

# Lecture 7

## The Hubble law

Schutz, Chap 12  
Vittorio, Chaps 1-2

### COSMOLOGICAL REDSHIFT

NOTE: every problem involving light propagation implies

- (1) something that emits a photon
- (2) the photon propagating on a null geodesic
- (3) the observer that receives the photon

In a FRW massive objects corresponding to emitters and receivers ("galaxies" in both cases) travel along time-like geodesics and have constant positions in comoving coordinates

$\Rightarrow L_{\text{los}}$  is constant

Let's have a photon propagating along the line of sight.

WE EXPLICIT C HERE

$$ds^2 = -c^2 dt^2 + a^2 dL_{\text{los}}^2 = 0 \Rightarrow \frac{cdt}{a(t)} = dL_{\text{los}}$$

The emitting galaxies gives two photons at  $t_{\text{em}}$  and  $t_{\text{em}} + \Delta t_{\text{em}}$ , observed at  $t_{\text{obs}}$  and  $t_{\text{obs}} + \Delta t_{\text{obs}}$

$$\int_{t_{\text{em}}}^{t_{\text{obs}}} \frac{cdt}{a(t)} = \int_{\text{GAL em}}^{\text{GAL obs}} dL_{\text{los}} = \int_{t_{\text{em}} + \Delta t_{\text{em}}}^{t_{\text{obs}} + \Delta t_{\text{obs}}} \frac{cdt}{a(t)}$$

This implies that:

$$\int_{t_{\text{em}}}^{t_{\text{em}} + \Delta t_{\text{em}}} \frac{cdt}{a(t)} \approx \frac{c \Delta t_{\text{em}}}{a(t_{\text{em}})} = \int_{t_{\text{obs}}}^{t_{\text{obs}} + \Delta t_{\text{obs}}} \frac{cdt}{a(t)} \approx c \Delta t_{\text{obs}}$$

given that  $a(t_{\text{obs}}) = 1$ . Given that frequencies are  $\propto t^{-1}$  and  $\lambda \propto v^{-1}$

$$\Rightarrow \frac{1}{a(t_{\text{em}})} = \frac{\lambda_{\text{obs}}}{\lambda_{\text{em}}} = 1 + \frac{\lambda_{\text{obs}} - \lambda_{\text{em}}}{\lambda_{\text{em}}} = 1 + z_{\text{em}}$$

The Schutz textbook gives a proof that is similar to that of gravitational redshift:

the metric  $ds^2 = -c^2 dt^2 + a^2(t) R_0^2 dX^2$  does not depend on  $X \Rightarrow P_X$  is conserved.

Using  $P^\alpha P_\alpha = 0$  we get:

$$0 = -P_0^2 + \frac{1}{R_0^2 \alpha^2(t)} P^2 x \Rightarrow P_0 \propto \frac{1}{\alpha(t)}$$

The photon energy in the frame of the observer, for which  $\vec{U} = (1, 0, 0, 0)$

$$\text{is : } E = -\vec{U} \cdot \vec{p} = -P_0 = h v_{\text{obs}} \Rightarrow h v_{\text{obs}} \propto \frac{1}{\alpha(t)} \Rightarrow \lambda_{\text{obs}} \propto \alpha$$

$$\Rightarrow \frac{\lambda_{\text{obs}}}{\lambda_{\text{em}}} = \frac{1}{\alpha(t_{\text{em}})} \quad \text{again}$$

But the basic interpretation is :  $\lambda$  expands with the Universe, as all lengths do.

## PROPER DISTANCE

Keep  $t = \text{const}$ ,  $\alpha(t) = \text{const}$ , the distance between two galaxies is:

$$d_p = \int_0^r \alpha(t) \frac{dr'}{\sqrt{1 - \frac{k}{R_0^2} r'^2}} = \alpha(t) f(r) = \alpha(t) R_0 \chi(r)$$

$$= \alpha(t) L_{\text{los}} \quad \text{(different definitions in different textbooks)}$$

This is of course NOT measurable

## COMOVING DISTANCE already discussed

$$d_c = f(r) = R_0 \chi(r) = L_{\text{los}}, d_p = \alpha(t) d_c$$

## HUBBLE PARAMETER

Let's call "velocity"  $v$  the rate at which the proper distance changes with time:

$$v = \frac{d}{dt} d_p = \dot{\alpha}(t) d_c = \frac{\dot{\alpha}}{\alpha} d_p$$

where the dot denotes time derivative.

Let's define the Hubble parameter as :

$$H(t) = \frac{\dot{\alpha}}{\alpha}$$

$$\text{then : } v = H d_p$$

but THIS IS NOT THE HUBBLE LAW in a strict sense because  $d_p$  is not measurable

To understand observable distances we adopt an OPERATIONAL APPROACH

We measure distances through:

- standard candles
- standard rulers

$$\rightarrow \text{Standard candle: } f = \frac{L}{4\pi d^2} = \frac{L}{S}$$

$f$  is the measured flux ( $T^{oi}$ )

$L$  the candle luminosity

$S$  the area of a surface at  $r=d$

$$S = 4\pi R_0^2 r^2 \quad \text{where } r \text{ is the coordinate distance}$$

Here we use  $R_0$  because the photons are observed at  $t=t_0$ .

