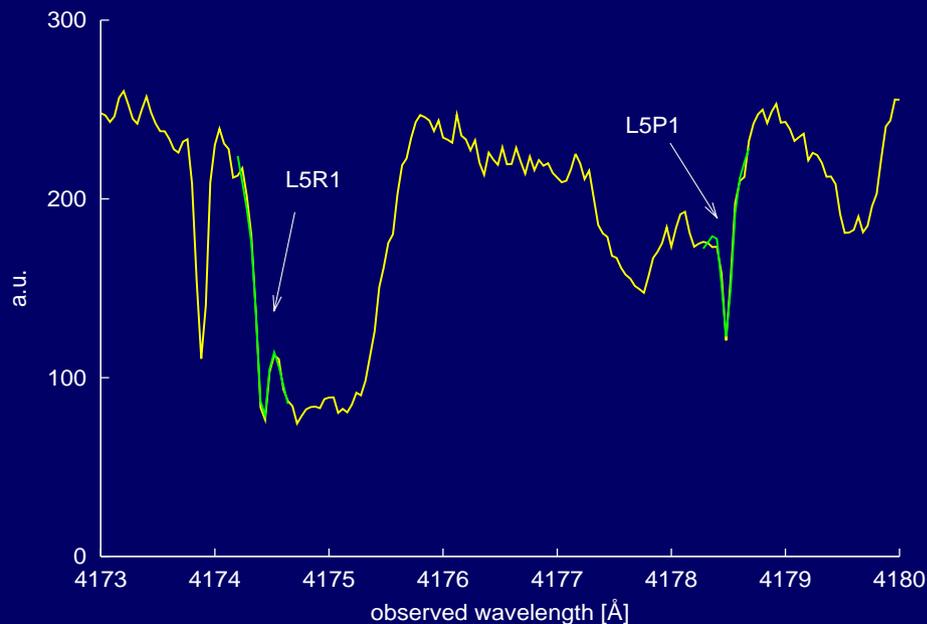
- H₂ lines in DLA systems

Quasar absorption line spectra - probing the universe

- H₂ lines in DLA systems
- contaminated continuum

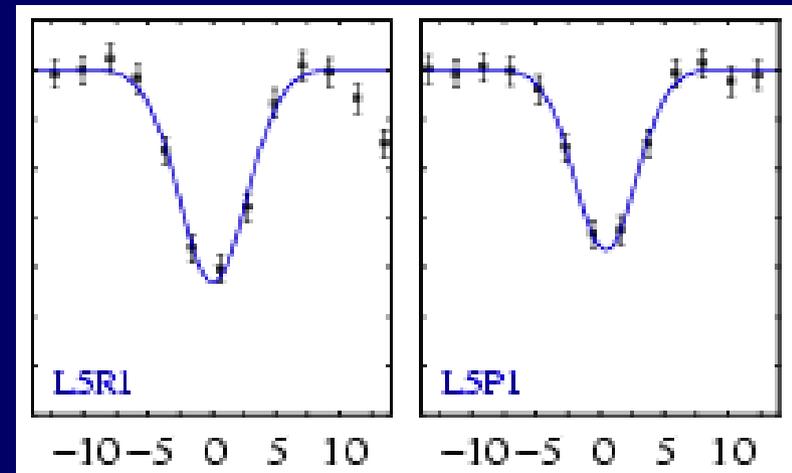
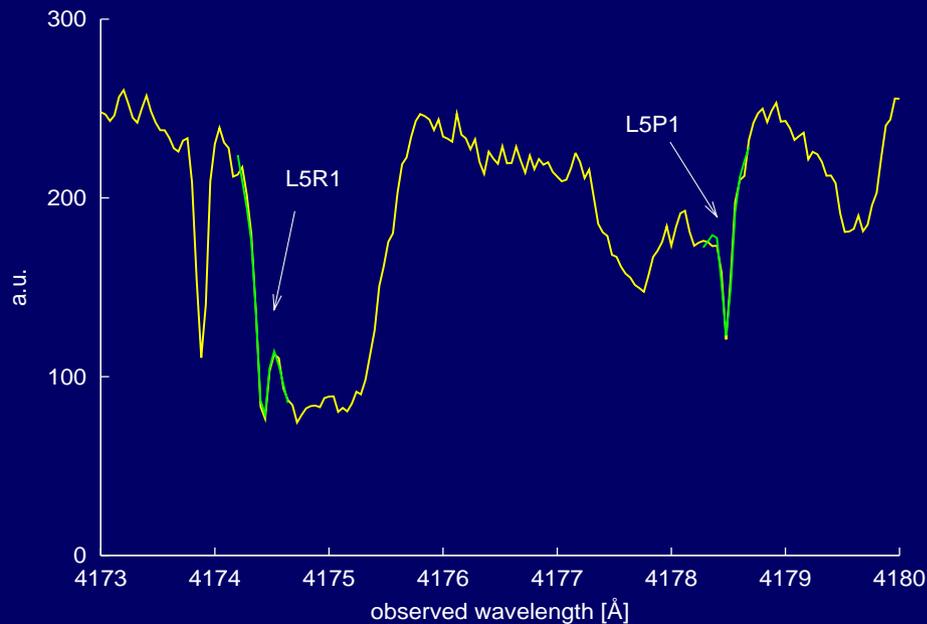
Quasar absorption line spectra - probing the universe

- H₂ lines in DLA systems
- contaminated continuum



Quasar absorption line spectra - probing the universe

- H₂ lines in DLA systems
- contaminated continuum



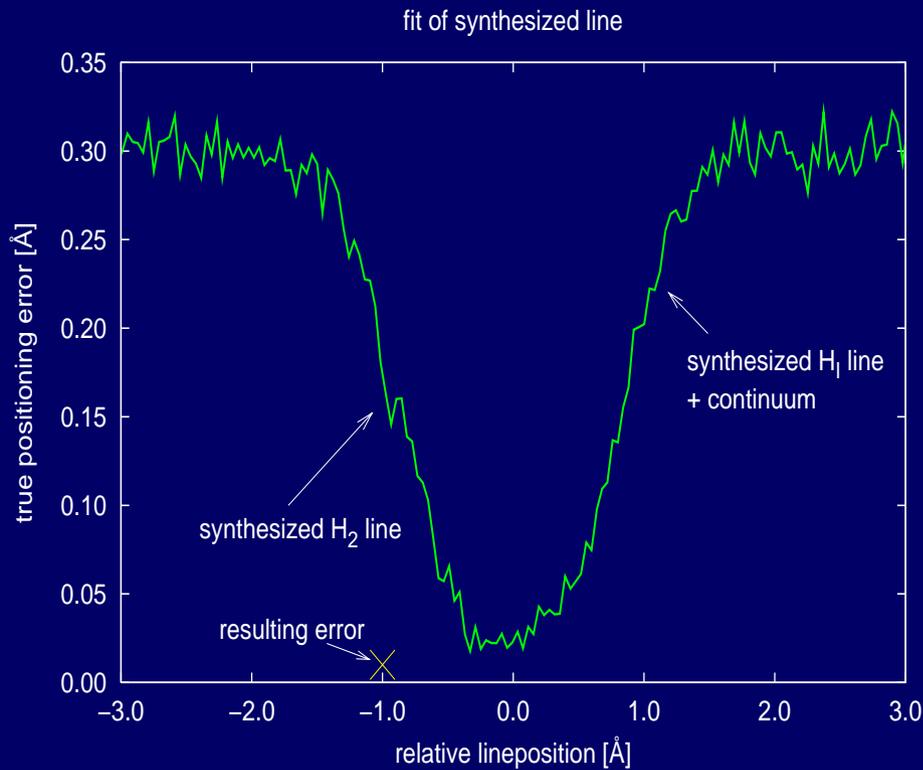
(Ivanchik et al. 2005)