$T^{oi}$  is affected twice by a transformation:

$$L \text{ is energy / time, } E = h\nu \propto \frac{1}{\alpha}, dt \propto \alpha$$

$$\Rightarrow L \propto \alpha^{-2}, L_{\text{obs}} = L_{\text{em}} \alpha^2$$

$$\Rightarrow f = \frac{L_{\text{em}}}{4\pi R_0^2 r^2} \alpha^2 = \frac{L_{\text{em}}}{4\pi d_L^2}$$

### LUMINOSITY DISTANCE

$$d_L = \frac{1}{\alpha} R_0 r = (1+z) R_0 r = (1+z) L_{\text{sky}}$$

$$\rightarrow \text{Standard ruler: } \theta = \frac{D}{d} \quad (\text{in radians})$$

$\theta$  is the angular extension of an object  
 $D$  is its physical size

Expansion is isotropic and keeps the angles constant, so we can compute this at fixed time

We use the time at emission, because  $D$  is defined at that time.

We use the computation of a great circle to obtain:

$$D = \theta \alpha R_0 r \Rightarrow \theta = \frac{D}{\alpha R_0 r} = \frac{D}{d_L}$$

### DIAMETER DISTANCE

$$d_D = \alpha R_0 r = \frac{1}{1+z} R_0 r = \frac{1}{1+z} L_{\text{sky}}$$

It follows that:  $d_L(1+z)^2 = d_L$

The relation between  $d_L$  and  $d_C$  is:

$$d_L = \begin{cases} (1+z) R_0 \sin\left(\frac{d_C}{R_0}\right) & k=1 \\ (1+z) d_C & k=0 \\ (1+z) R_0 \sinh\left(\frac{d_C}{R_0}\right) & k=-1 \end{cases}$$

and similarly for  $d_D$ , with  $(1+z)^{-1}$  in place of  $(1+z)$

NB: a peculiarity of diameter distance is that it does not necessarily increase with  $z$  or  $d_C$

### COSMOLOGICAL DIMMING

The surface brightness of an object is its luminosity per unit area:

$$I = \frac{L}{D^2}, \text{ if } L \propto \frac{1}{r} \text{ and } D \propto \frac{1}{r} \text{ then } I = \text{const}$$

$$\text{but in cosmology: } I = I_0 \left(\frac{d_D}{d_L}\right)^2 = I_0 (1+z)^{-4}$$

### HUBBLE LAW

Let's Taylor-expand  $\alpha(t)$  to first order:

$$\alpha(t) = \alpha(t_0) + (t-t_0) \dot{\alpha}(t_0) + \mathcal{O}((t-t_0)^2)$$

$$\text{Calling } H_0 = \frac{\dot{\alpha}(t_0)}{\alpha(t_0)} = \dot{\alpha}(t_0) \quad (\alpha(t_0) = 1)$$

$$\frac{1}{1+z} = \alpha(t) \approx 1 + H_0(t-t_0) + \mathcal{O}((t-t_0)^2)$$

$$1+z \approx 1 - H_0(t-t_0) \Rightarrow z \approx H_0(t_0 - t)$$

Photons travel along null geodesics:

$$ds^2 = -c^2 dt^2 + \alpha^2(t) \frac{R^2 dr^2}{1-k^2 r^2} = 0$$

$$\int_t^{t_0} \frac{cdt}{\alpha(t)} = \int_0^r \frac{R dr'}{\sqrt{1-k^2 r'^2}}$$

The right-hand side is expanded as:

$$\int_0^t \frac{R_0 c'}{\sqrt{1-k r'^2}} = R_0 c + \mathcal{O}(1/R_0 c t^3)$$

while the left-hand side is:

$$\int_t^{t_0} \frac{c dt}{\Omega(t)} = c(t_0 - t) + \mathcal{O}((t_0 - t)^2) \approx \frac{cz}{H_0}$$

$$\begin{aligned} \text{To the same first order: } d_L &\approx d_D \approx d_p \approx d_c \approx R_0 c \\ \Rightarrow c z &= H_0 d_L \end{aligned}$$

This shows that:

- +  $z$  IS NOT a Doppler redshift
- +  $v = cz$  IS NOT a velocity in the usual sense
- +  $z$  is a gravitational redshift due to time dilation in an expanding Universe

## SUPERLUMINAL MOTIONS

Consider the "Hubble Law" in the form:

$$v = H d_p$$

this is how unobservable proper distance varies with time.

It is possible that:  $v = H d_p > c$

This is NOT a violation of SR, that applies to objects passing through our (local) inertial frame, so that their velocity can be measured

## DECELERATION PARAMETER

The relations above can be expanded to second order

$$a(t) \approx 1 + H_0(t-t_0) - \frac{1}{2} q_0 H_0^2 (t-t_0)^2 + \dots$$

$$q_0 = - \frac{\ddot{a}(t_0)}{\dot{a}(t_0)^2} \quad \text{Deceleration parameter}$$

$$z \approx H_0(t_0 - t) + \left(1 + \frac{1}{2} q_0\right) H_0^2 (t-t_0)^2 + \dots$$

$$\int_t^{t_0} \frac{c dt}{a(t)} \approx C \left[ (t_0 - t) + \frac{1}{2} H_0 (t_0 - t)^2 + \dots \right]$$

$$\approx \frac{C}{H_0} \left[ z - \frac{1}{2} (1 + q_0) z^2 + \dots \right]$$

$$d_L = (1+z) R_0 z \approx \frac{C}{H_0} \left[ z + \frac{1}{2} (1 - q_0) z^2 + \dots \right]$$

This last, observable relation makes it possible to measure  $q_0$  from the non linearity of Hubble law at  $z \approx 1$

We will see that this parameter is related to the matter/energy content of the universe, that is responsible for its deceleration - or acceleration.