Quasar absorption line spectra - probing the universe

simulated fits to estimate accuracy

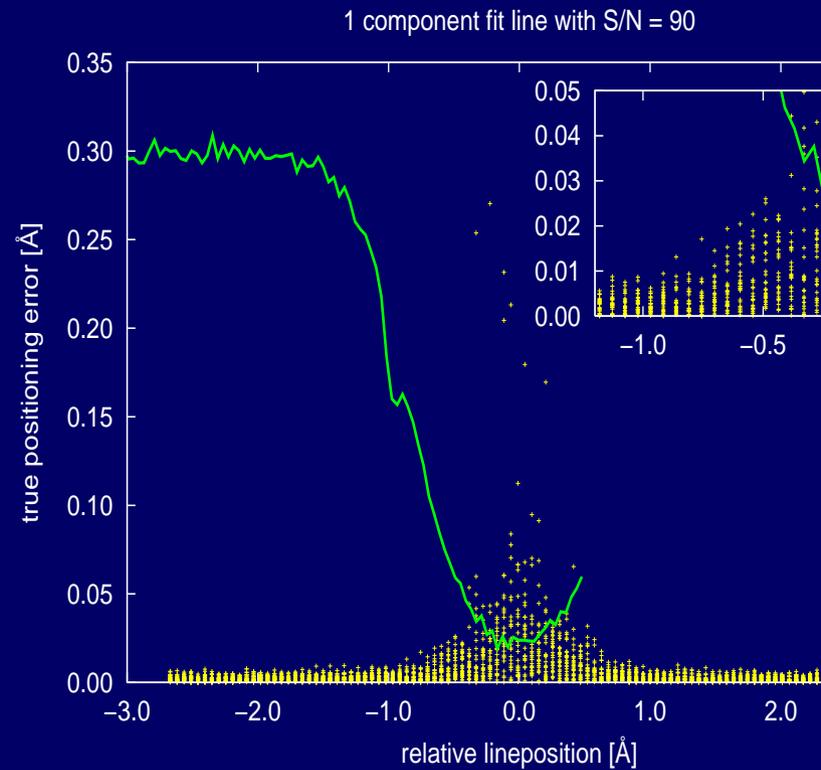
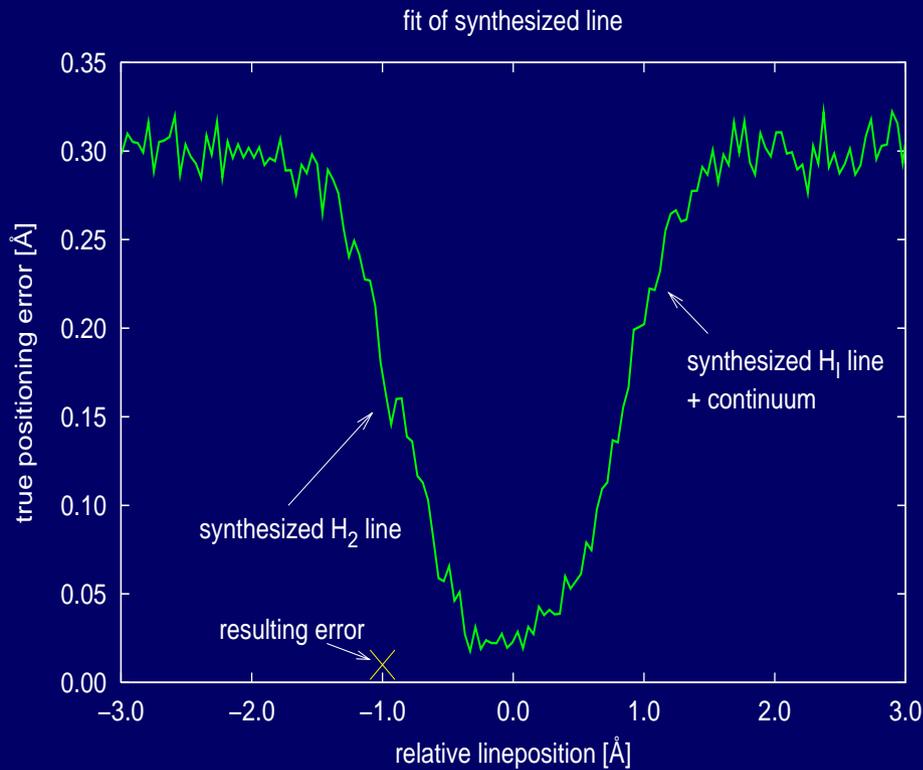
Quasar absorption line spectra - probing the universe

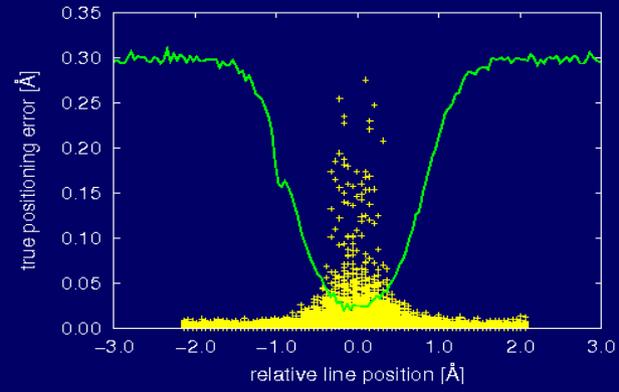
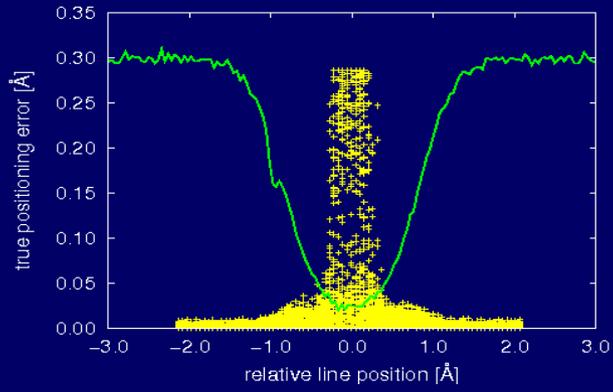
simulated fits to estimate accuracy

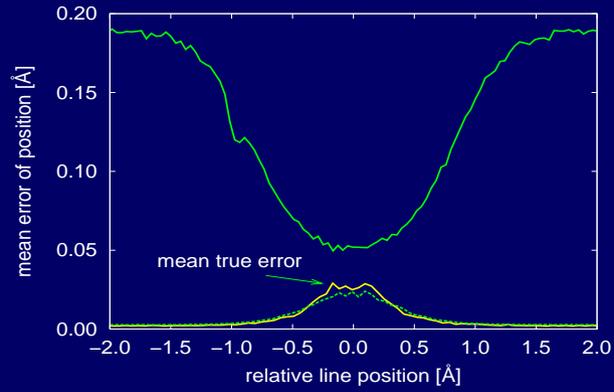
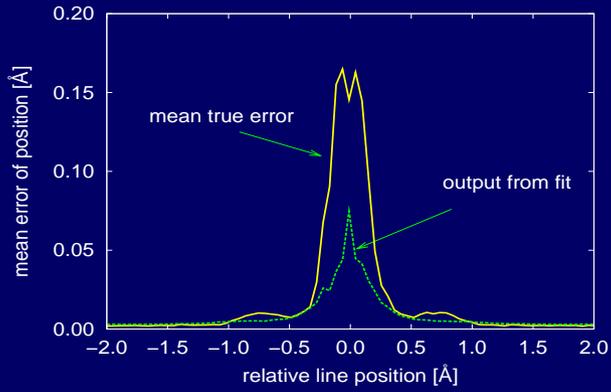
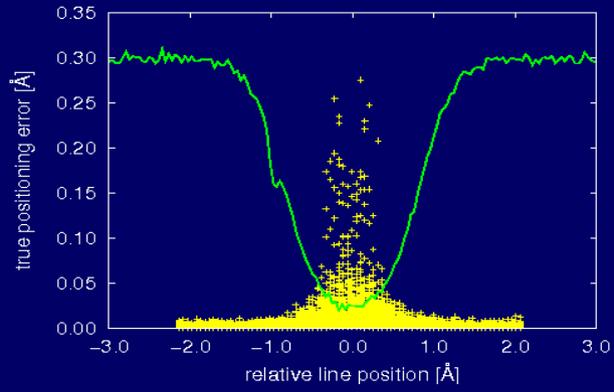
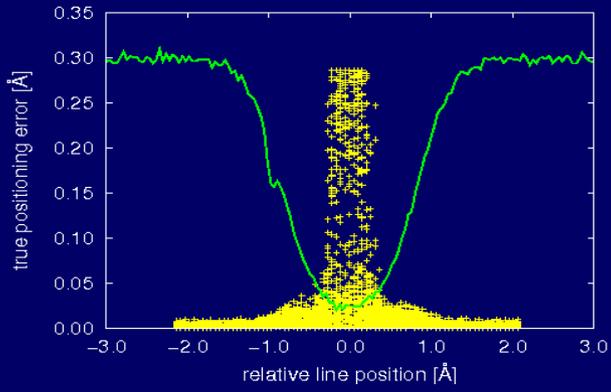


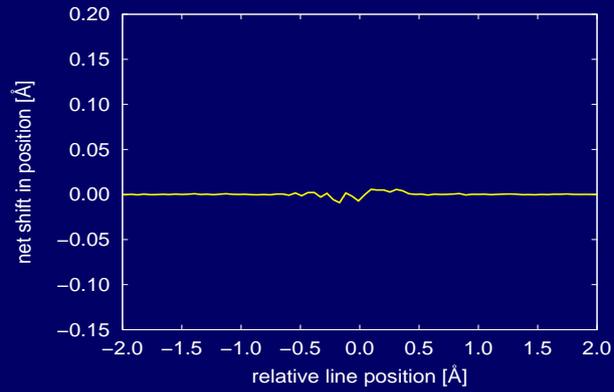
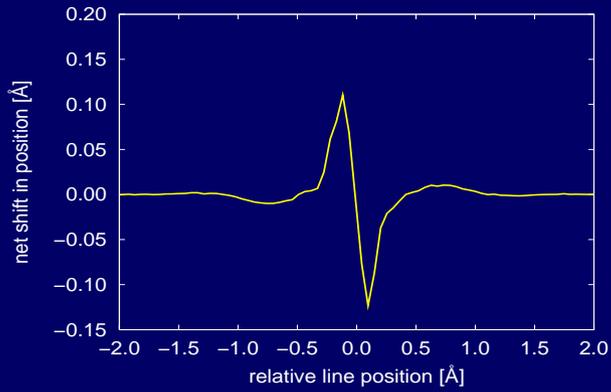
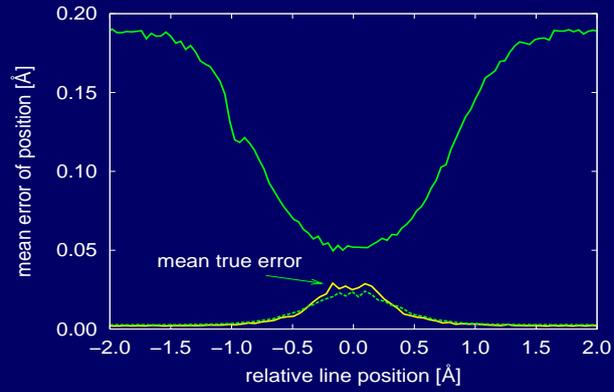
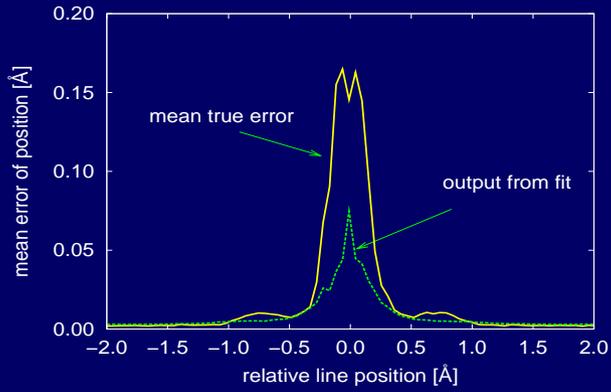
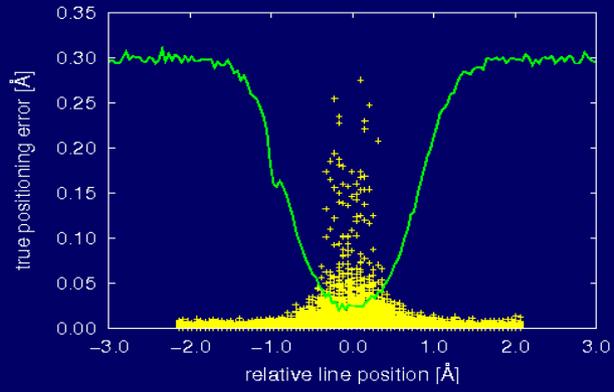
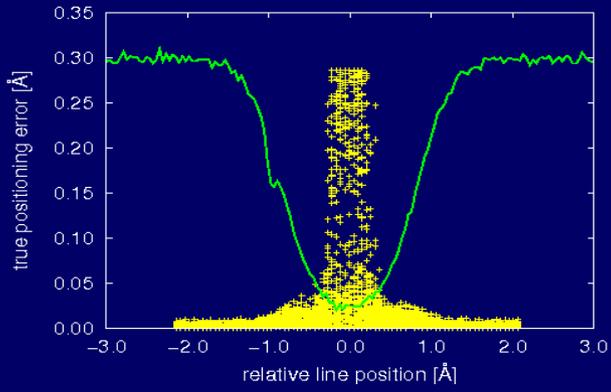
Quasar absorption line spectra - probing the universe

simulated fits to estimate accuracy



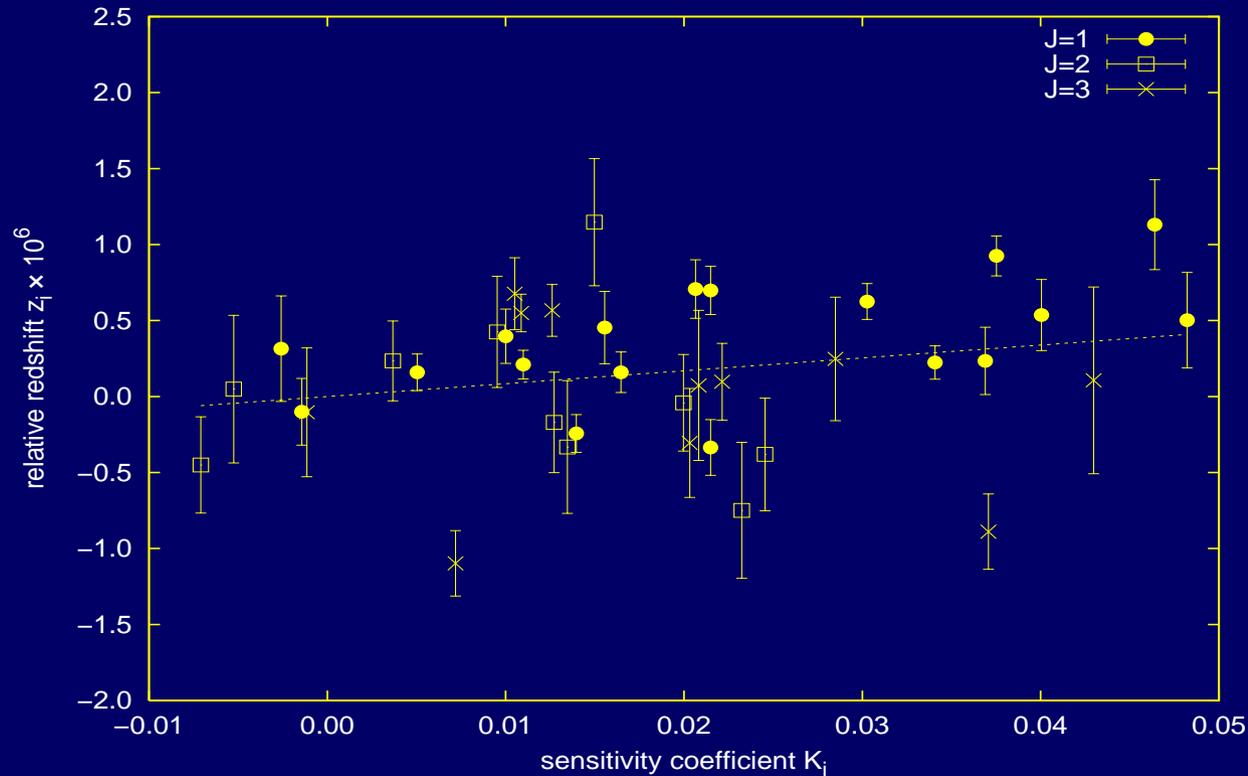






$$b = (1 + z_{\text{abs}}) \times \frac{\Delta\mu}{\mu}$$

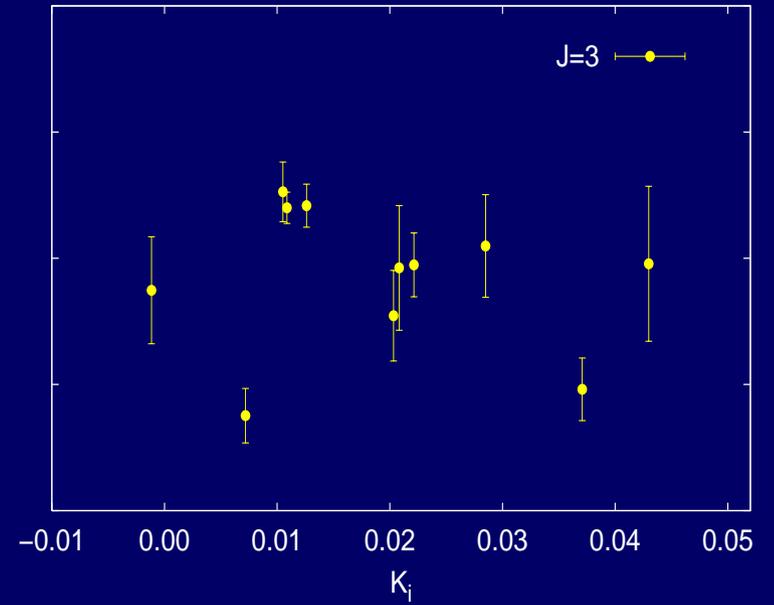
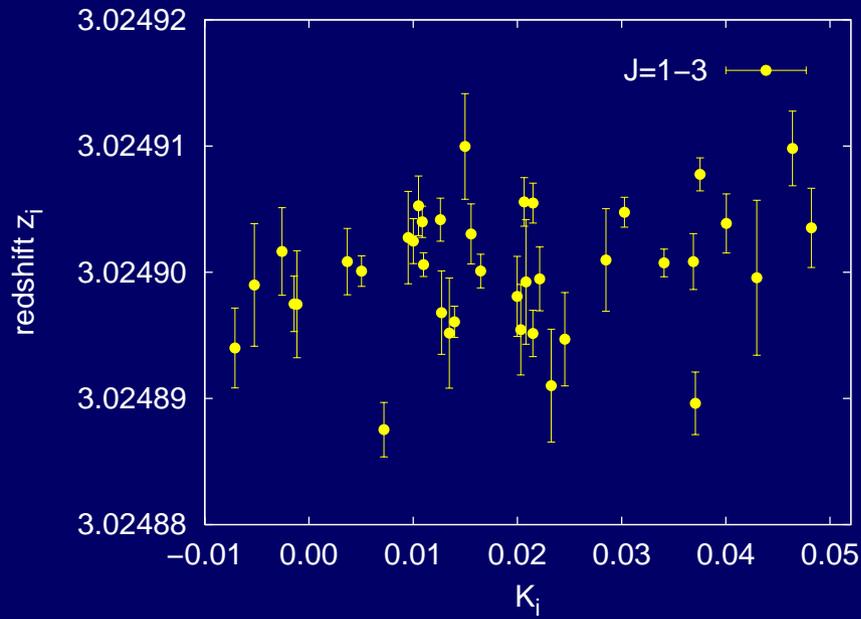
$$b = (1 + z_{\text{abs}}) \times \frac{\Delta\mu}{\mu}$$



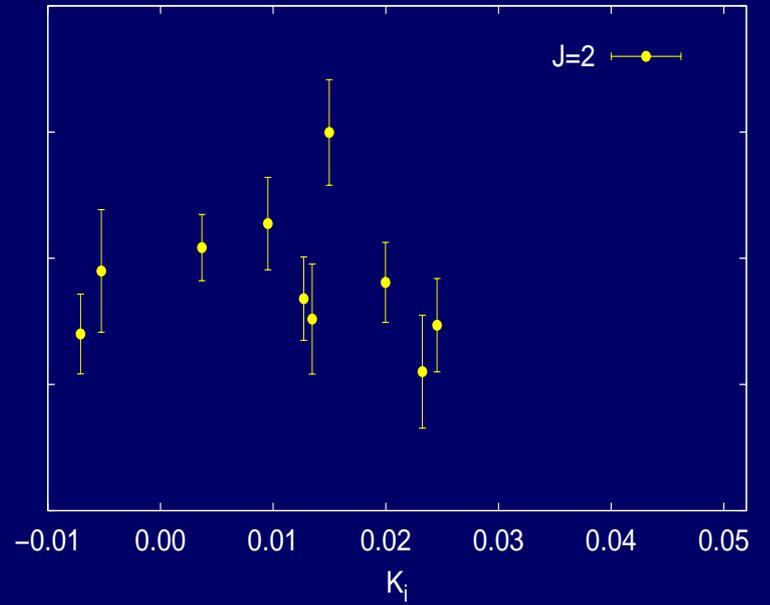
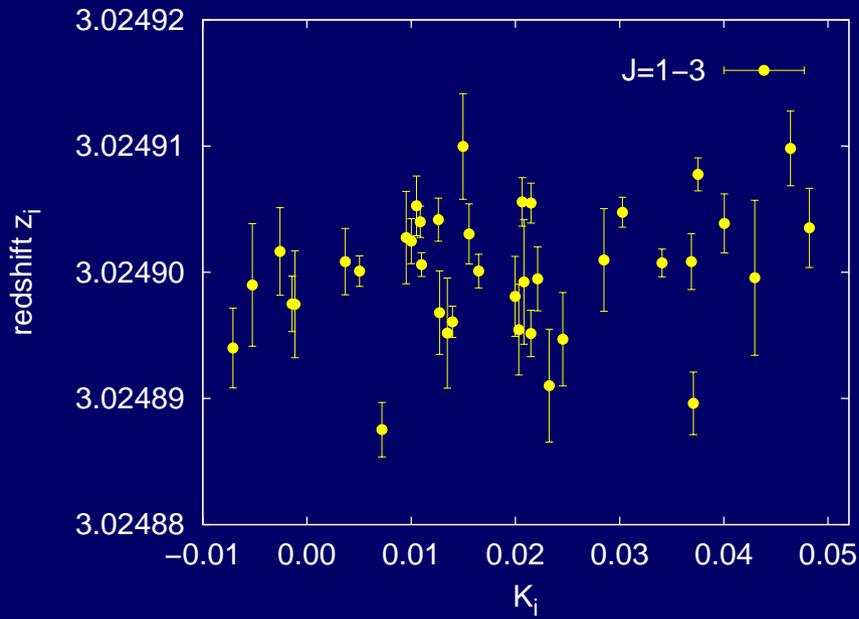
corresponding to $\frac{\Delta\mu}{\mu} = 2.1 \pm 1.4 \times 10^{-5}$

(Reinhold et al. 2006: $\frac{\Delta\mu}{\mu} = 2.0 \pm 0.6 \times 10^{-5}$)

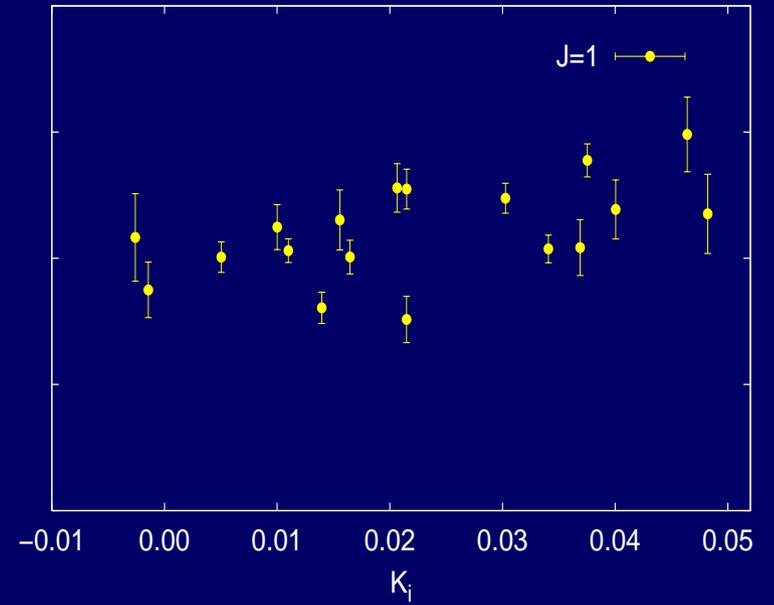
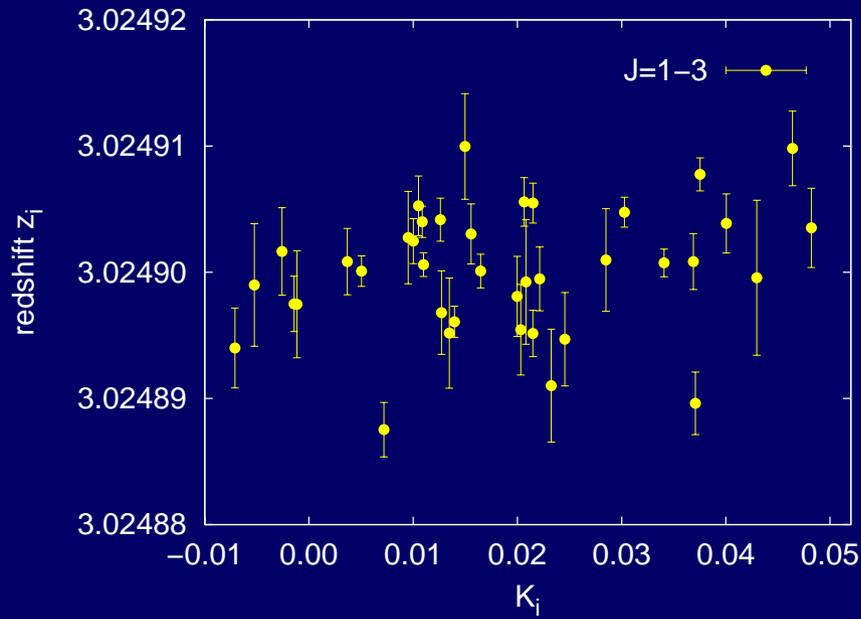
$$b = (1 + z_{\text{abs}}) \times \frac{\Delta\mu}{\mu}$$



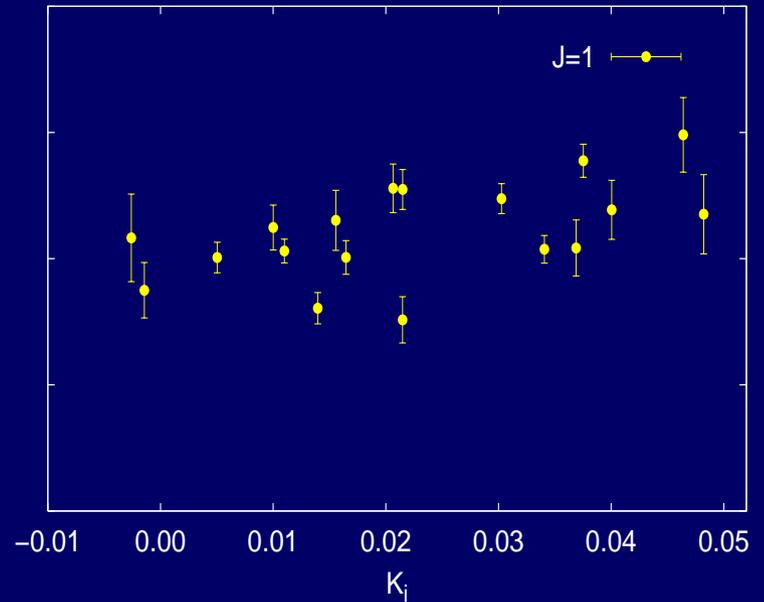
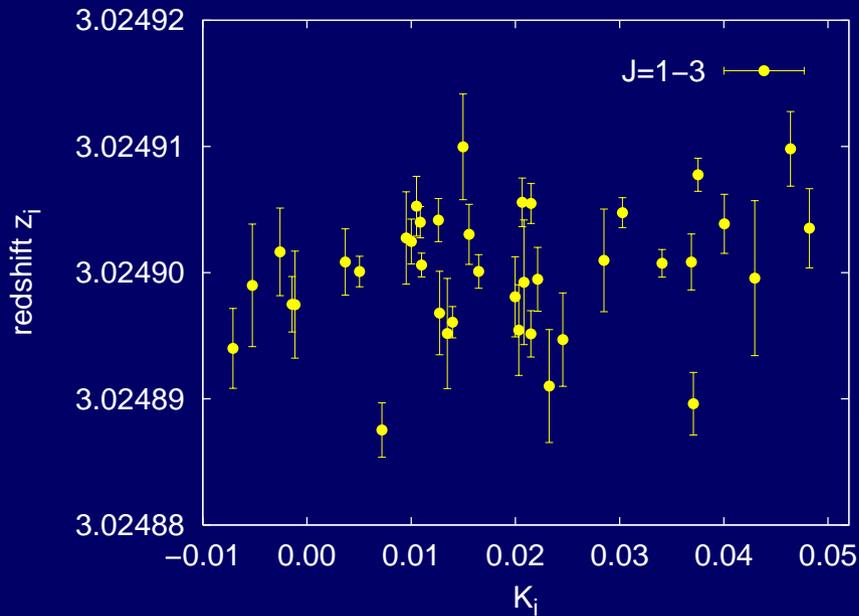
$$b = (1 + z_{\text{abs}}) \times \frac{\Delta\mu}{\mu}$$



$$b = (1 + z_{\text{abs}}) \times \frac{\Delta\mu}{\mu}$$

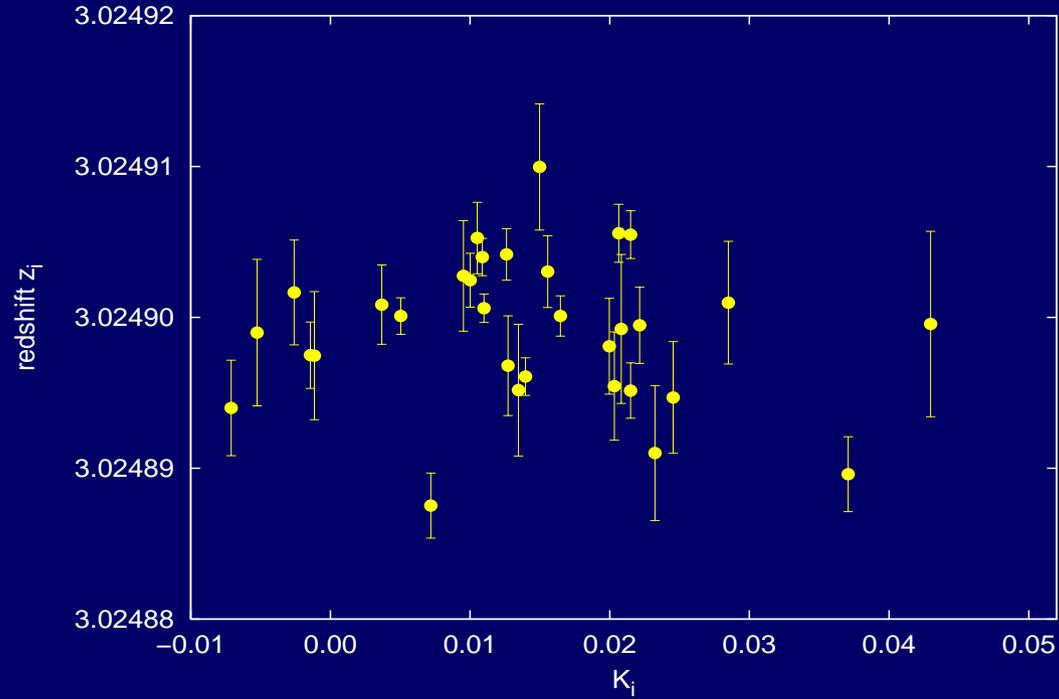


$$b = (1 + z_{\text{abs}}) \times \frac{\Delta\mu}{\mu}$$

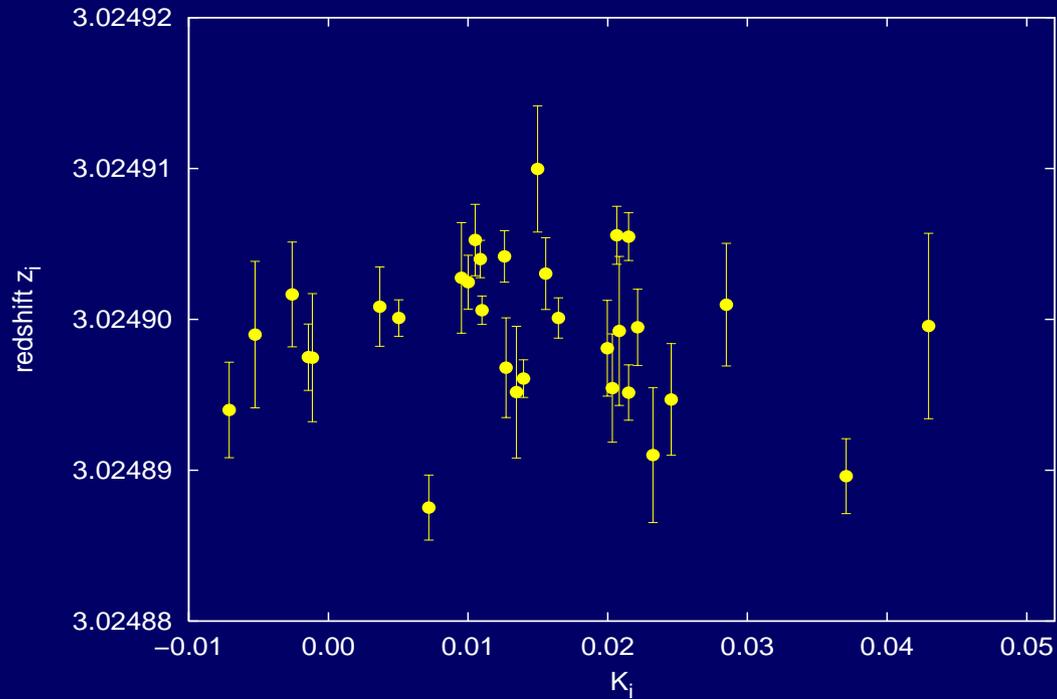


Merely transitions with high vibrational quantum numbers in the first rotational level contribute to a positive result

News or noise?

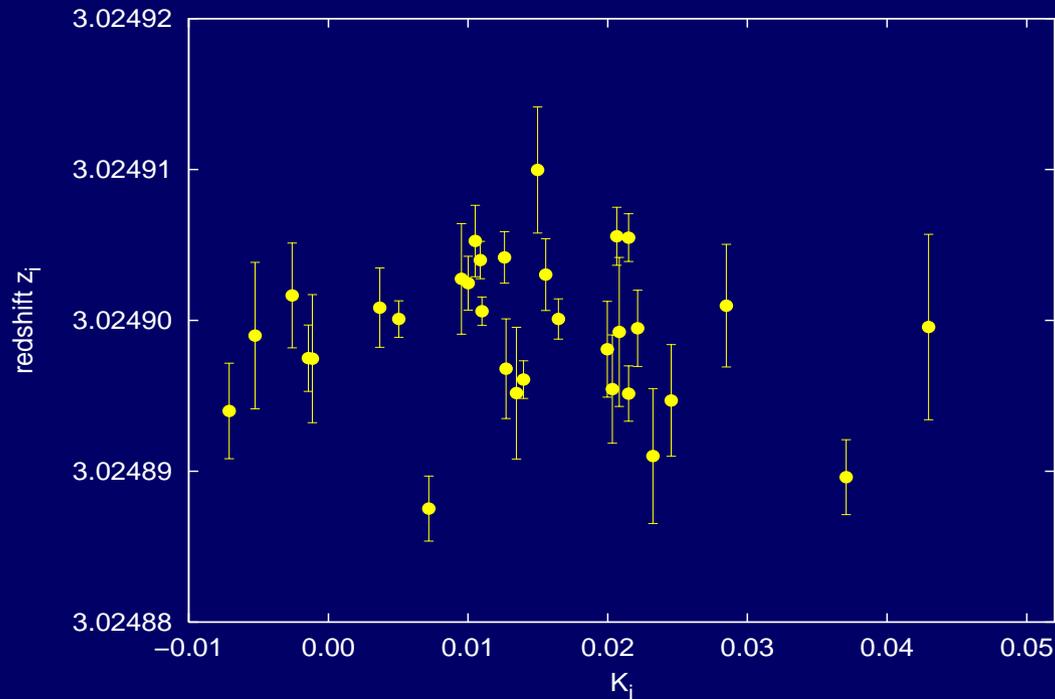


News or noise?



no detectable correlation in a 85% subset

News or noise?



no detectable correlation in a 85% subset

$$|\Delta\mu/\mu| \leq 4.9 \times 10^{-5} \text{ over the period of } \approx 11.5 \text{ Gyr}$$

Outlook

Outlook

- line lists of required accuracy just available
⇒ increased need for high resolution

Outlook

- line lists of required accuracy just available
⇒ increased need for high resolution
- in general attach more importance to data reduction

Outlook

- line lists of required accuracy just available
⇒ increased need for high resolution
- in general attach more importance to data reduction
- search for more quasar spectra with DLA and H₂ signatures

Outlook

- line lists of required accuracy just available
⇒ increased need for high resolution
- in general attach more importance to data reduction
- search for more quasar spectra with DLA and H₂ signatures
- further simulations of detectability of variation

Outlook

- line lists of required accuracy just available
⇒ increased need for high resolution
- in general attach more importance to data reduction
- search for more quasar spectra with DLA and H₂ signatures
- further simulations of detectability of variation
- better understanding of the nature of DLAs

The Ratio of Proton and Electron Masses

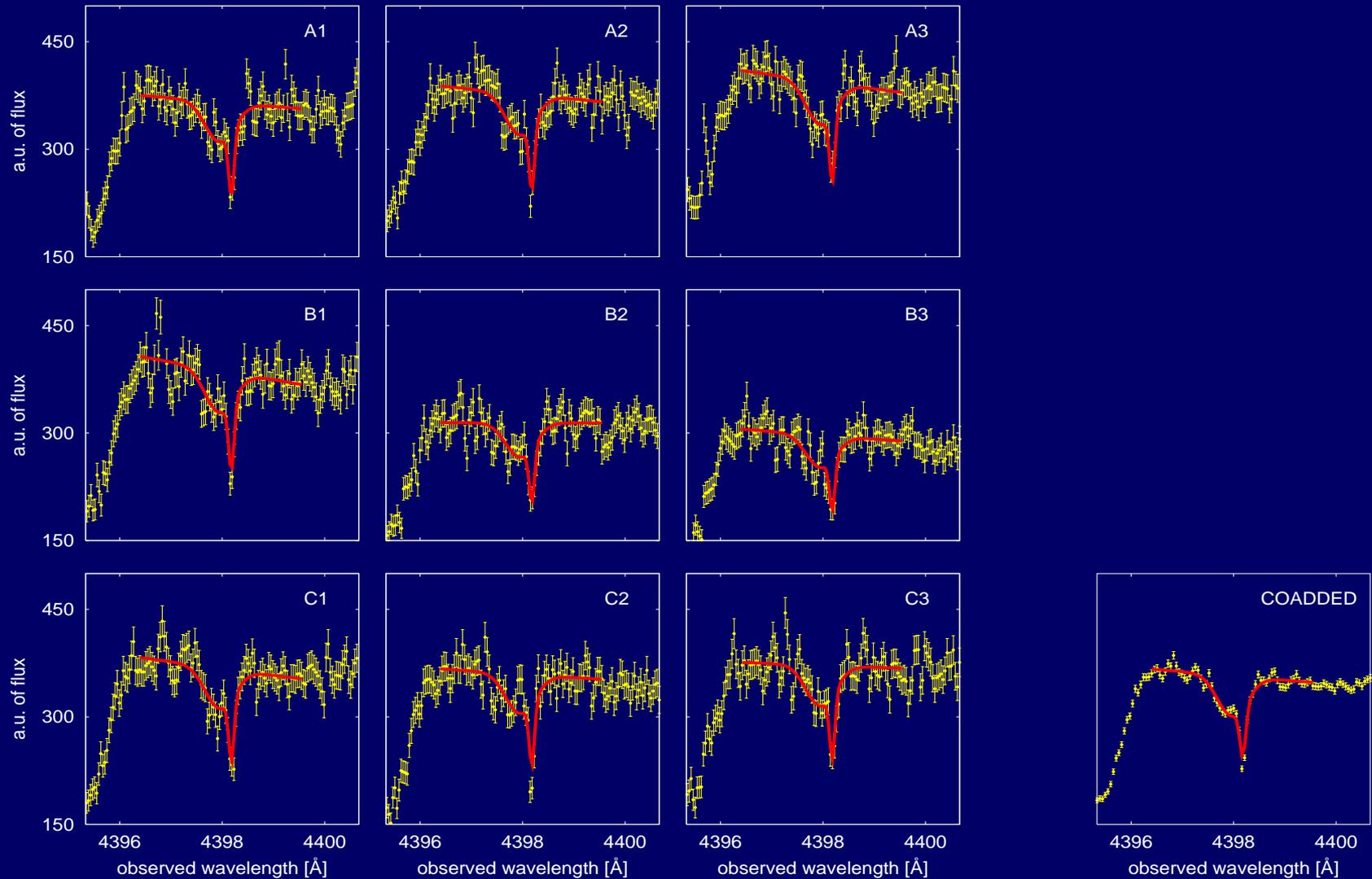
FRIEDRICH LENZ

Düsseldorf, Germany

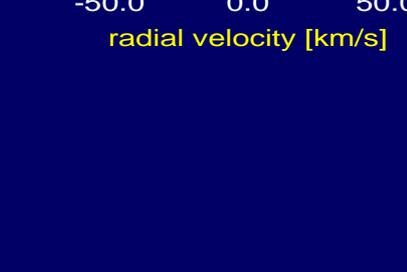
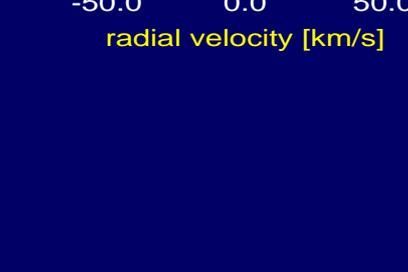
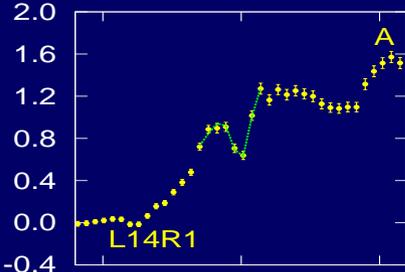
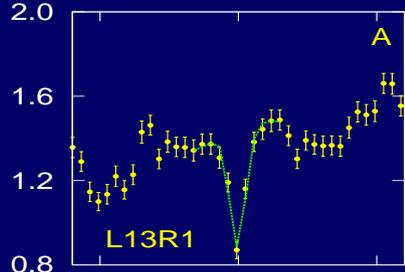
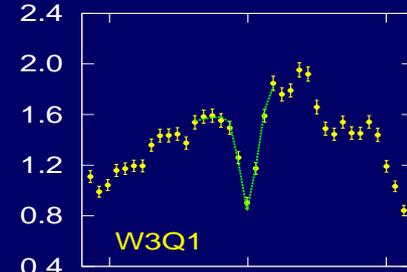
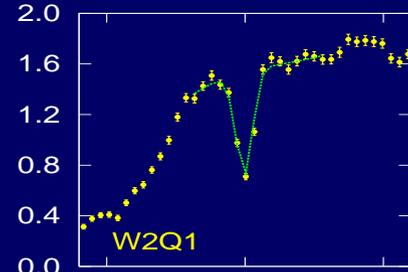
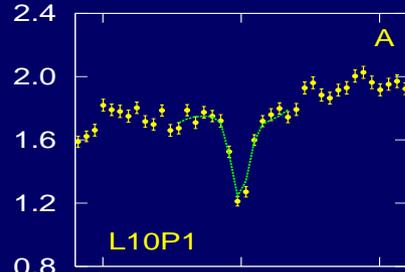
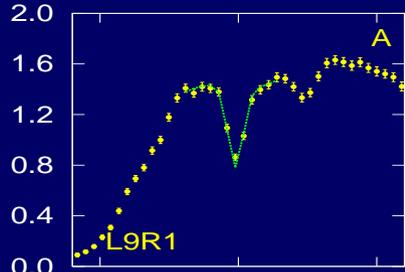
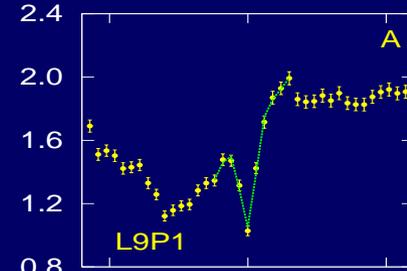
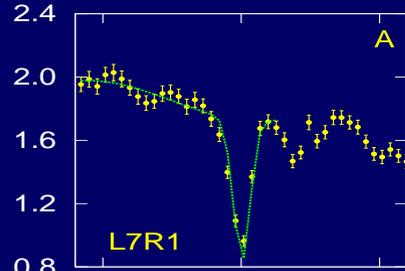
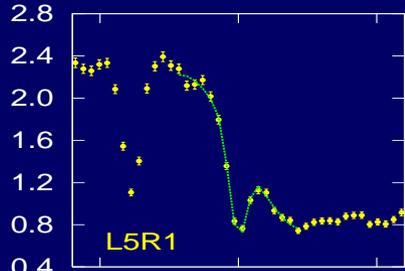
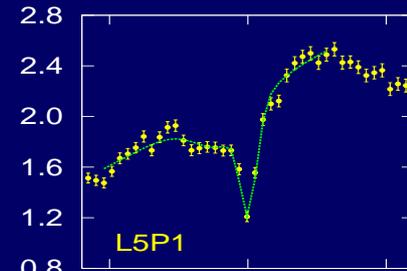
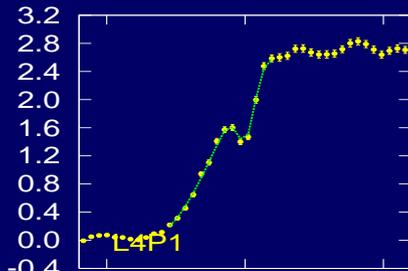
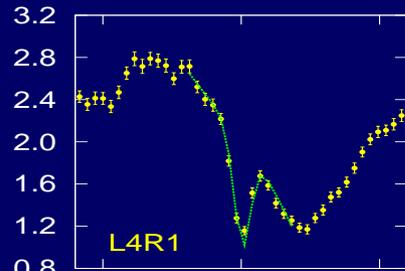
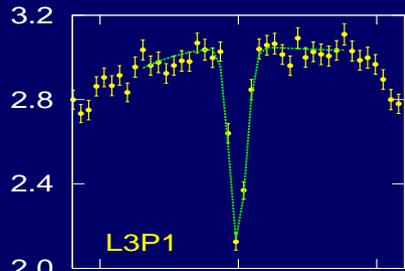
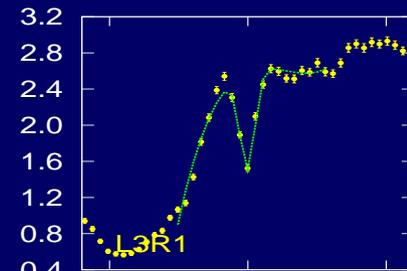
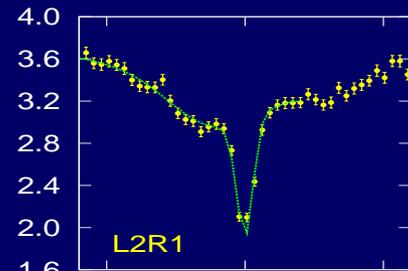
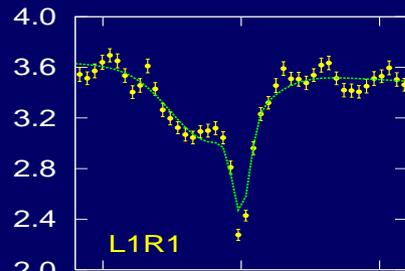
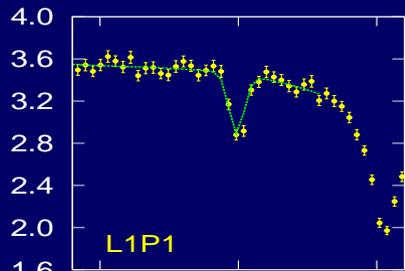
(Received April 5, 1951)

THE most exact value at present¹ for the ratio of proton to electron mass is 1836.12 ± 0.05 . It may be of interest to note that this number coincides with $6\pi^5 = 1836.12$.

¹ Sommer, Thomas, and Hipple, *Phys. Rev.* **80**, 487 (1950).



Nine separately observed spectra with errorbars
and exemplary fit of L1R1.



radial velocity [km/s]

radial velocity [km/s]

radial velocity [km/s]

radial velocity [km/s]